High-energy neutrinos, cosmic rays, and gamma rays from GRBs

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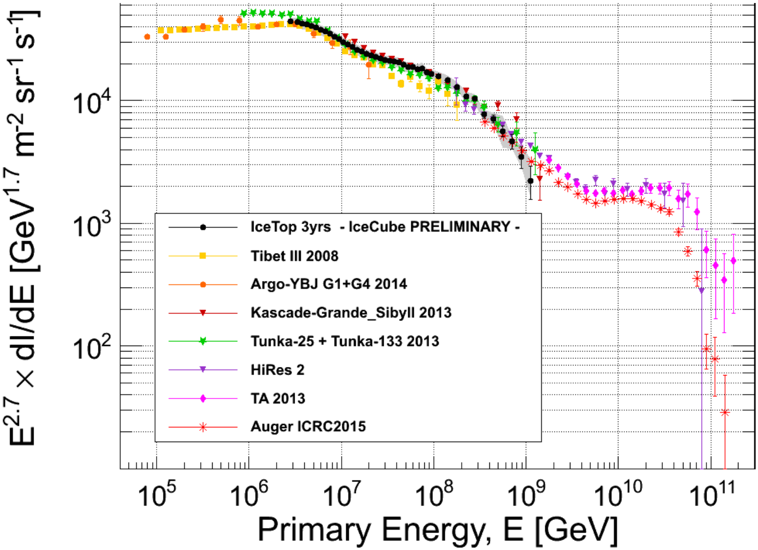
May 25, 2017

THE OHIO STATE UNIVERSITY

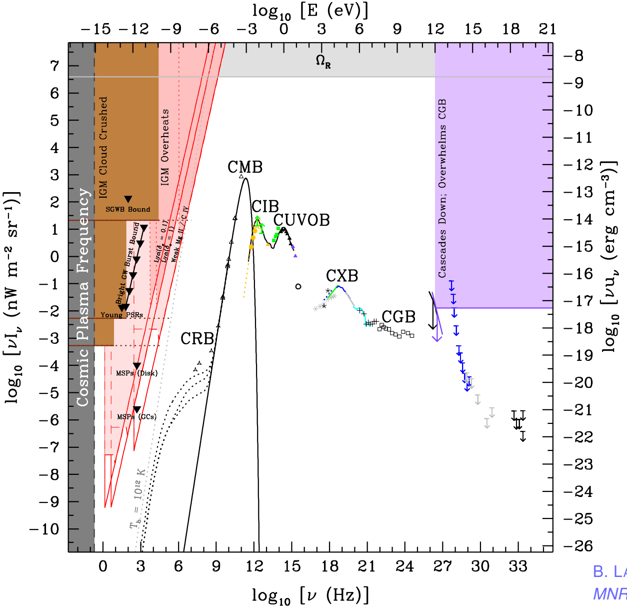
CENTER FOR COSMOLOGY AND ASTROPARTICLE PHYSICS

Ultra-high-energy cosmic rays

After 50+ years of UHECR measurements —



The electromagnetic sky

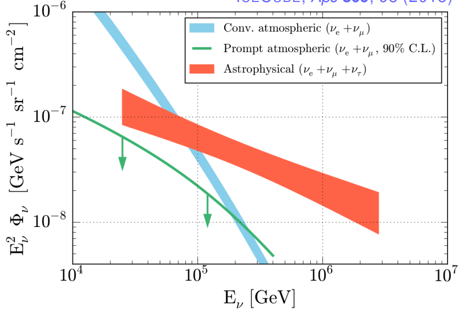


B. Lacki, *MNRAS* **406**, 863 (2010)

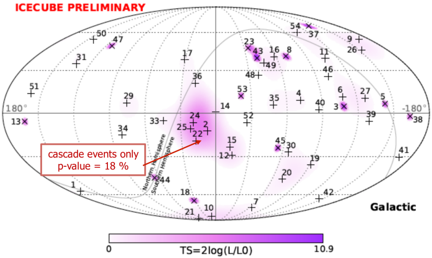
A discovery and a mystery

IceCube has discovered a diffuse astrophysical neutrino flux —

ICECUBE, *ApJ* **809**, 98 (2015)



Energies: 25 TeV – 2 PeV



O. BOTNER, IPA 2015

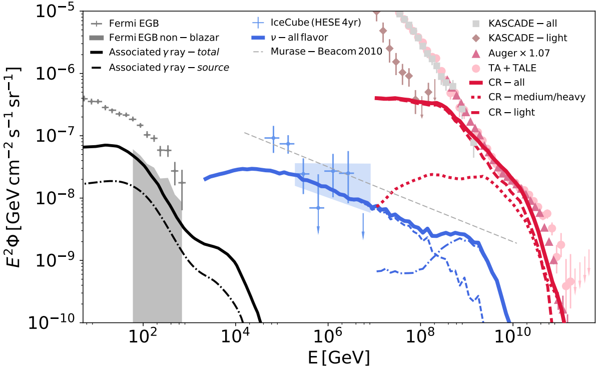
Isotropy of arrivals hints at extragalactic sources

No sources identified yet

Two profound observations from the discovery

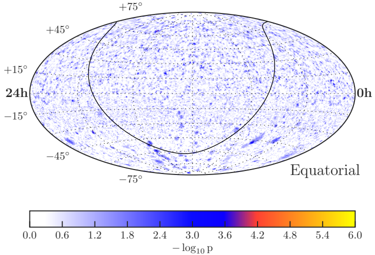
1. Diffuse fluxes of gamma rays, UHECRs, neutrinos are comparable

FANG & MURASE, 1704.00015 ▶

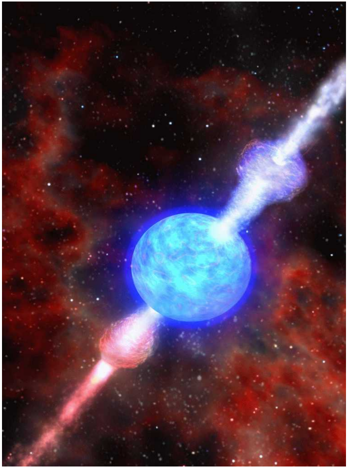


2. There is no (evident) dominant neutrino-bright source or many bright sources

ICECUBE, *ApJ* 2017 [1609.04981]



Why are GRBs natural ν source candidates?



1. They are bright
   - $10^{52} - 10^{53}$ erg in gamma rays
   - Gamma rays up to $\sim$ 100 GeV

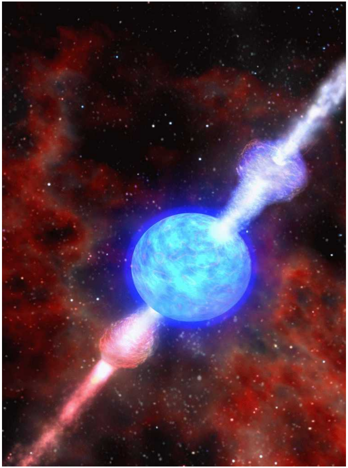
   Implies possible acceleration of protons to high energies

2. They have fast time-variability
   - Structure at 0.01 s scale
   - Compact objects

   Implies high proton and photon number densities

Original expectation:
GRBs produce copious high-energy neutrinos via $p\gamma$ interactions

Why are GRBs natural ν source candidates?



1. They are bright
   - $10^{52} - 10^{53}$
   - Gamma ray

10^{20} erg	H bomb
10^{26} erg	killer asteroid
10^{40} erg	Death Star
10^{33} erg s^{-1}	Sun
10^{41} erg s^{-1}	supernova
10^{45} erg s^{-1}	galaxy

Implies possible acceleration of protons to high energies

2. They have fast time-variability
   - Structure at 0.01 s scale
   - Compact objects

Implies high proton and photon number densities

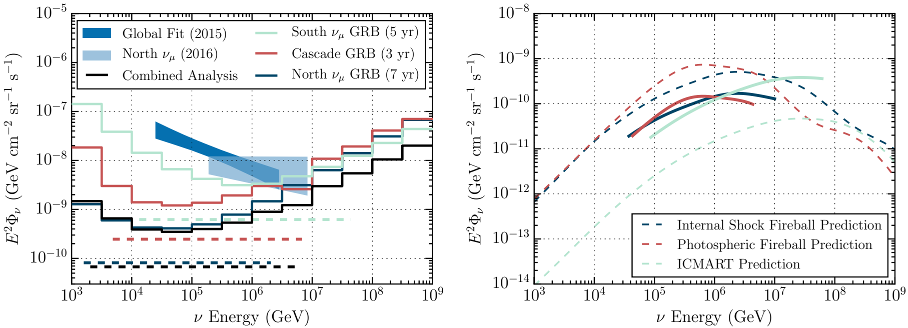
Original expectation:
GRBs produce copious high-energy neutrinos via $p\gamma$ interactions

Our expectations, unrewarded?

IceCube has not found neutrinos associated to GRBs

- 3 yr of showers (all flavors) + 7 yr of upgoing tracks + 5 of downgoing tracks
- 1172 GRBs
- No statistically significant coincidence

Only $\lesssim$ 0.4% of the diffuse flux can be from prompt GRB emission



IceCube, 1702.06868

The  in the room

Why is it still interesting to look for GRB neutrinos?

1. Best candidates for joint high-energy e.m.–neutrino emission
2. Potential sources of ultra-high-energy cosmic rays

Also . . .

3. "Choked" bursts might contribute sizeably to the diffuse flux

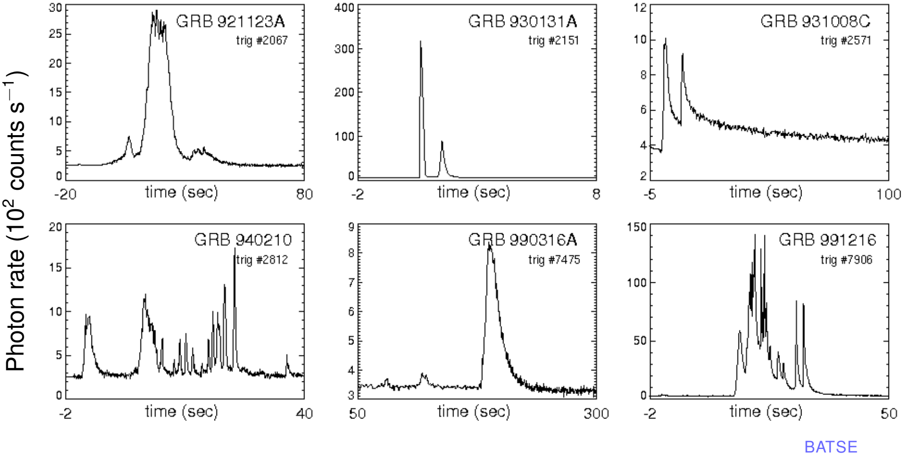
   e.g., [P. Mészaros, E. Waxman 2001] [N. Senno, K. Murase, P. Mészaros 2016] [I. Tamborra, S. Ando 2016]

4. Neutrinos from GRB afterglows expected at $\sim$ EeV

   e.g., [K. Murase 2007] [S. Razzaque, L. Yang 2015]

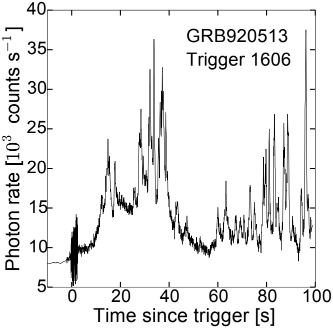
So, every bit of insight into how GRBs make neutrinos helps

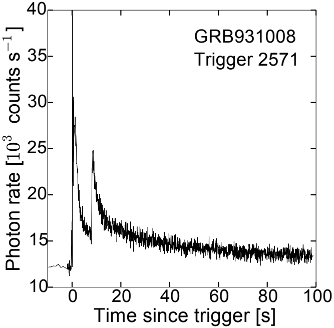
GRBs — a zoo of light curves



Which GRB is brighter in neutrinos?

Just from looking at these gamma-ray light curves,





can we tell which GRB is likely bright in neutrinos?

Which GRB is brighter in neutrinos?

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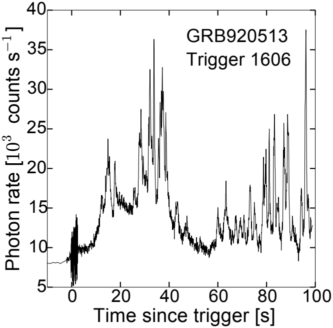
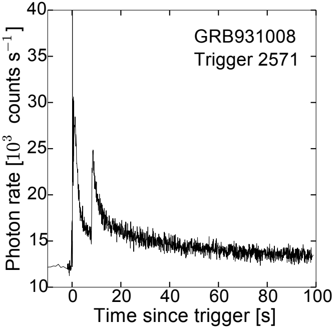


Fast time variability



Slow pulse + fast variability

can we tell which GRB is likely bright in neutrinos?

