Anomalous radio signals: ultra-high-energy neutrinos or new physics?

Mauricio Bustamante

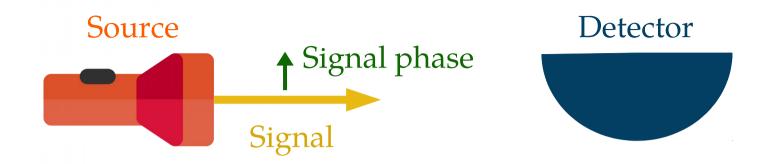
Niels Bohr Institute, University of Copenhagen

UNIVERSITY OF COPENHAGEN

NBIA N-Talk Copenhagen, November 23, 2018

VILLUM FONDEN





Reflecting medium

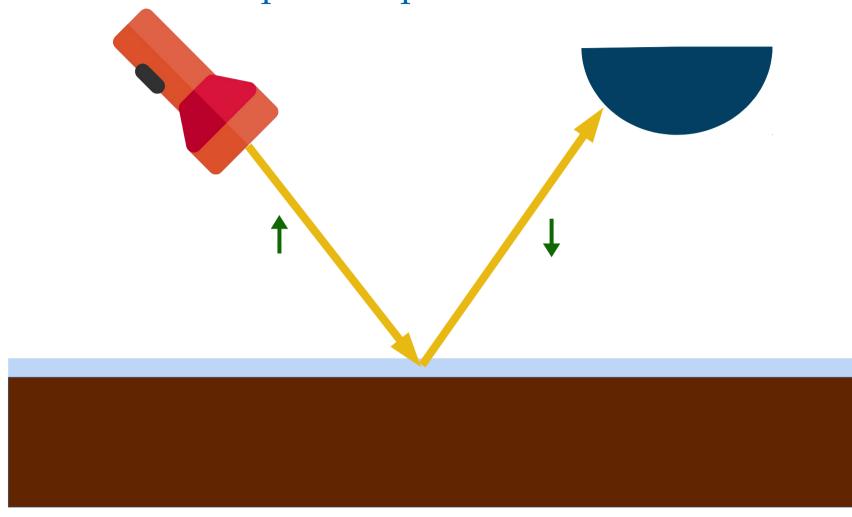
Thick layer of opaque substrate

Case 1: No reflection → no phase flip





Case 2: Reflection → phase flip



But what if we saw this?

► Phase is not flipped, so there was no reflection

► *But* the signal comes from below!

Did a signal go through the substrate?

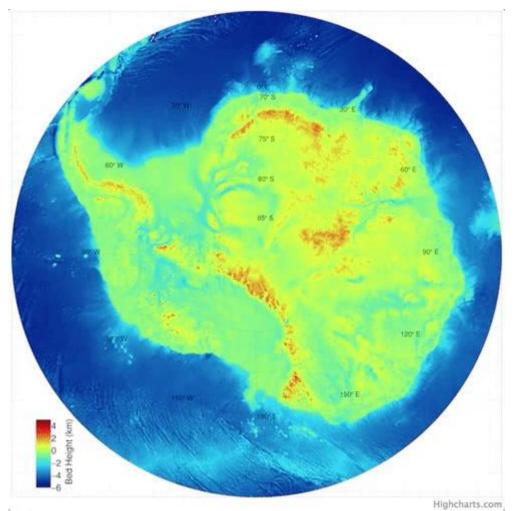
ANITA: Antarctic Impulsive Transient Antenna