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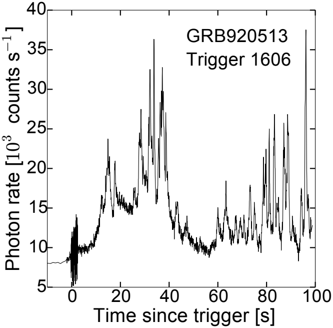

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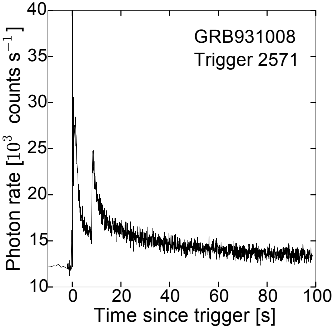
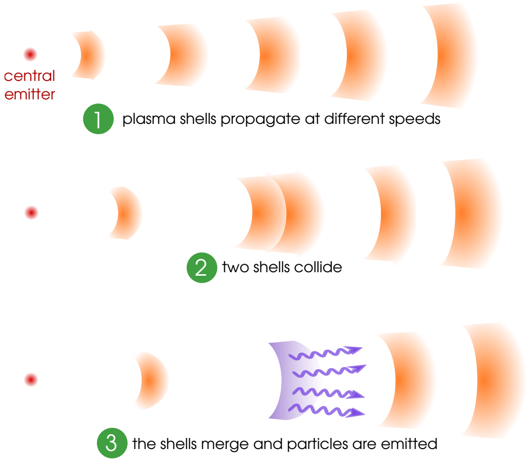


Slow pulse + fast variability

can we tell which GRB is likely bright in neutrinos?

(**Answer:** yes, the one on the left)

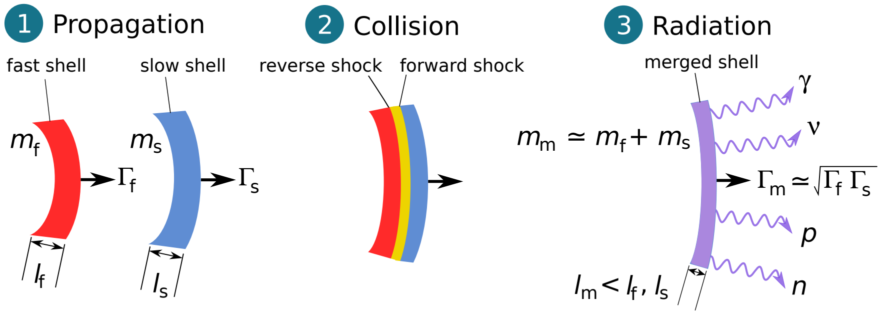
The fireball model — internal collisions



Energy in photons and electrons $\approx$ Energy in magnetic fields

Energy in relativistic protons $\gtrsim 10 \times$ Energy in photons and electrons

Anatomy of an internal collision



Part of the initial kinetic energy radiated as γ's, ν's, p's, and n's:

$$E_{\rm coll}^{\rm iso} = \left(E_{{\rm kin},f}^{\rm iso} - E_{{\rm kin},m}^{\rm iso}\right) + \left(E_{{\rm kin},m}^{\rm iso} - E_{{\rm kin},s}^{\rm iso}\right)$$

1/12	1/12	5/6
$\underbrace{\epsilon_e E_{\rm coll}^{\rm iso}}$	$\underbrace{\epsilon_B E_{\rm coll}^{\rm iso}}$	$\underbrace{\epsilon_p E_{\rm coll}^{\rm iso}}$
energy in photons	energy in magnetic fields	energy in baryons

Producing ν's, CRs, γ rays

power law $\sim E^{-\alpha_p}$ broken power law

$$p\,\gamma \rightarrow \Delta^+(1232) \rightarrow \begin{cases} n\pi^+\,, & \text{BR} = 1/3 \\ p\pi^0\,, & \text{BR} = 2/3 \end{cases}$$

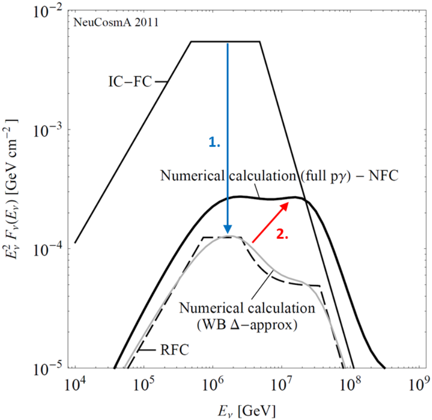
$$\pi^+ \rightarrow \mu^+\nu_\mu \rightarrow \bar{\nu}_\mu e^+ \nu_e \nu_\mu$$

$$\pi^0 \rightarrow \gamma\gamma$$

$$n \text{ (escapes)} \rightarrow p e^- \bar{\nu}_e$$

We use NeuCosmA:

- More production channels
- Detailed particle interactions
- Flavor mixing



energy in neutrinos $\propto$ energy in gamma rays

S. Hümmer, P. Baerwald, W. Winter, *PRL* **108**, 231101 (2012)

What are the ingredients?

To calculate the ν flux from a GRB, we need:

- Its gamma-ray luminosity L_γ^{iso} [erg s^{-1}] [measured]
- Its variability timescale t_v [s], from the light curve [measured]
- Break energy of its photon spectrum $\epsilon_{\gamma,\text{break}}$ [MeV] [measured]
- Its redshift z [(sometimes) measured]
- The bulk Lorentz factor of its jet Γ [estimated]
- The energy in electrons, magnetic field, protons [estimated]

Now let us cook up the neutrinos ▶

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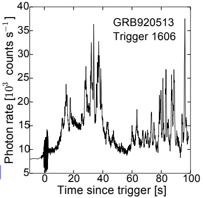

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- The energy in electrons, magnetic field, protons [estimated]

Now let us cook up the neutrinos ▶

Normalizing neutrinos with observed gamma rays

For each GRB,

energy in neutrinos $\propto$ energy in gamma rays

$$\int_0^\infty dE_\nu E_\nu F_\nu(E_\nu) = \frac{1}{8} \underbrace{\left[1 - \left(1 - \langle x_{p\to\pi}\rangle\right)^{\Delta R/\lambda_{p\gamma}}\right]}_{f_\pi:\ \text{fraction of total } p \text{ energy given to } \pi\text{'s}} \frac{1}{f_e} \int_{1\ \text{keV}}^{10\ \text{MeV}} d\epsilon_\gamma \epsilon_\gamma F_\gamma(\epsilon_\gamma)$$

ΔR: size of the emitting region
$\lambda_{p\gamma}$: mean free path for $p\gamma$ interactions
$\langle x_{p\to\pi}\rangle$: avg. fraction of p energy transferred to a π in one interaction
f_e^{-1}: ratio of energy in protons to energy in photons ("baryonic loading")

$$\text{Optical depth to } p\gamma: \frac{\Delta R}{\lambda_{p\gamma}} = \left(\frac{L_\gamma^{\text{iso}}}{10^{52}\ \text{erg s}^{-1}}\right)\left(\frac{0.01}{t_v}\right)\left(\frac{10^{2.5}}{\Gamma}\right)^4\left(\frac{\text{MeV}}{\varepsilon_{\gamma,\text{break}}}\right)$$

Cooking up the neutrinos (simple example)

Observed gamma-ray fluence [GeV^{-1} cm^{-2}]

$$F_\gamma(\varepsilon_\gamma) \propto \begin{cases} (\varepsilon_\gamma/\varepsilon_{\gamma,\mathrm{br}})^{-1} & , \varepsilon_\gamma < \varepsilon_{\gamma,\mathrm{br}} = 1\ \mathrm{MeV} \\ (\varepsilon_\gamma/\varepsilon_{\gamma,\mathrm{br}})^{-2.2} & , \varepsilon_\gamma \geqslant \varepsilon_{\gamma,\mathrm{br}} \end{cases}$$

+

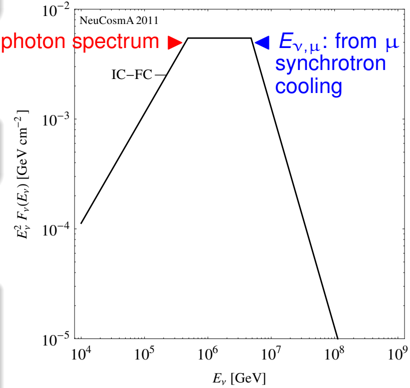
Assumed proton spectrum in the source

$$N'_p(E_p) \propto E'^{-2}_p$$

=

Neutrinos from $p\gamma$, via Δ resonance

$$F_\nu(E_\nu) \propto \begin{cases} \left(\frac{E_\nu}{E_{\nu,\mathrm{br}}}\right)^{-\alpha_\nu} & , E_\nu < E_{\nu,\mathrm{br}} \\ \left(\frac{E_\nu}{E_{\nu,\mathrm{br}}}\right)^{-\beta_\nu} & , E_{\nu,\mathrm{br}} \leqslant E_\nu < E_{\nu,\mu} \\ \left(\frac{E_\nu}{E_{\nu,\mathrm{br}}}\right)^{-\beta_\nu}\left(\frac{E_\nu}{E_{\nu,\mu}}\right)^{-2} & , E_\nu \geqslant E_{\nu,\mu} \end{cases}$$

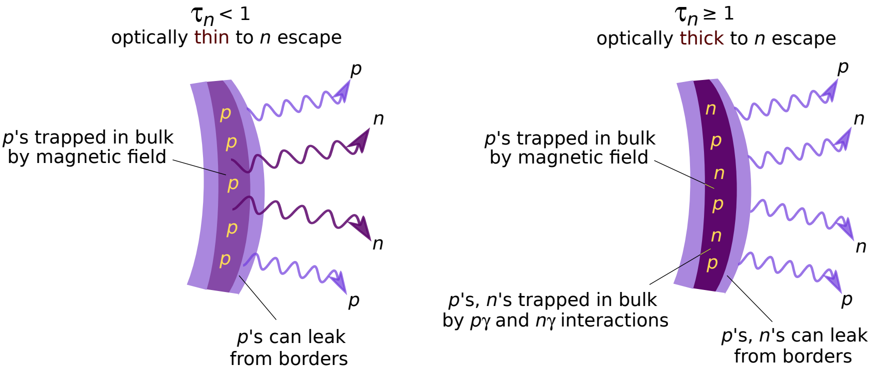


E. Waxman, J. N. Bahcall, *PRL* **78**, 2292 (1997)
D. Guetta *et al.*, *Astropart. Phys.* **20**, 429 (2004)

Two forms of UHECR escape

- Neutrons, produced together with ν's; or
- Protons that leak out without producing ν's in the source

Relative contributions determined by $\tau_n \equiv \left(t_{p\gamma}^{-1} / t_{\rm dyn}^{-1} \right) \Big|_{E'_{p,\max}}$



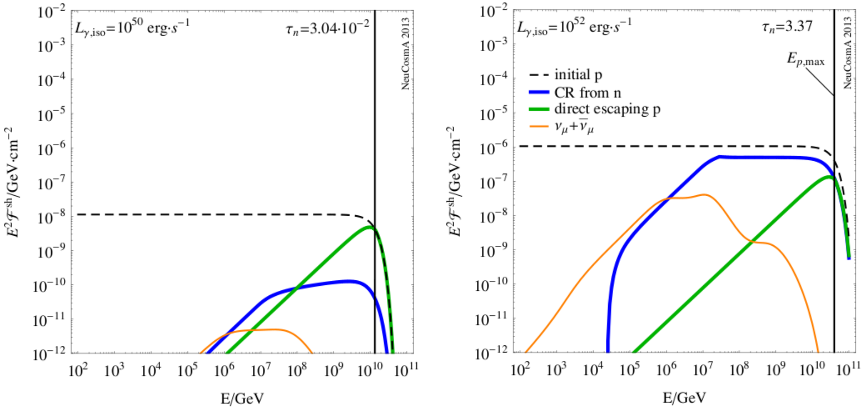
P. Baerwald, MB, W. Winter, *ApJ* **768**, 186 (2013)
P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)
See also: H. He *et al.*, *ApJ* **752**, 29 (2012)

A two-component model of UHECR emission

Sample neutrino fluences —

Optically **thin** source

Optically **thick** source

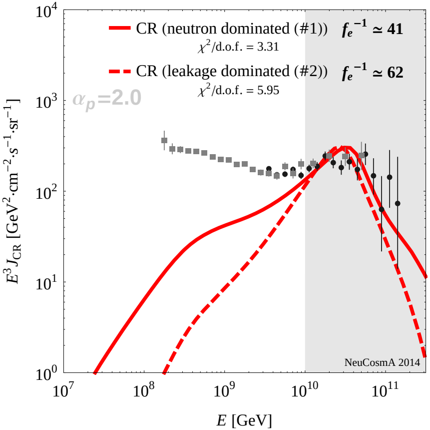


P. Baerwald, MB, W. Winter, *ApJ* **768**, 186 (2013)

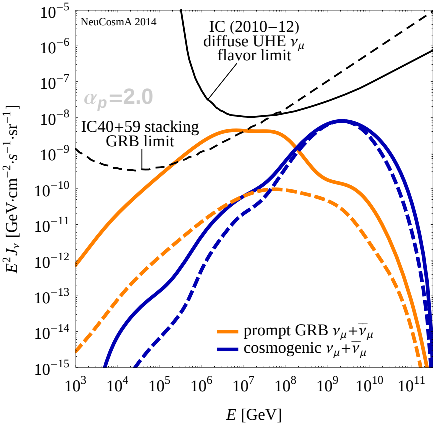
Diffuse fluxes at Earth

Neutron model vs. two-component model:
prompt and cosmogenic ν's

UHECRs



Neutrinos



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Using UHECR + neutrino constraints

Constrain parameter space using UHECRs and ν upper bounds:

P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)

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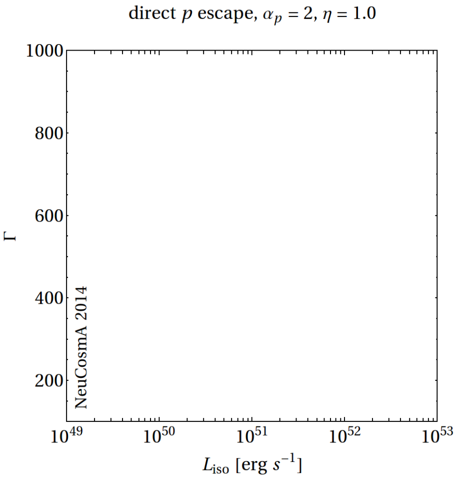

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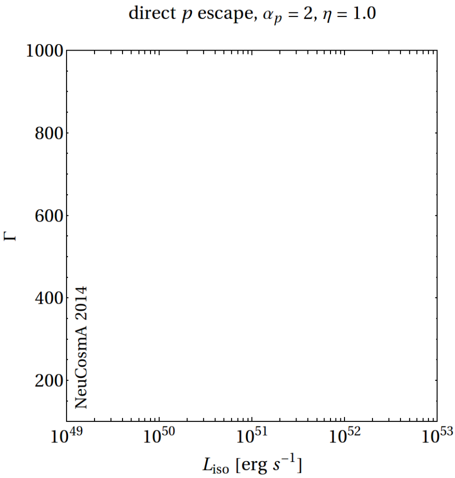

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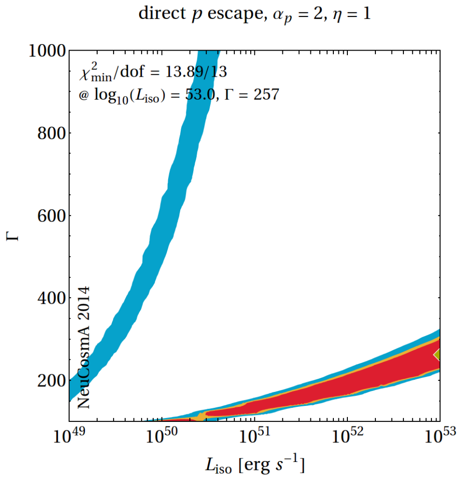

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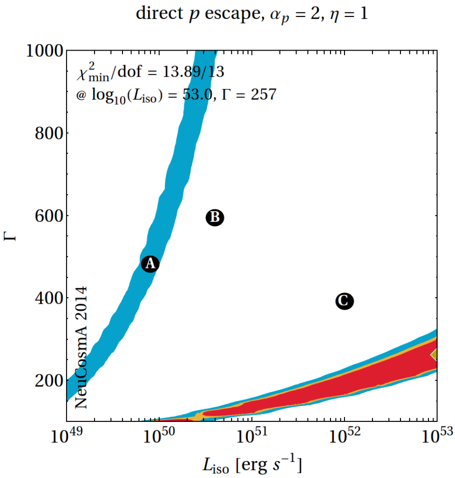

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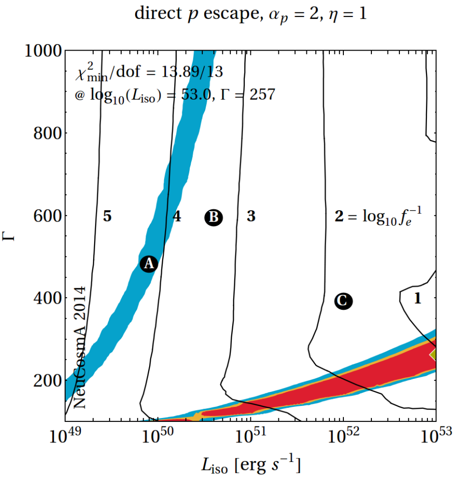

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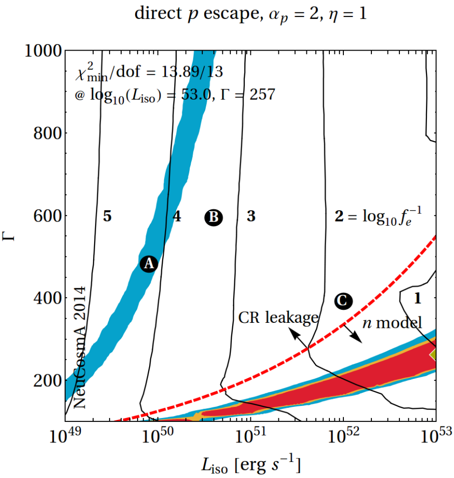
4. Find the baryonic loading f_e^{-1} at each point
5. Identify the region corresponding to pure n escape and to n escape + CR leakage



P. BAERWALD, MB, W. WINTER, *Astropart. Phys.* **62**, 66 (2015)

Using UHECR + neutrino constraints

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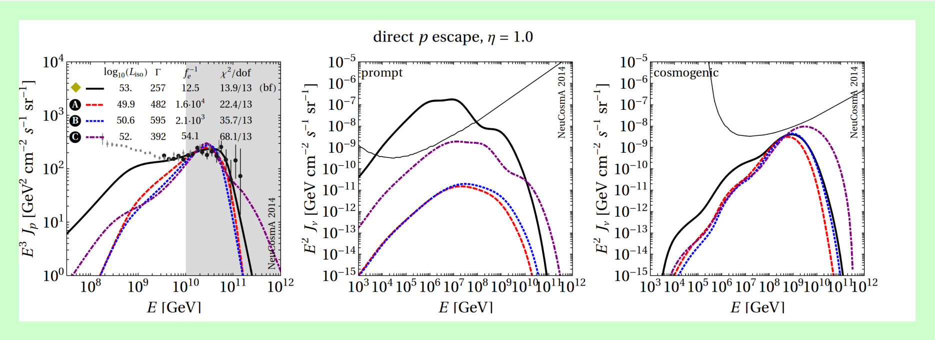
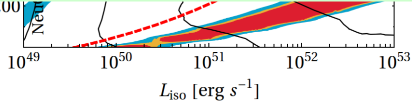
direct p escape, $\alpha_p = 2$, $\eta = 1$





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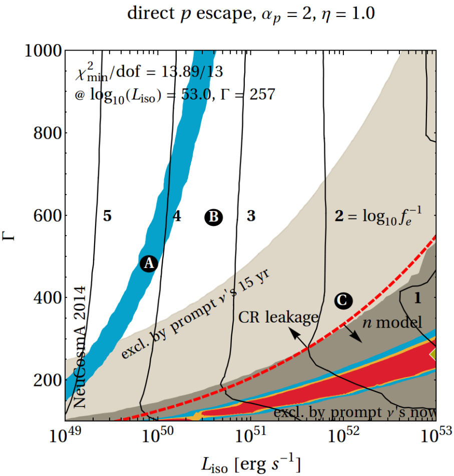
6. Find where the number of prompt ν_μ's is > 2.44, *i.e.*, the excluded at 90% C.L.



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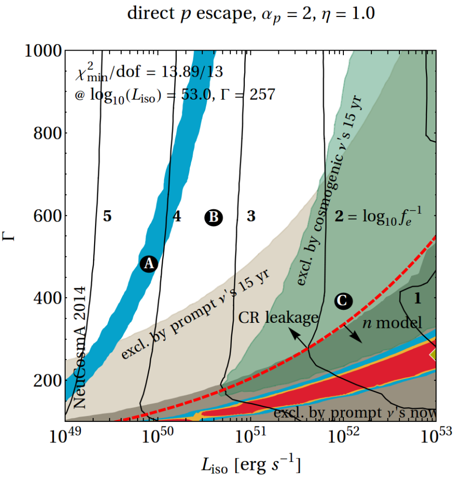
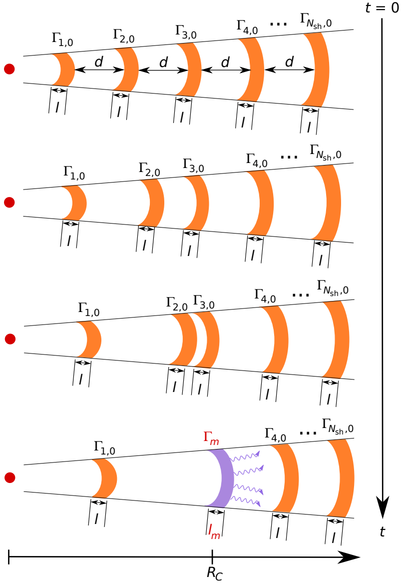
5. Identify the region corresponding to pure *n* escape and to *n* escape + CR leakage
6. Find where the number of prompt ν_μ's is > 2.44, *i.e.*, the excluded at 90% C.L.
7. After 15 yr of exposure and no detection, cosmogenic neutrinos also help



P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)

An evolving fireball

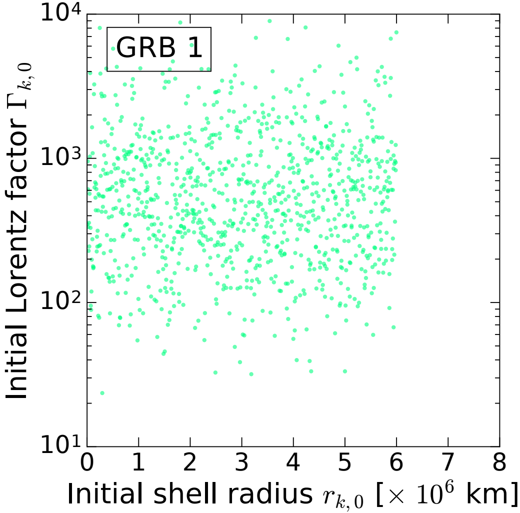
- The fireball expands with time
- Particle densities $\propto R_C^{-2}$
- $\sim$ 1000 shells
- Different speeds, masses
- They collide at different radii



S. Kobayashi, T. Piran, R. Sari, *ApJ* **490**, 92 (1997)
F. Daigne, R. Mochkovitch, *MNRAS* **296**, 275 (1998)

Initial distribution of shell speeds

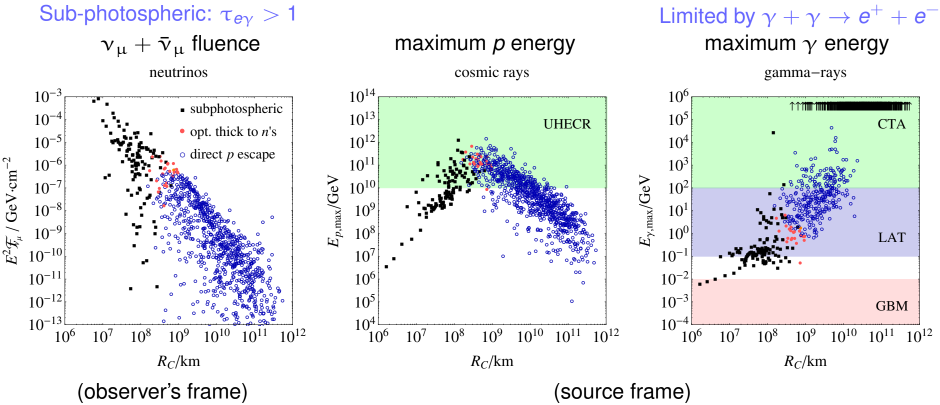
Broad distribution — slow and fast shells find each other early



MB, Baerwald, Murase, Winter, *Nat. Commun.* 2015

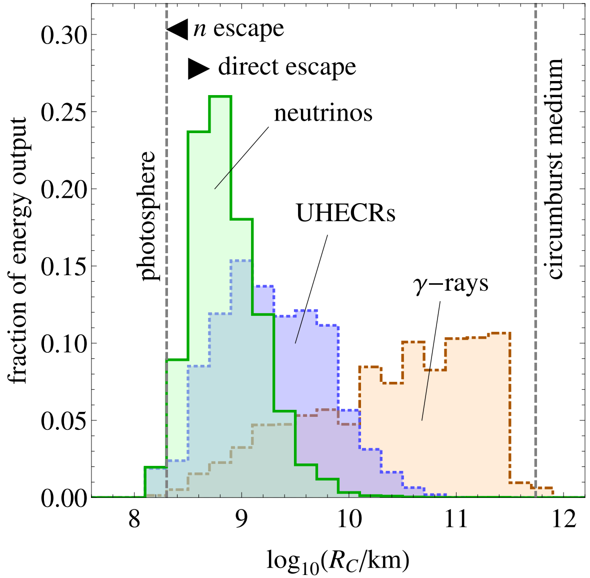
Tracking each collision individually

Each collision occurs in a different emission regime —



MB, Baerwald, Murase, Winter, *Nat. Commun.* 2015

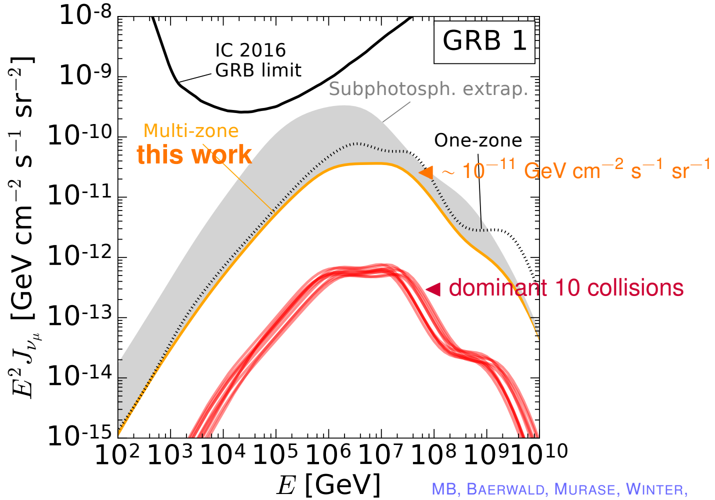
Different particles come from different jet regions