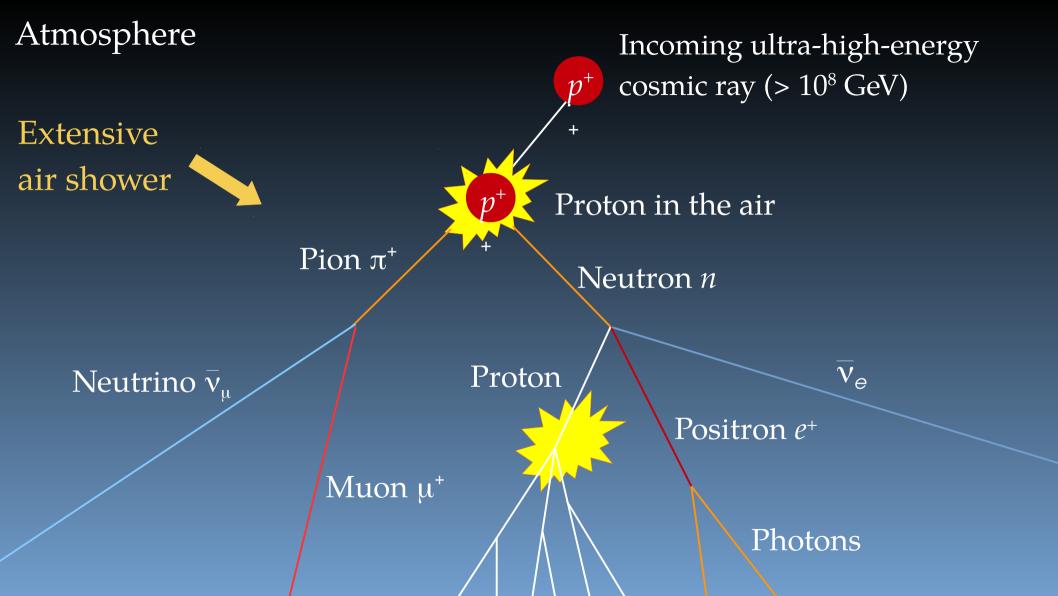


Photo by Spencer Kleir

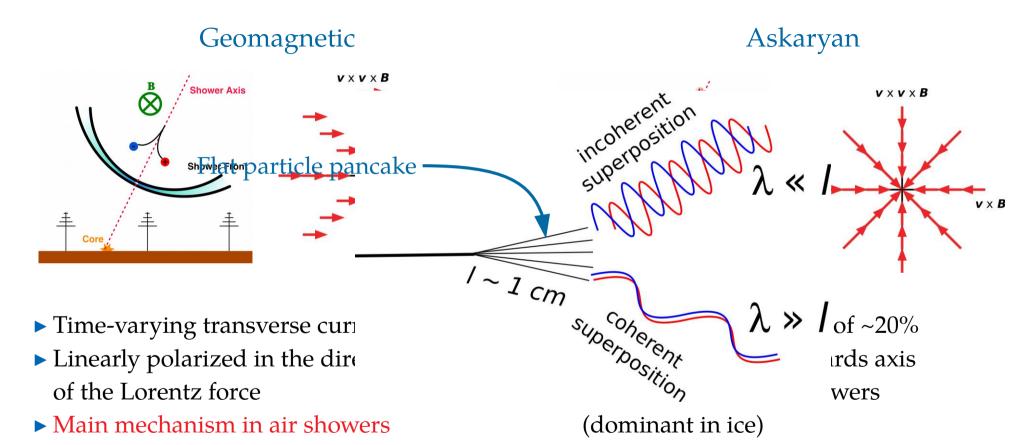
Photo by Brian Hill/U. Hawaii-Manoa