MB, Baerwald, Murase, Winter, *Nat. Commun.* 2015

A robust minimal diffuse ν flux from GRBs

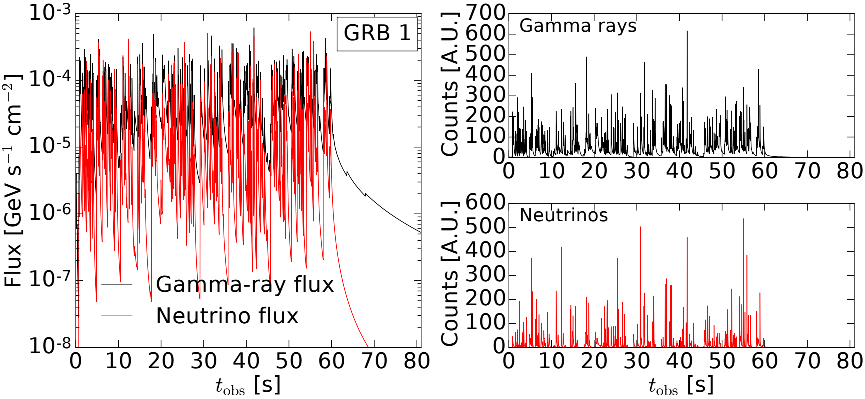
- Take the simulated burst as stereotypical
- Quasi-diffuse neutrino flux, assuming 667 identical GRBs per year:



MB, Baerwald, Murase, Winter, *Nature Commun.* 2015

Synthetic light curves

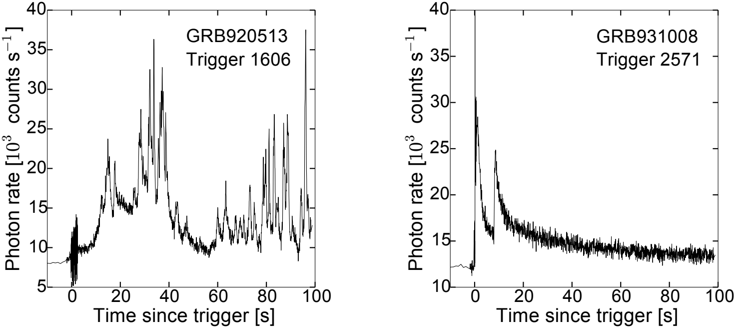
Each collision emits a particle pulse
— their superposition yields a synthetic light curve:



MB, BAERWALD, MURASE, WINTER
Nature Commun. 2015

Which GRB is brighter in neutrinos?

Just from looking at these gamma-ray light curves,



can we tell which GRB is likely bright in neutrinos?

Which GRB is brighter in neutrinos?

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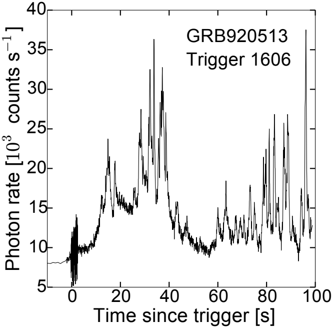
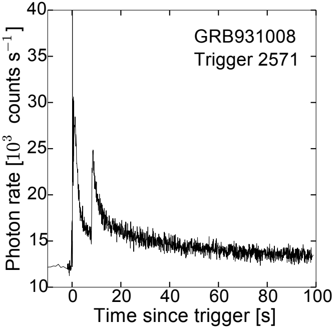


Fast time variability



Slow pulse + fast variability

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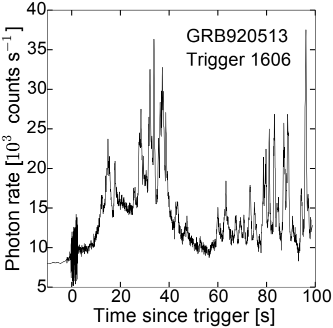

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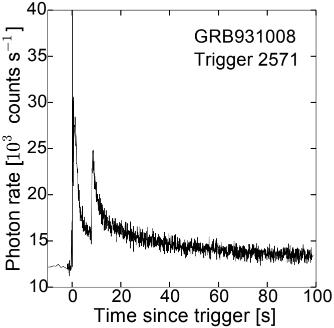


Slow pulse + fast variability

can we tell which GRB is likely bright in neutrinos?

Answer: yes, the one on the left

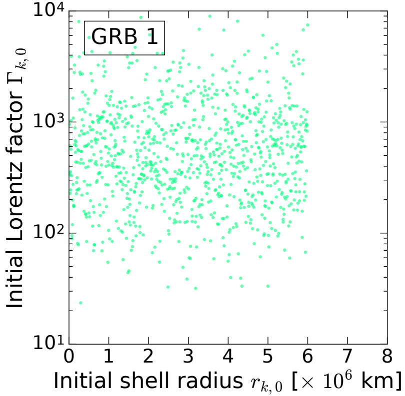
What makes a GRB bright in neutrinos?

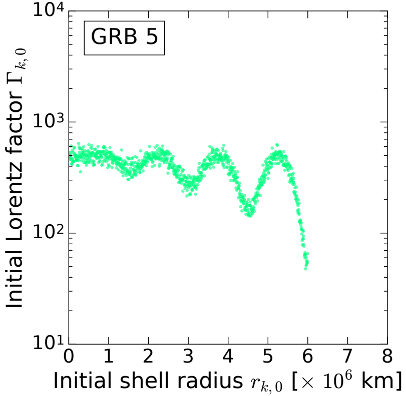
Undisciplined **GRB engine**

- Broad Γ distribution
- *E.g.*, engine emits shells with log-normal Γ distribution

Disciplined **GRB engine**

- Narrow Γ distribution
- *E.g.*, engine emits shells with oscillating Γ





MB, Heinze, Murase, Winter, *ApJ* 2017

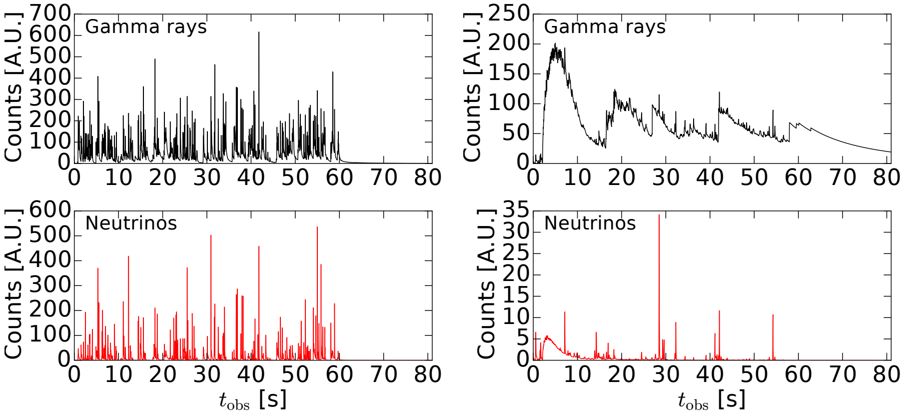
Light curves

Undisciplined GRB engine

- Fast variability dominates
- No broad pulses

Disciplined GRB engine

- Broad pulses dominate
- Fast variability on top



MB, Heinze, Murase, Winter, *ApJ* 2017

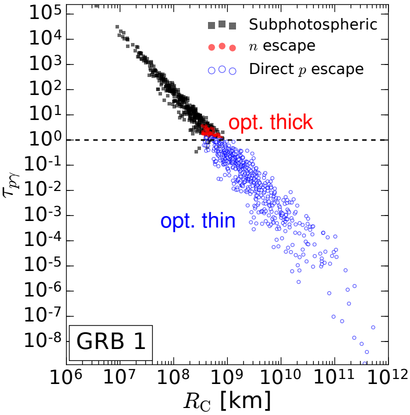
How many optically thick collisions?

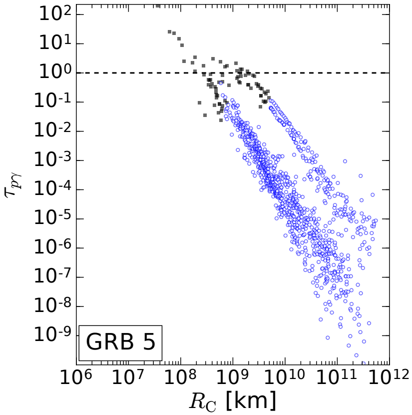
Undisciplined GRB engine

- Shells with very different speeds
- Collide quickly, close to center
- High p and γ densities
- Some optically thick collisions

Disciplined GRB engine

- Shells with similar speeds
- Collide far from center
- Low p and γ densities
- No optically thick collisions

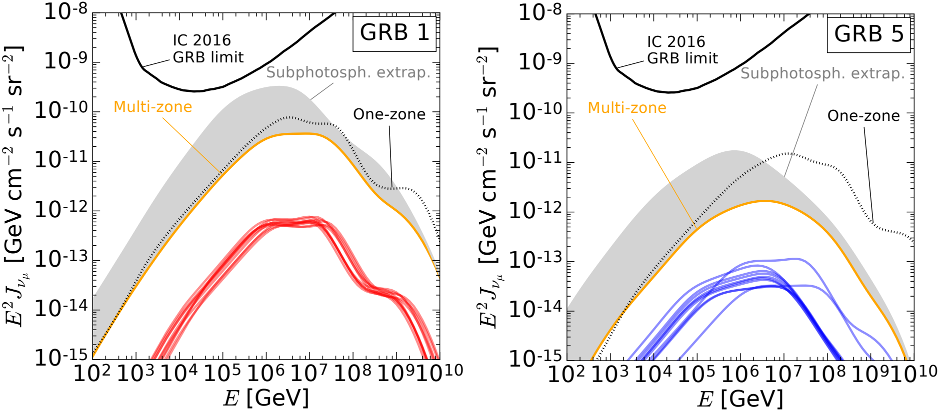




MB, Heinze, Murase, Winter, *ApJ* 2017

So which burst is neutrino-bright?

Undisciplined GRB engine	Disciplined GRB engine
$\sim 10^{-11}$ GeV cm^{-1} s^{-1} sr^{-1}	$\sim 10^{-12}$ GeV cm^{-1} s^{-1} sr^{-1}



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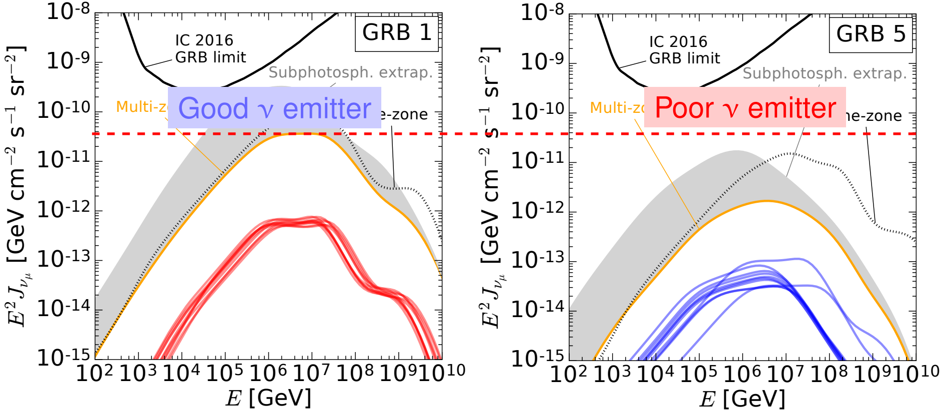

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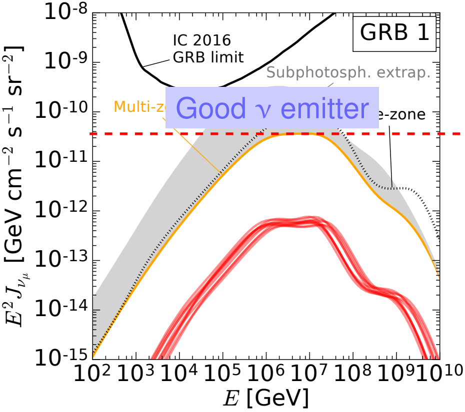
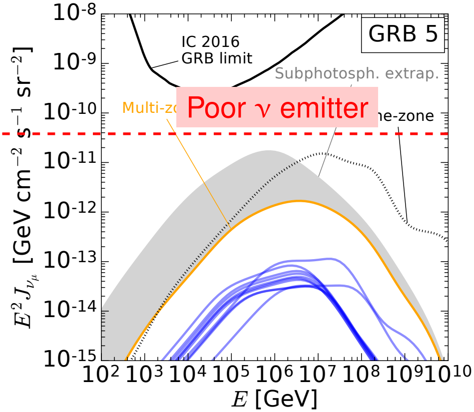
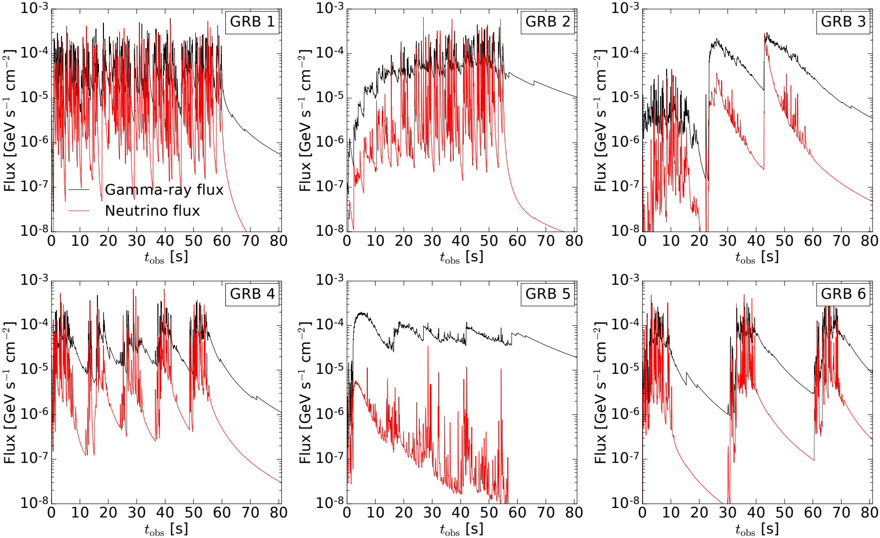