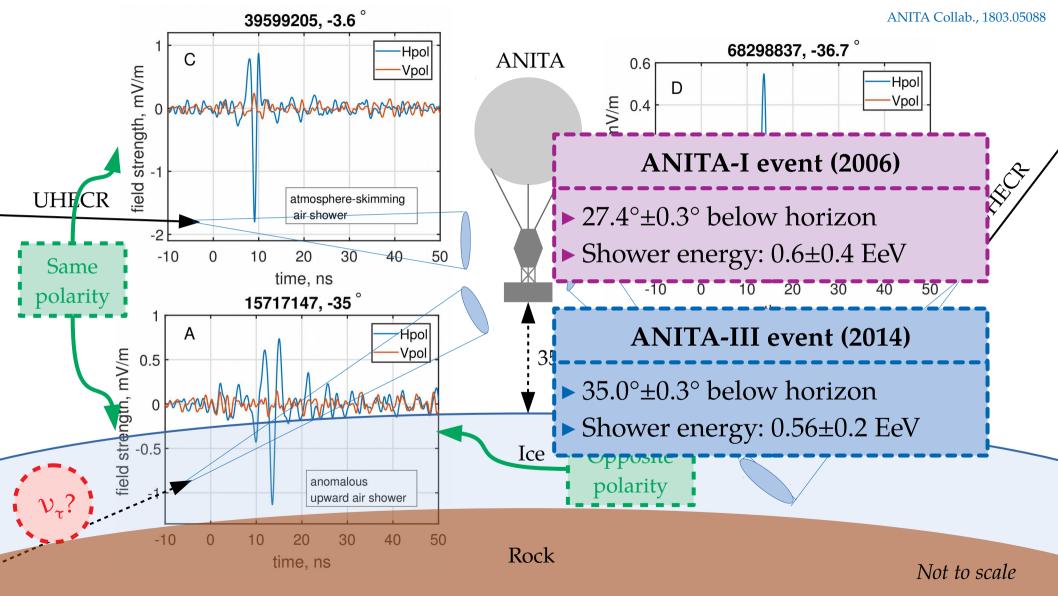


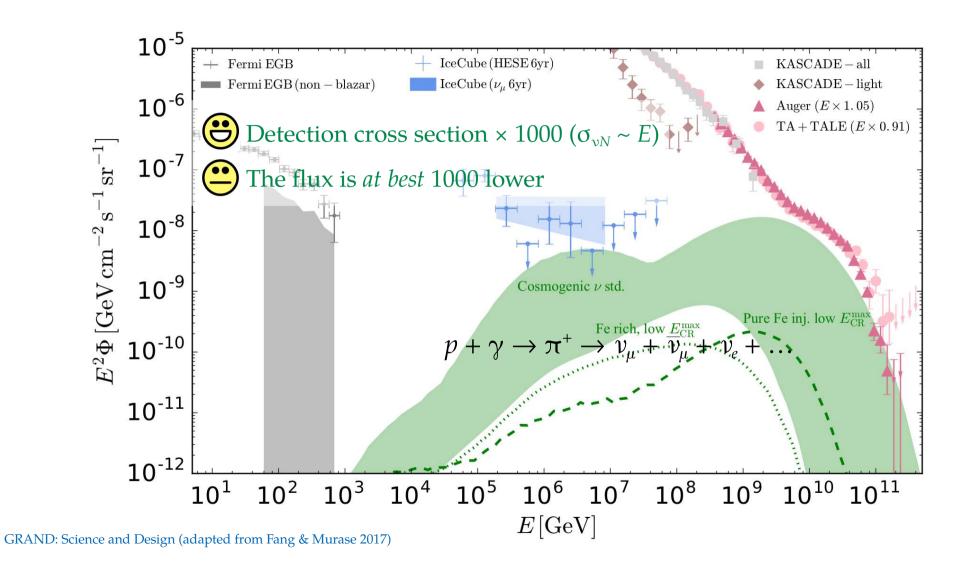
ANITA-IV flight path (Credit: UCL)

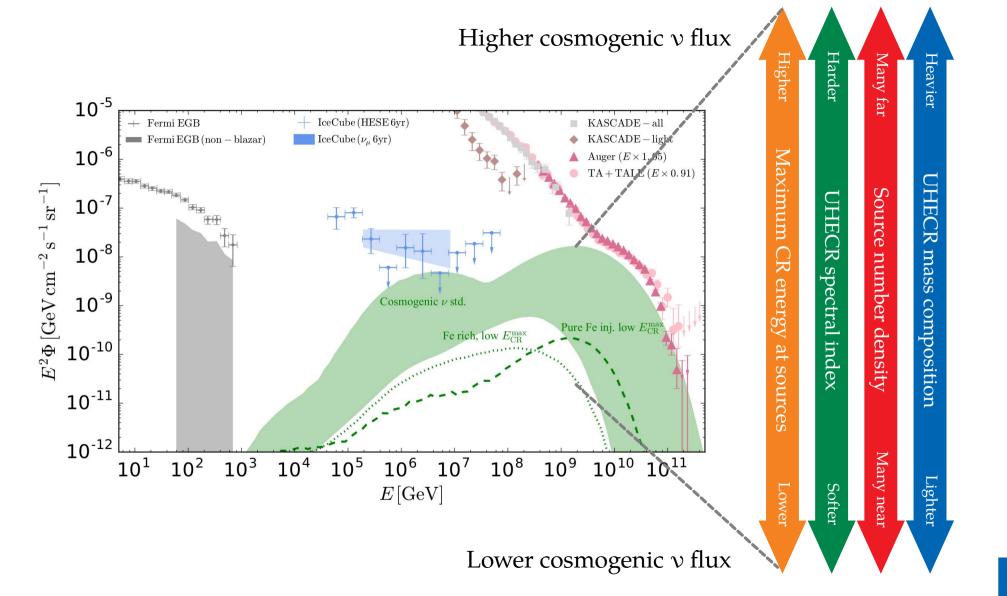


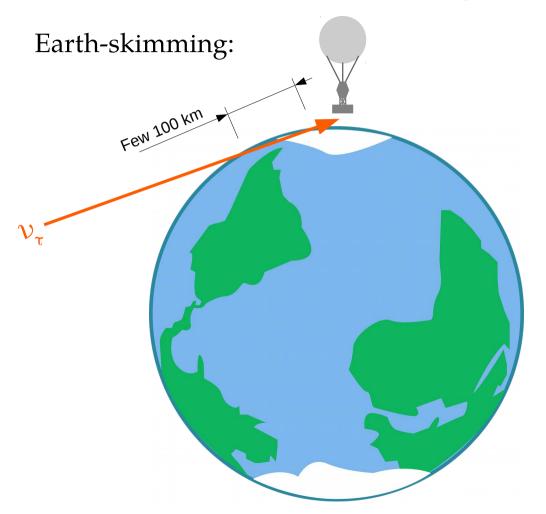
Radio emission from air showers



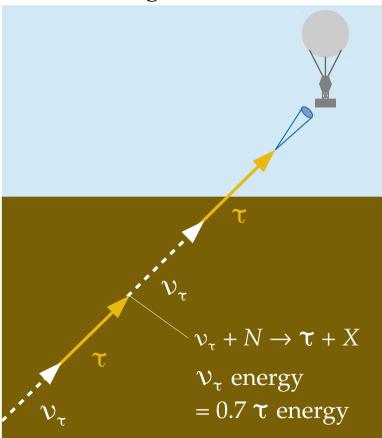


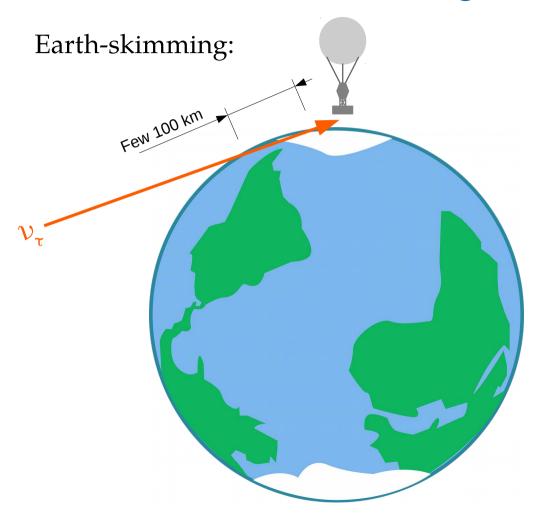




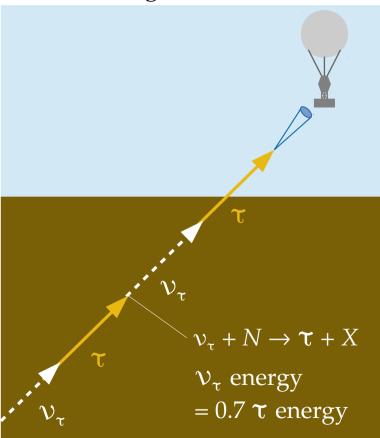


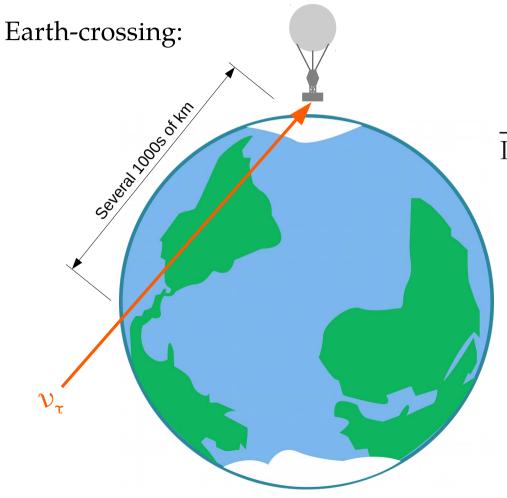
Neutrino regeneration:





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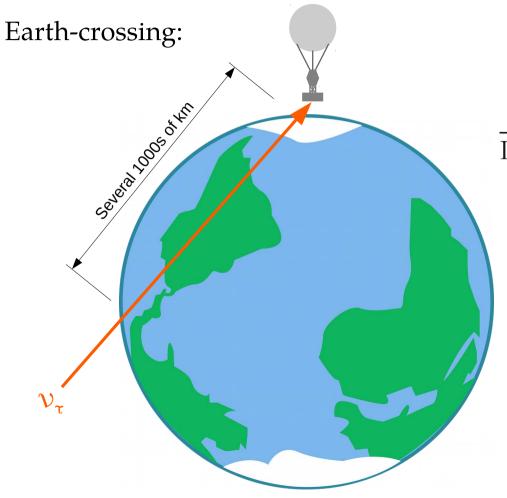




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$$\frac{\text{Chord inside Earth}}{\text{Interaction length in Earth}} = \frac{7000 \text{ km}}{390 \text{ km}} = 18$$

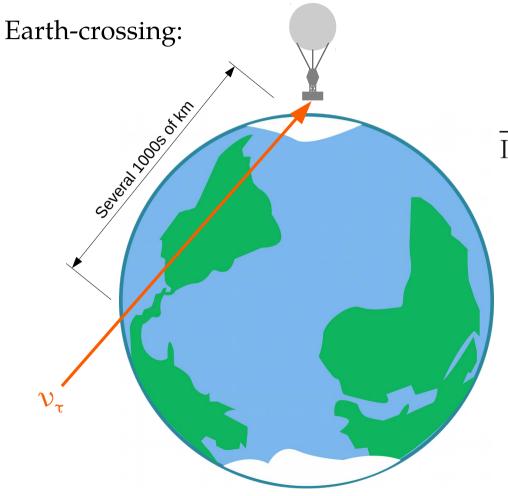
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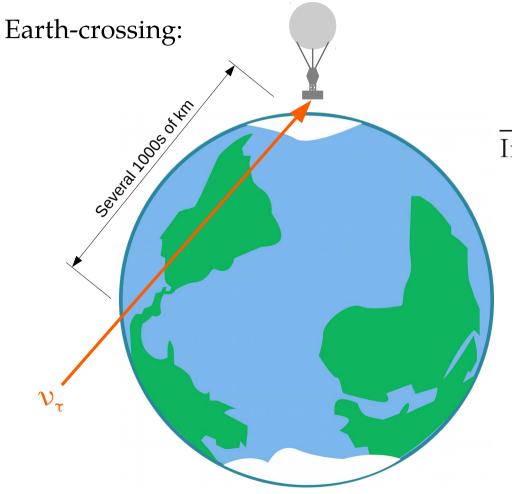
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Solution: Make EeV cross section 10 times smaller

Problem: Within the SM, only factor-of-few suppression is possible



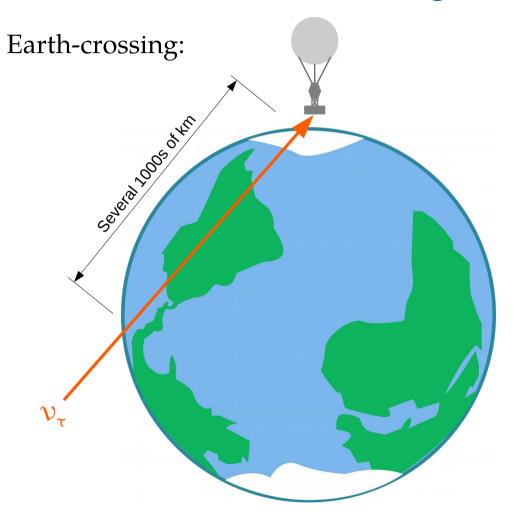
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Problems with diffuse-flux interp.