An undisciplined engine makes a GRB neutrino-bright

Using our new insight

MB, Heinze, Murase, Winter, *ApJ* 2017



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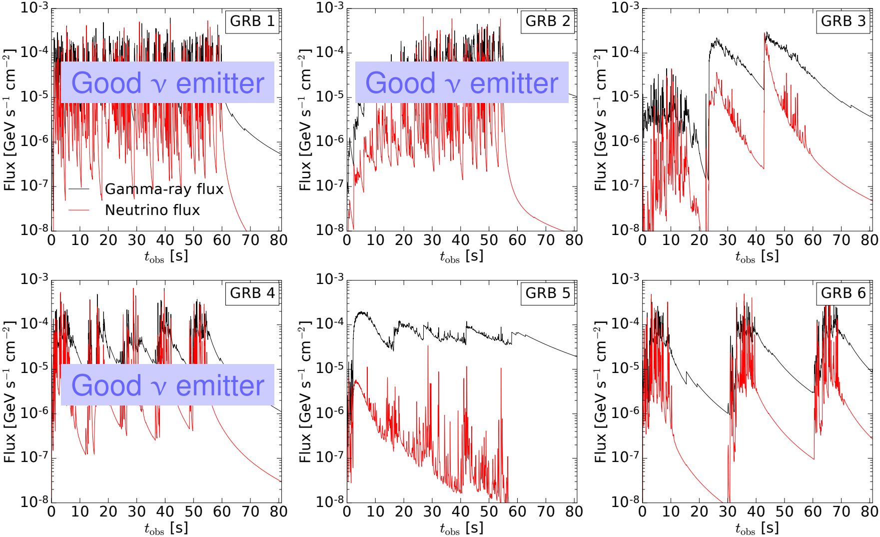
MB, Heinze, Murase, Winter, *ApJ* 2017



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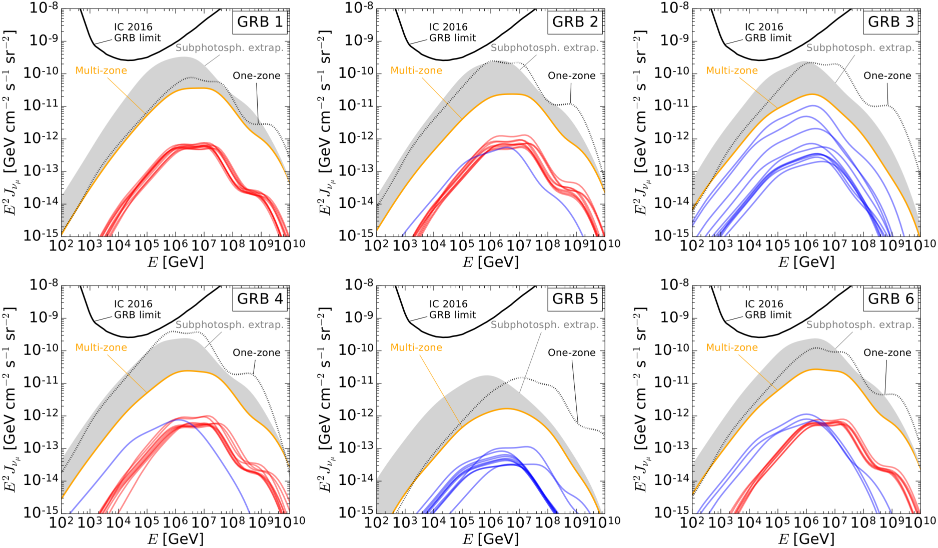
MB, Heinze, Murase, Winter, *ApJ* 2017



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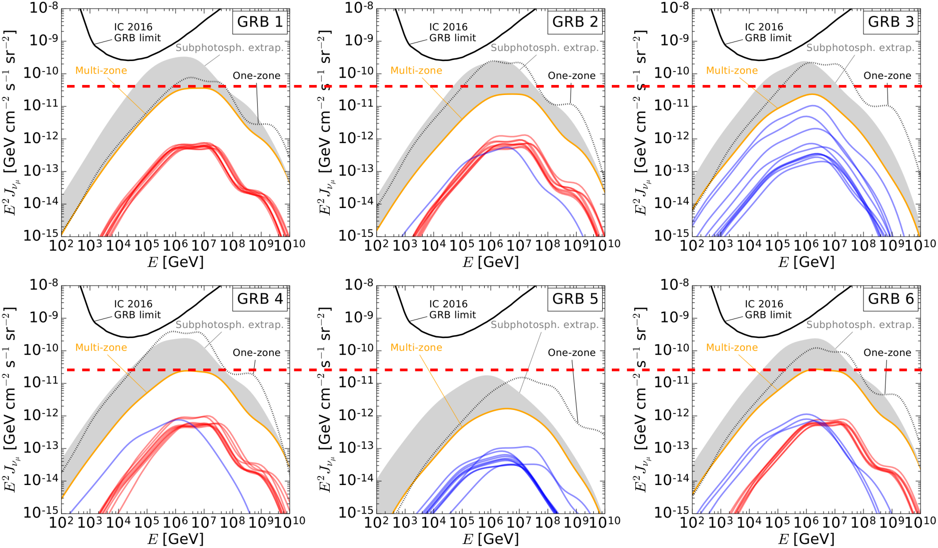

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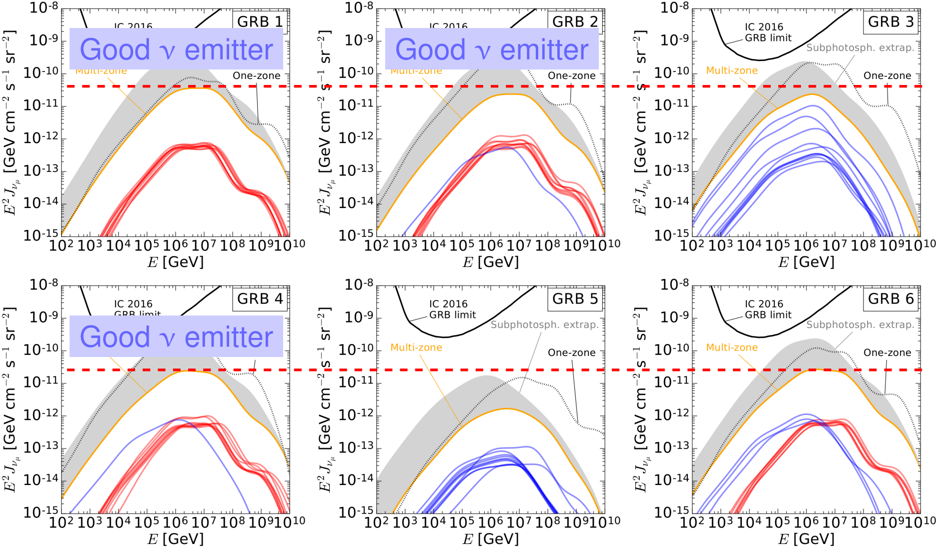

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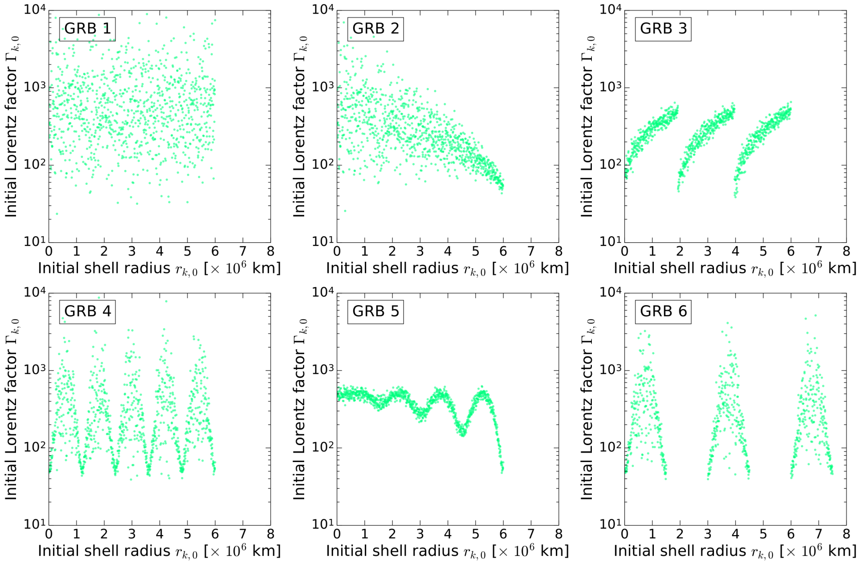
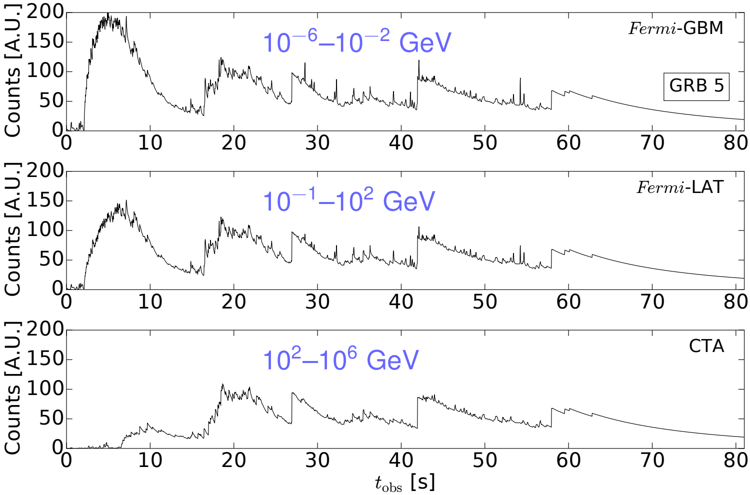
MB, Heinze, Murase, Winter, *ApJ* 2017



Time delays in gamma-ray light curves

Neutrino-weak bursts show time delays in different energy bands —



MB, Heinze, Murase, Winter, *ApJ* 2017 [1606.02325]
See also: Bošnjak, Daigne, *A&A* 2014 [1404.4577]

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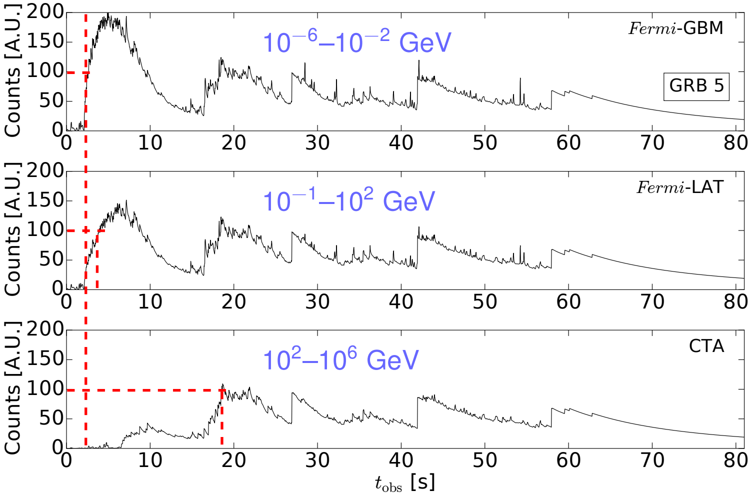
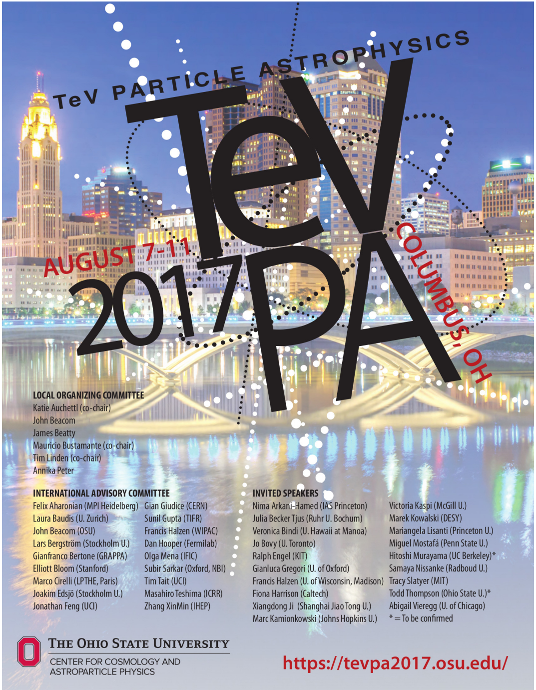


MB, Heinze, Murase, Winter, *ApJ* 2017 [1606.02325]
See also: Bošnjak, Daigne, *A&A* 2014 [1404.4577]

Conclusions . . . and the future

- Neutrino emission from bright GRBs is lower than we expected
- Still candidates for joint e.m.-ν detection, UHECR production
- Possibly the first extragalactic neutrino sources to be resolved
- More realistic understanding of GRB ν production
- We need next-gen neutrino telescopes (IceCube-Gen2, KM3NeT)



TeVPA 2017

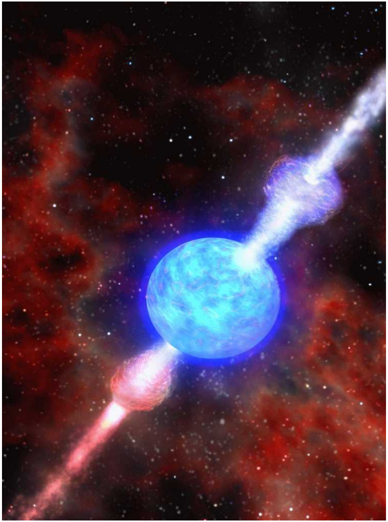
tevpa2017.osu.edu

- August 7–11, Columbus, OH
- Deadline for registration and abstract submission: June 2
- Pre-meeting mini-workshops on Sunday, August 7
- Ample room for parallels: *we welcome your talks!*

Backup slides

GRBs at a glance

- Duration: 0.1 s to > 100 s
- Two populations:
  - Short (< 2 s): NS, BH mergers
  - Long (> 2 s): hypernovae
- Accretion onto a compact object
- Isotropically distributed
- Far: most at a few Gpc ($z \approx 1$)
- Gamma-ray energies: > 1 keV
- Radiated energy: 10^{52}–10^{53} erg



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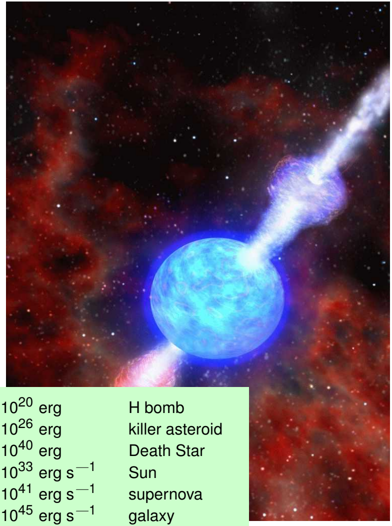
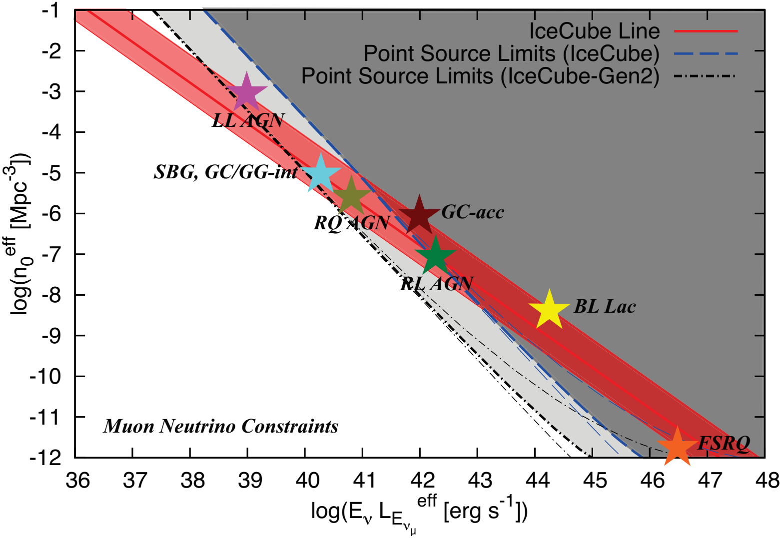


10^{20} erg	H bomb
10^{26} erg	killer asteroid
10^{40} erg	Death Star
10^{33} erg s^{-1}	Sun
10^{41} erg s^{-1}	supernova
10^{45} erg s^{-1}	galaxy

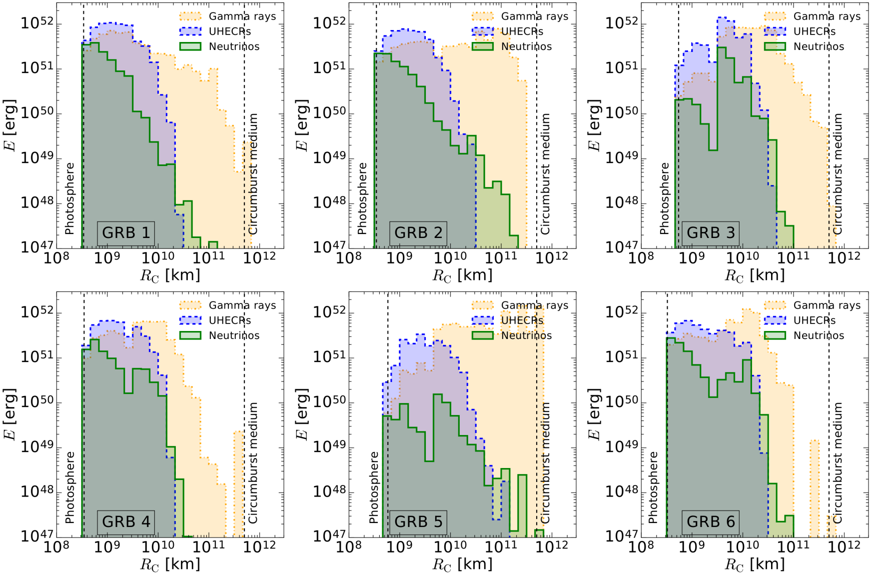
Needed: many weak (steady) sources

So far, neutrinos are *not* correlated to a few bright steady sources —



MURASE & WAXMAN, *PRD* 2016 [1607.01601]

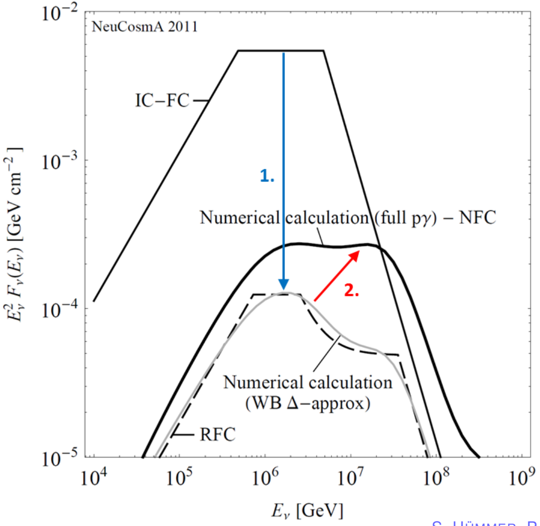
Particle production inside the GRB jet



MB, Heinze, Murase, Winter, *ApJ* 2017 [1606.02325]

Refining the neutrino spectrum — NeuCosmA

More production channels, more complete particle-physics treatment



For example, GRB080603A:

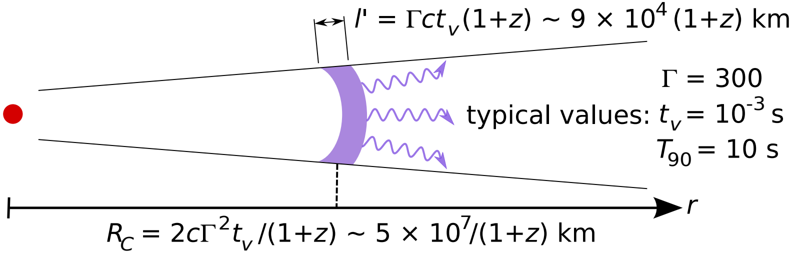
1. Correction to analytical model (IC-FC → RFC)
2. Change due to full numerical calculation

IC-FC: IceCube-Fireball Calculation
RFC: Revised Fireball Calculation
NFC: Numerical Fireball Calculation

S. Hümmer, P. Baerwald, W. Winter, *PRL* **108**, 231101 (2012)

A simplifying assumption: identical collisions

All internal collisions are identical and occur at the same radius —



$N_{\rm coll} \approx T_{90}/t_v \sim 100$–$1000$ identical collisions

- Calculate particle emission spectra once
- Then multiply by $N_{\rm coll}$
- Γ is the average speed of shells

A two-component model of CR emission

Optical depth:

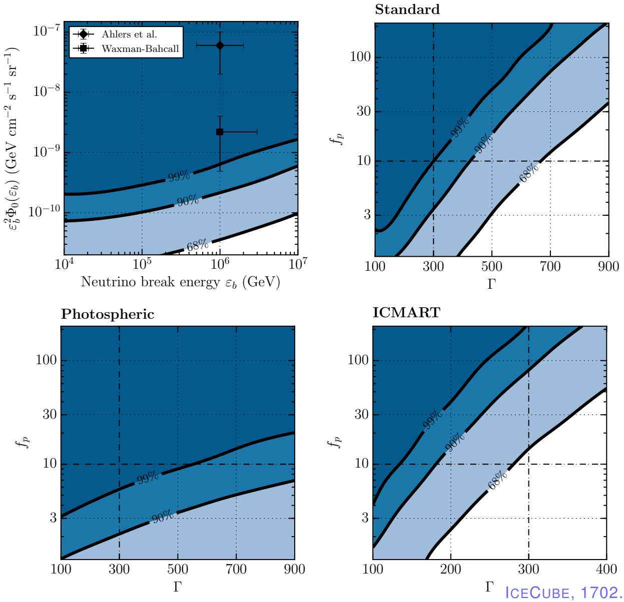
$$\tau_n = \left. \frac{t_{p\gamma}^{-1}}{t_{\text{dyn}}^{-1}} \right|_{E_{p,\max}} = \begin{cases} \lesssim 1\,, & \text{optically } \textbf{thin} \text{ source} \\ > 1\,, & \text{optically } \textbf{thick} \text{ source} \end{cases}$$

Particles can escape from within a shell of thickness λ'_{mfp}:

$$\left.\begin{aligned} \lambda'_{p,\text{mfp}}\left(E'\right) &= \min\left[\Delta r', R'_L\left(E'\right), ct'_{p\gamma}\left(E'\right)\right] \\ \lambda'_{n,\text{mfp}}\left(E'\right) &= \min\left[\Delta r', ct'_{p\gamma}\left(E'\right)\right] \end{aligned}\right\} f_{\text{esc}} = \frac{\lambda'_{\text{mfp}}}{\Delta r'}$$

fraction of escaping particles

Limits on GRB ν production model parameters



IceCube, 1702.06868

NeuCosmA: (revised) GRB particle emission – I

Two ingredients:

$$\underbrace{N'_p\left(E'_p\right)}_{\text{proton density at the source [GeV}^{-1}\text{ cm}^{-3}]} \overset{\text{NeuCosmA}}{\otimes} \underbrace{N'_\gamma\left(E'_\gamma\right)}_{\text{photon density at the source}}$$

$$= \underbrace{Q'_\nu\left(E'_\nu\right)}_{\text{emitted neutrino spectrum [GeV}^{-1}\text{ cm}^{-3}\text{ s}^{-1}]}$$

- Photons (same shape as observed at Earth):

$$N'_\gamma\left(E'_\gamma\right) = \begin{cases} \left(E'_\gamma/E'_{\gamma,\text{break}}\right)^{-1} & , \; E'_{\gamma,\text{min}} = 0.2\text{ eV} \leqslant E'_\gamma < E'_{\gamma,\text{break}} = 1\text{ keV} \\ \left(E'_\gamma/E'_{\gamma,\text{break}}\right)^{-2.2} & , \; E'_\gamma \geqslant E'_{\gamma,\text{break}} \\ 0 & , \; \text{otherwise} \end{cases}$$

- Protons: $N'_p\left(E'_p\right) \propto E_p'^{-\alpha_p} e^{-E'_p/E'_{p,\text{max}}} \quad \left(\alpha_p \gtrsim 2\right)$

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$$t'_{\text{acc}}\left(E'_{p,\text{max}}\right) = \min\left[t'_{\text{dyn}}, t'_{\text{syn}}\left(E'_{p,\text{max}}\right), t'_{p\gamma}\left(E'_{p,\text{max}}\right)\right]$$

- Protons: $N'_p\left(E'_p\right) \propto E_p'^{-\alpha_p} e^{-E'_p/E'_{p,\text{max}}} \quad (\alpha_p \gtrsim 2)$

NeuCosmA: (revised) GRB particle emission – II

Normalize the particle densities at the source —

- Photons:

$$\underbrace{\text{photon energy density per collision}}_{\int E'_\gamma\, N'_\gamma\left(E'_\gamma\right)\, dE'_\gamma} = \frac{\overbrace{\text{total gamma-ray energy of burst}}^{E^{\text{iso},\prime}_{\gamma,\text{tot}} \sim 10^{53}\ \text{erg (from observed fluence)}}}{\underbrace{\text{number of collisions}}_{N_{\text{coll}}} \cdot \underbrace{\text{volume of one collision}}_{V'_{\text{iso}}}}$$

- Protons:

baryonic loading (energy in p's / energy in e's + γ's), *e.g.*, 10

$$\underbrace{\text{proton energy density per collision}}_{\int E'_p\, N'_p\left(E'_p\right)\, dE'_p} = \frac{1}{f_e} \cdot \text{photon energy density per collision}$$

NeuCosmA: (revised) GRB particle emission – III

Injected/ejected spectrum of secondaries (π, K, n, ν, etc.):

$x \equiv E'/E'_p$

$y \equiv E'_p E'_\gamma / \left(m_p c^2\right)$

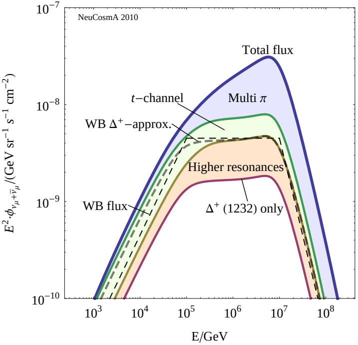
$$Q'\left(E'\right) = \int_{E'}^{\infty} \frac{dE'_p}{E'_p} N'_p\left(E'_p\right) \int_0^{\infty} c\, dE'_\gamma\, N'_\gamma\left(E'_\gamma\right) R\left(x, y\right)$$

response function

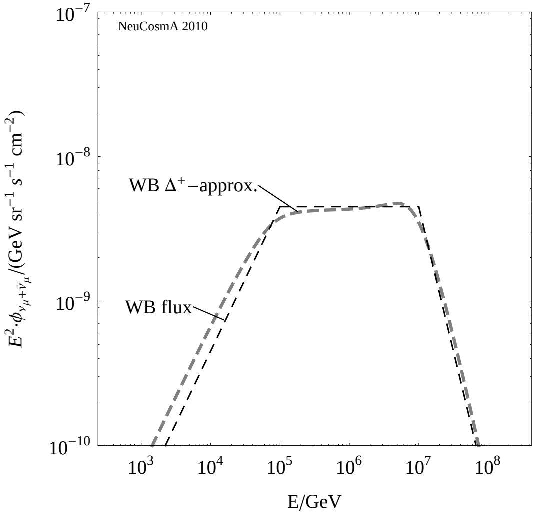
R contains cross sections, multiplicities for different channels

What does NeuCosmA include?