- ► Flux needs to be 10⁸ times larger
- No events seen closer to horizon



Transient astrophysical event?

- ► ANITA-1 event: none associated
- ► ANITA-3 event:
 - ▶ Type-Ia SN2014dz (z = 0.017)
 - ▶ Within 1.9°, 5 hours before event
 - ▶ Probability of chance SN: 3×10^{-3}
 - \triangleright v luminosity must exceed bolometric luminosity of 4 \times 10⁴² erg s⁻¹

By diffuse flux limits

- ▶ Propagate EeV v_{τ} in Earth
- ► Same directions as AAEs*
- $\triangleright v_{\tau}$ interact and regenerate
- ▶ Probability of τ exiting: 10^{-8} – 10^{-7}
- ▶ Diffuse flux needed is ~10⁶ higher than Auger/IceCube upper limits
- p-value for a Standard-Model v: $p = 2.8 \times 10^{-12} (7\sigma)$

D. Fox et al., 1809.09615

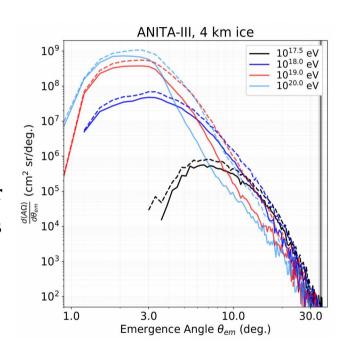
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By direction arguments



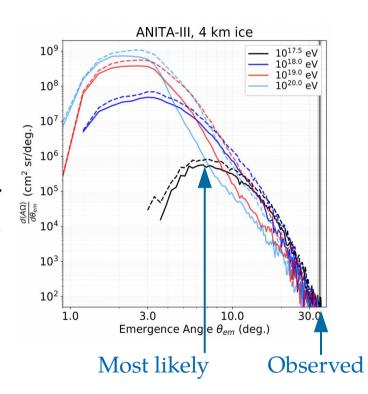
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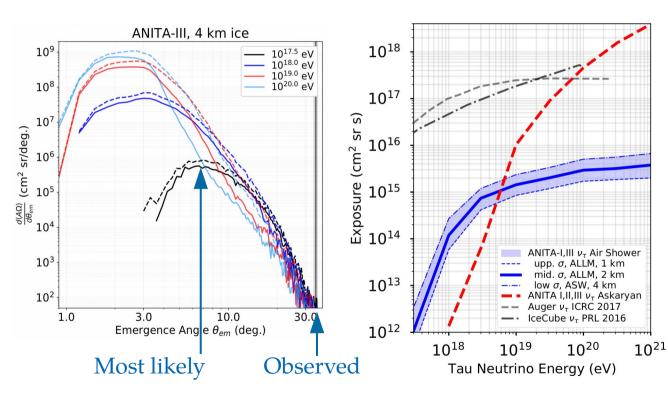
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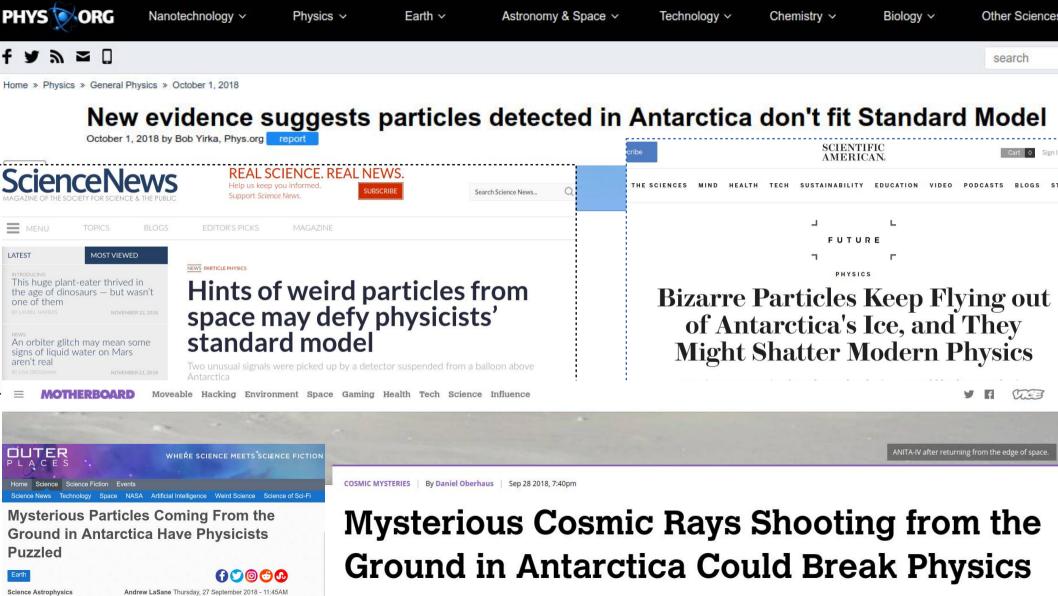
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By comparison to others