- $p\gamma \to \Delta^+(1232) \to \pi^0, \pi^+, \ldots$
- extra K, n, π^-, multi-π prod. modes
- synchrotron losses of secondaries
- adiabatic cooling
- full photon spectrum
- neutrino flavor transitions



NeuCosmA – the full photohadronic cross section

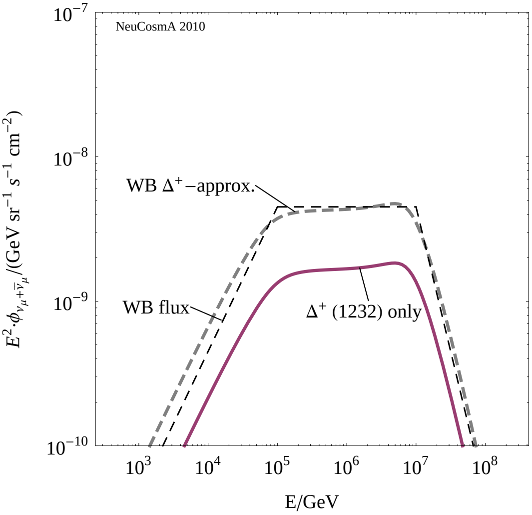


NeuCosmA – the full photohadronic cross section

Contributions to $(\nu_\mu + \bar{\nu}_\mu)$ flux from $\pi^\pm$ decay divided in:

- $\Delta(1232)$-resonance

P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

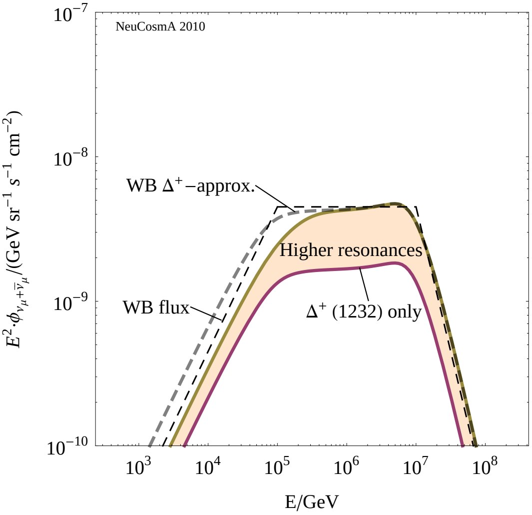


NeuCosmA – the full photohadronic cross section

Contributions to $(\nu_\mu + \bar{\nu}_\mu)$ flux from $\pi^\pm$ decay divided in:

- $\Delta(1232)$-resonance
- Higher resonances

P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)



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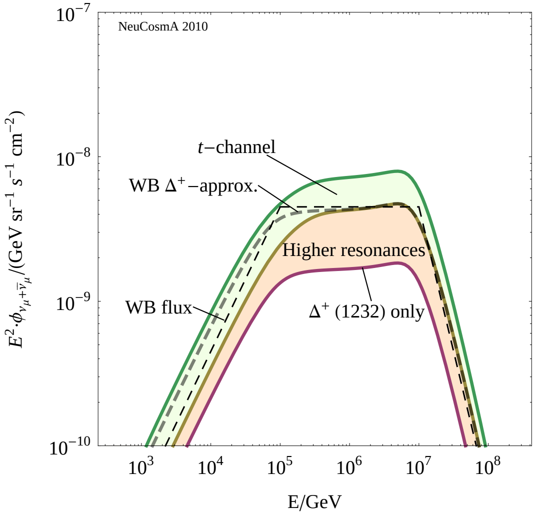
P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

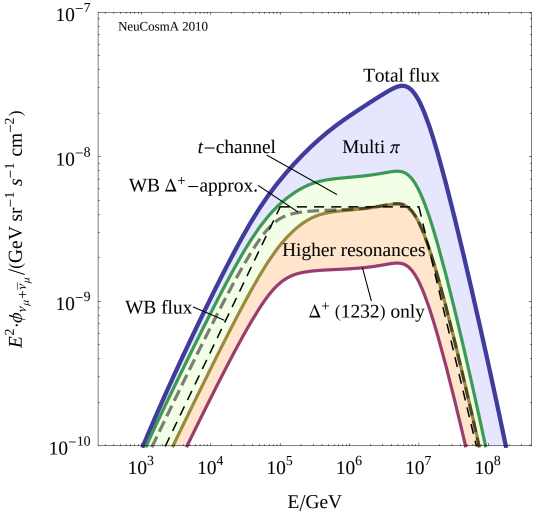


NeuCosmA – the full photohadronic cross section

Contributions to $(\nu_\mu + \bar{\nu}_\mu)$ flux from $\pi^\pm$ decay divided in:

- $\Delta(1232)$-resonance
- Higher resonances
- *t*-channel (direct production)
- High energy processes (multiple π)

P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)



Especially "Multi π" contribution leads to change of flux shape; neutrino flux higher by up to a factor of 3 compared to WB treatment

NeuCosmA – further particle decays

$$
\begin{aligned}
\pi^+ &\rightarrow \mu^+ + \nu_\mu \\
&\quad \mu^+ \rightarrow e^+ + \nu_e + \bar{\nu}_\mu \\
\\
\pi^- &\rightarrow \mu^- + \bar{\nu}_\mu \\
&\quad \mu^- \rightarrow e^- + \bar{\nu}_e + \nu_\mu \\
\\
K^+ &\rightarrow \mu^+ + \nu_\mu \\
\\
n &\rightarrow p + e^- + \bar{\nu}_e
\end{aligned}
$$

NeuCosmA – further particle decays

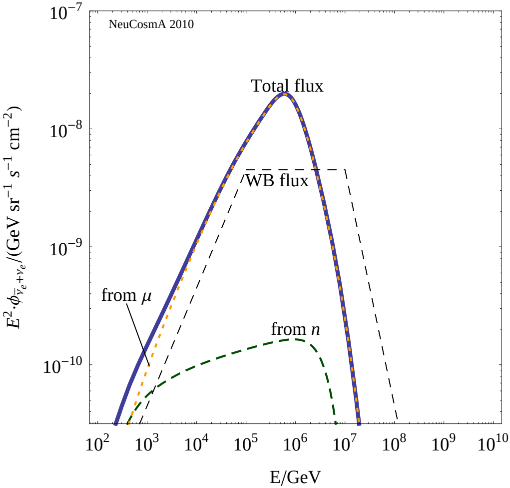
$$\pi^+ \rightarrow \mu^+ + \nu_\mu$$
$$\mu^+ \rightarrow e^+ + \nu_e + \bar{\nu}_\mu$$

$$\pi^- \rightarrow \mu^- + \bar{\nu}_\mu$$
$$\mu^- \rightarrow e^- + \bar{\nu}_e + \nu_\mu$$

$$K^+ \rightarrow \mu^+ + \nu_\mu$$

$$n \rightarrow p + e^- + \bar{\nu}_e$$

Resulting ν_e flux (at the observer)



P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

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$$\mu^- \rightarrow e^- + \bar{\nu}_e + \nu_\mu$$

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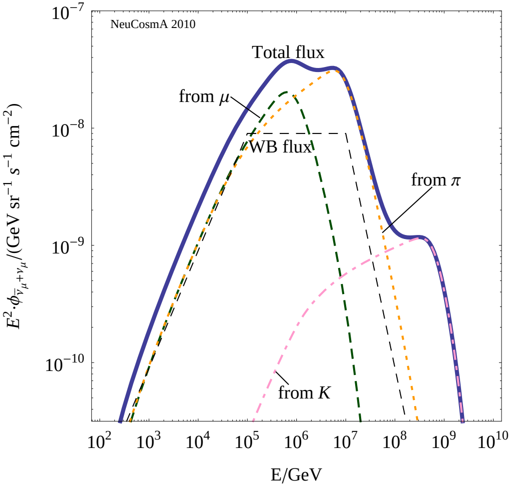
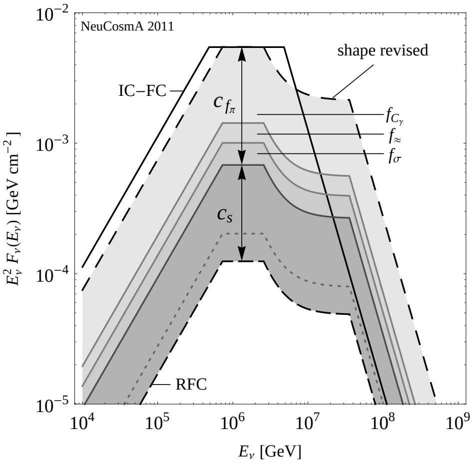
Resulting ν_μ flux (at the observer)



P. BAERWALD, S. HÜMMER, AND W. WINTER, *Phys. Rev.* **D83**, 067303 (2011)

NeuCosmA – how the neutrino spectrum changes – I



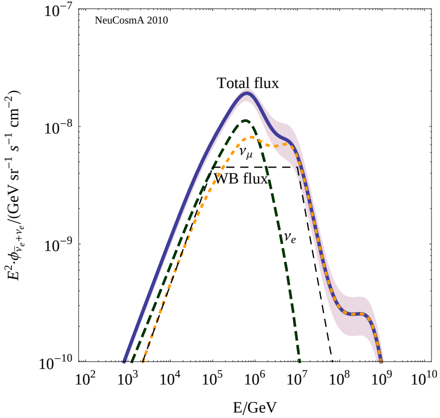
S. HÜMMER, P. BAERWALD, W. WINTER, *Phys. Rev. Lett.* **108**, 231101 (2012)

Corrections to the analytical model:

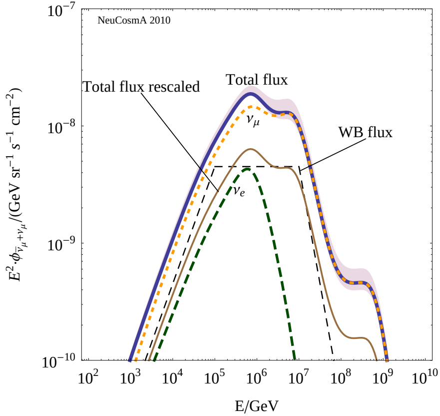
- shape revised:
  - shift of first break (correction of photohadronic threshold)
  - different cooling breaks for μ's and π's
  - $(1+z)$ correction on the variability scale of the GRB
- Correction cf_π to π prod. efficiency:
  - f_{C_γ}: full spectral shape of photons
  - $f_\approx = 0.69$: rounding error in analytical calculation
  - $f_\sigma \simeq 2/3$: from neglecting the width of the Δ-resonance
- Correction c_S:
  - energy losses of secondaries
  - energy dependence of the mean free path of protons

NeuCosmA – neutrino spectra including flavor mixing

Electron neutrino spectrum



Muon neutrino spectrum



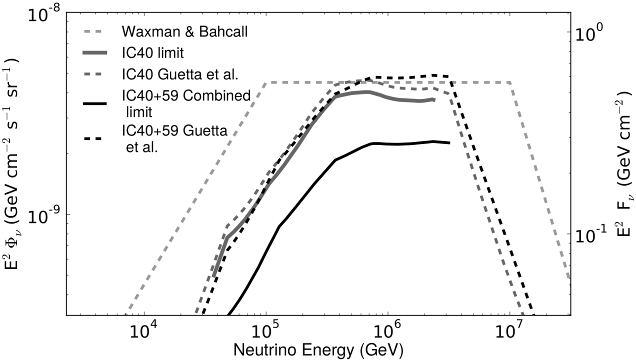
P. Baerwald, S. Hümmer, W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

Characteristic double peak structure from μ and π decay in both flavors, additional peak from K^+ decay at 10^8 to 10^9 GeV

Neutron model of UHECR emission under tension?