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So, is ANITA seeing the first ultra-high-energy neutrinos?

Most likely not



So what is ANITA seeing?

- ► Transition radiation [Motloch et al., PRD 2017]:
 - \blacktriangleright Wide-angle emission of radio waves at ice-air interface could make horizontal v_{τ} look upgoing
 - ▶ Assessment: Needs too large a diffuse flux of v_{τ} , because transition radiation is a small effect
- ► Sterile neutrinos [Cherry & Shoemaker, 1802.01611; Huang, 1804.05362]:
 - ▶ Sterile neutrinos propagate in Earth, then convert $v_s \rightarrow v_r$
 - ▶ Assessment: Model predicts more (unseen) events at shallower angles
- ▶ Dark matter decay in Earth core [Anchordoqui et al., 1803.11554]:
 - ▶ Decay of 480-PeV sterile right-handed ν_r (relic DM) trapped "puffy" Earth core: $\nu_r \rightarrow \text{Higgs} + \nu_\tau$
 - ► Assessment: Viable, but exotic explanation
- ► Staus (NLSP) [Fox *et al.*, 1809.09615]:
 - ▶ Made by UHECRs, decays to τ prior to exiting; can satisfy: $\sigma \sim \sigma_{vN}/1000$, $\tau \sim 10$ ns (m/500 GeV)
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Decay length ~ Earth radius

What are you taking home?

- ► ANITA has seen (so far) 2 upgoing anomalous events
- ► A Standard-Model explanation is incompatible at 7σ
- ► Transient sources unlikely
- ► There are a few viable BSM models that could explain the events
- ► We need more data! ANITA-IV analysis is underway

Backup slides

UHECR sources are hard to find:

- ► Magnetic deflection: few × 10°, composition-dependent
- ► GZK horizon: $p\gamma_{\text{CMB}}$ scattering kills CRs with > 40 EeV after 100 Mpc

How to find them?

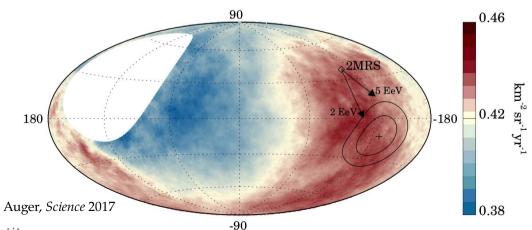
- ► Look for UHECR (> few EeV) arrival-direction anisotropies
- ▶ Look for UHE (> 10 EeV) gamma rays and neutrinos from $p\gamma$

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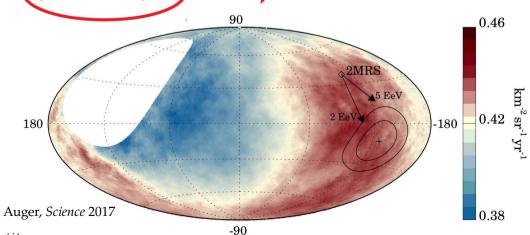
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How to find them?

Can only find nearby sources because of gamma-ray opacity of Universe

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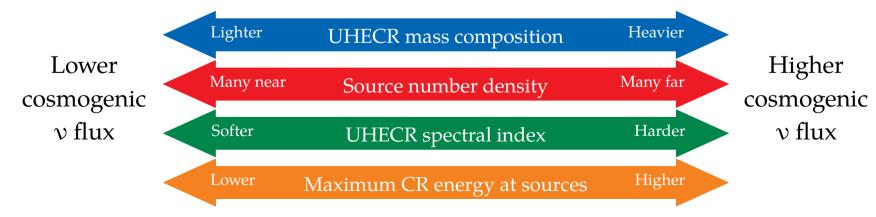
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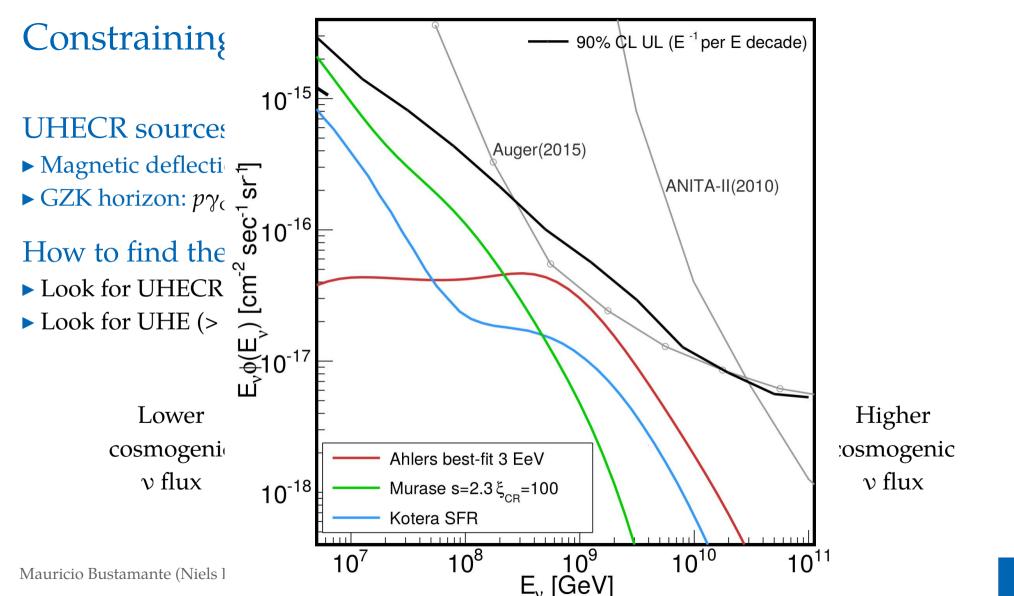
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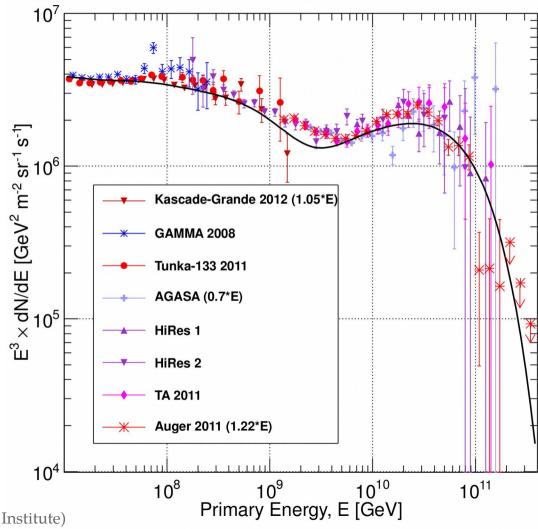
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50+ years of UHECR observations –



Abundant, but hardly interacting ▼ Flux (cm⁻² s⁻¹ sr⁻¹ MeV⁻¹)
10²⁰
10¹⁶
10⁸
10⁴ Abundant and detected Cosmological v Solar v Supernova burst (1987A) Reactor anti-v Background from old supernovae 10^{-4} Terrestrial anti-v 10^{-8} Atmospheric v 10^{-12} 10^{-16} v from AGN ■ Rare but detected 10^{-20} Cosmogenic 10^{-24}

 10^{9}

GeV

 10^{6}

MeV

 10^{3}

keV

eV

 10^{12}

TeV

1015

PeV

Neutrino energy

1018

EeV

 10^{-28}

 10^{-6}

μeV

 10^{-3}

meV

Rare but detected
Very rare,
not detected yet