In 2012, IceCube ruled this analytical version of the fireball model –

- assumed a fixed baryonic loading of 10
- extrapolated diffuse ν flux from 117–215 GRBs ("quasi-diffuse")
- analytical calculation – in tension with upper bounds



ICECUBE, *Nature* **484**, 351 (2012)

M. AHLERS *et al. Astropart. Phys.* **35**, 87 (2011)

D. GUETTA *et al. Astropart. Phys.* **20**, 429 (2004)

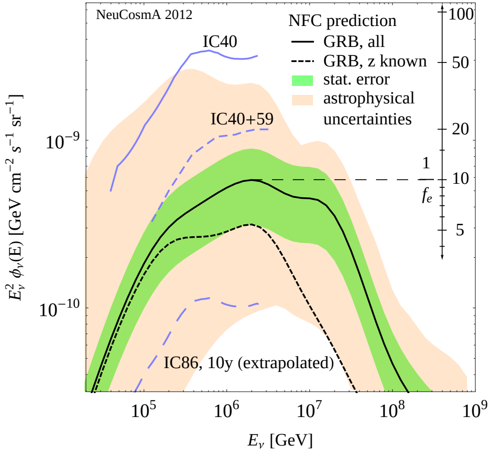
The new prediction of the quasi-diffuse GRB ν flux

Repeat the IceCube GRB neutrino analysis, with NeuCosmA —

- Same GRB sample and parameters
- Calculate ν fluence for each burst and stacked fluence $F_\nu(E_\nu)$
- Quasi-diffuse flux ($N_{\mathrm{GRB}} = 117$):

$$\phi_\nu(E_\nu) = F_\nu(E_\nu)\,\frac{1}{4\pi}\,\frac{1}{N_{\mathrm{GRB}}}\,\frac{667\ \text{bursts}}{\text{yr}}$$

Flux ~ 1 order of magnitude lower!

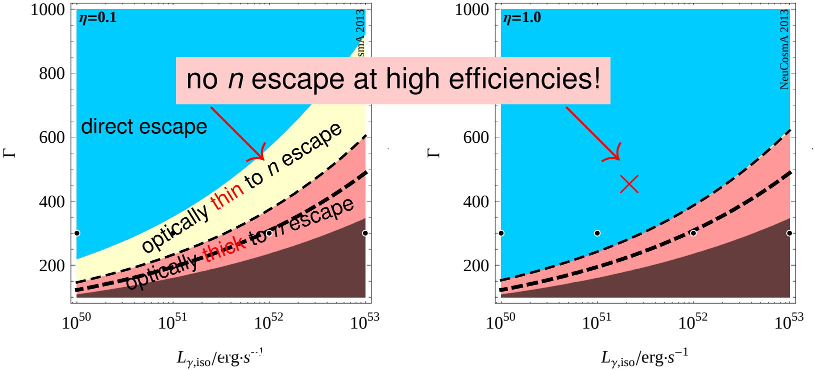


S. Hümmer, P. Baerwald, W. Winter, *PRL* **108**, 231101 (2012)

We *need* direct proton escape

Scan of the GRB emission parameter space –

acceleration efficiency ⟶ $\eta = 0.1$ $\eta = 1.0$



P. Baerwald, MB, and W. Winter, *ApJ* **768**, 186 (2013)

we need high efficiencies ⇒ direct proton escape *is* required

Cosmogenic neutrinos

We have seen that protons interact with the cosmological photon fields (CMB, etc.), *e.g.*,

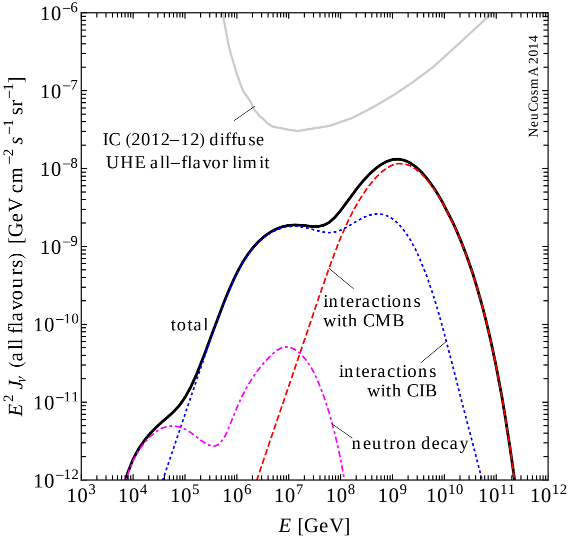
$$p + \gamma \to \Delta^+ \to \pi^+ + n\,,$$

and neutrinos are created in the decays of the secondaries:

$$\pi^+ \to \mu^+ + \nu_\mu$$
$$\mu^+ \to \bar{\nu}_\mu + \nu_e + e^+$$
$$n \to p + e^- + \bar{\nu}_e$$

These are called *cosmogenic neutrinos*

Cosmogenic neutrinos



P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)

ν's in the GRB internal shock model

Secondary injection of neutrons, neutrinos ($\mathrm{GeV}^{-1}\ \mathrm{cm}^{-3}\ \mathrm{s}^{-1}$)

$$Q'\left(E'\right) = \int_{E'}^{\infty} \frac{dE'_p}{E'_p} N'_p\left(E'_p\right) \int_0^{\infty} c\,d\varepsilon' N'_\gamma\left(\varepsilon'\right) R\left(E', E'_p, \varepsilon'\right)$$

Normalisation to the observed GRB photon flux F_γ

$$\int \varepsilon' N'_\gamma\left(\varepsilon'\right) d\varepsilon' = \frac{E'^{\mathrm{sh}}_{\mathrm{iso}}}{V'_{\mathrm{iso}}} \propto F_\gamma\ , \quad \int E'_p N'_p\left(E'_p\right) dE'_p = \frac{1}{f_e} \frac{E'^{\mathrm{sh}}_{\mathrm{iso}}}{V'_{\mathrm{iso}}} \propto \frac{F_\gamma}{f_e}$$

Fluence per shell, at Earth ($\mathrm{GeV}^{-1}\ \mathrm{cm}^{-2}$)

$$\mathcal{F}^{\mathrm{sh}} = t_v V'_{\mathrm{iso}} \frac{(1+z)^2}{4\pi d_L^2} Q'$$

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- Photon density, shock rest frame (GeV^{-1} cm^{-3}):

$$N'_\gamma\left(\varepsilon'\right) \propto \begin{cases} \left(\varepsilon'\right)^{-\alpha_\gamma}, & \varepsilon'_{\gamma,\text{min}} = 0.2\ \text{eV} \leqslant \varepsilon' \leqslant \varepsilon'_{\gamma,\text{break}} \\ \left(\varepsilon'\right)^{-\beta_\gamma}, & \varepsilon'_{\gamma,\text{break}} \leqslant \varepsilon' \leqslant \varepsilon'_{\gamma,\text{max}} = 300 \times \varepsilon'_{\gamma,\text{min}} \end{cases}$$

$$\varepsilon'_{\gamma,\text{break}} = \mathcal{O}\left(\text{keV}\right), \alpha_\gamma \approx 1, \beta_\gamma \approx 2$$

- Proton density:

$$N'_p\left(E'_p\right) \propto \left(E'_p\right)^{-\alpha_p} \times \exp\left[-\left(E'_p/E'_{p,\text{max}}\right)^2\right] \quad \left(\alpha_p \approx 2\right)$$

Maximum proton energy limited by energy losses:

$$t'_{\text{acc}}\left(E'_{p,\text{max}}\right) = \min\left[t'_{\text{dyn}}, t'_{\text{syn}}\left(E'_{p,\text{max}}\right), t'_{p\gamma}\left(E'_{p,\text{max}}\right)\right]$$

ν's in the GRB internal shock model

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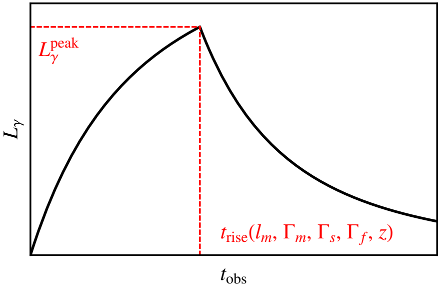
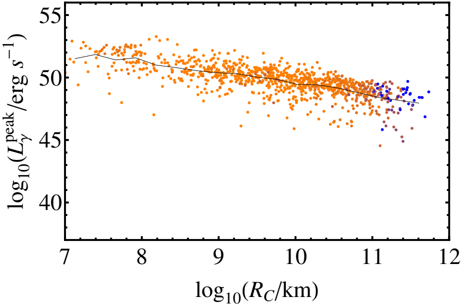
Gamma-ray and neutrino pulses

A fast-rise-exponential-decay (FRED) gamma-ray pulse is emitted in every collision:

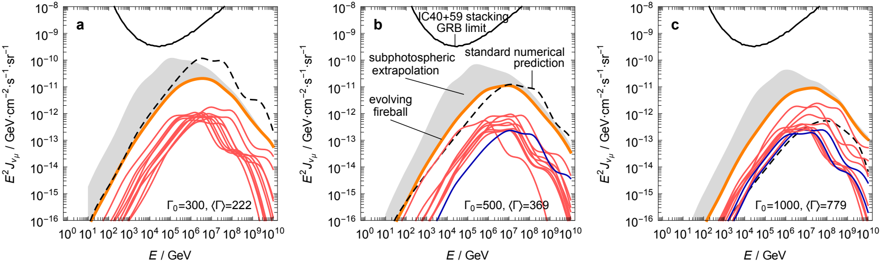




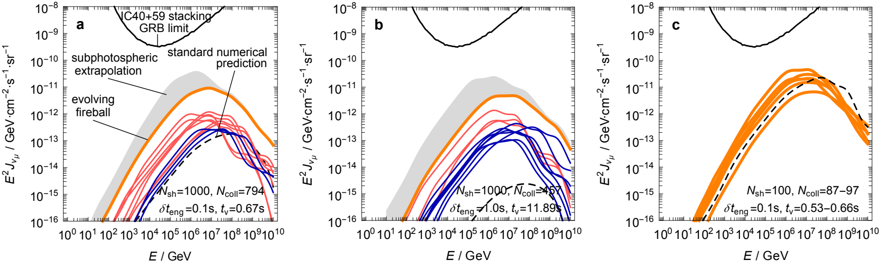
$L_\gamma^{\rm peak} \sim R_C^{-2}$

The prediction *is* robust

Simulations show only weak dependence of the **flux** on the boost Γ . . .

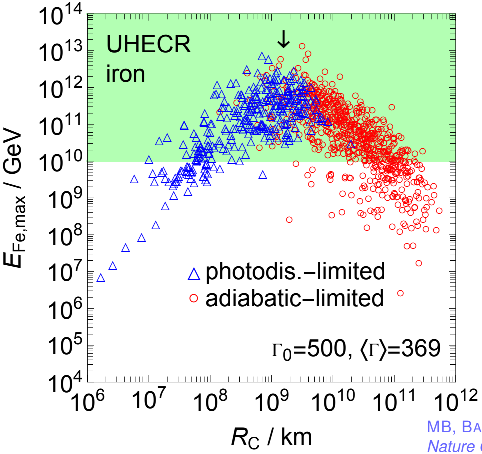


. . . and on the GRB engine variability time δt_{eng}



Accelerating iron

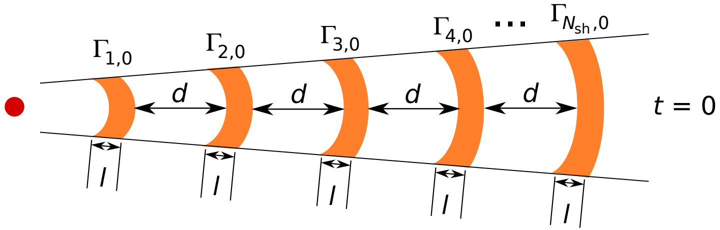
- Photodisintegration destroys nuclei close to the center ($\sim 10^8$ km) *e.g.*, ANCHORDOQUI *et al.*, *Astropart. Phys.* **29**, 1 (2008)
- However, they can survive at large radii:



MB, BAERWALD, MURASE, WINTER *Nature Commun.* **6**, 6783 (2015)

Initialising the burst simulation

Initial number of plasma shells in the jet: $\gtrsim 1000$



Initial values of shell parameters:

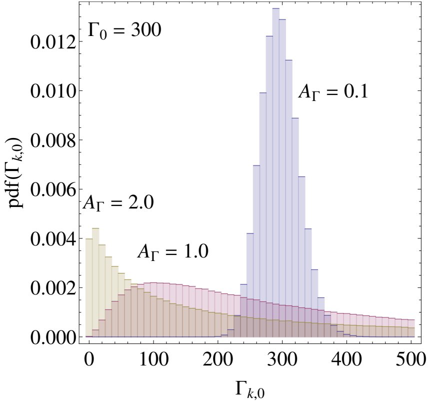
- Width of shells and separation between them: $l = d = c \cdot \delta t_{eng}$
- Equal kinetic energy for all shells ($\sim 10^{52}$ erg)
- Shell speeds $\Gamma_{k,0}$ follow a distribution (log-normal or other)

Initial distribution of shell speeds

Distribution of initial shell speeds (Lorentz factors):

$$\ln\left(\frac{\Gamma_{k,0}-1}{\Gamma_0-1}\right) = A_\Gamma \cdot x$$

x follows a Gaussian distribution, $P(x)\,dx = dx\; e^{-x^2/2}/\sqrt{2\pi}$



$A_\Gamma < 1$

Speeds too similar, collisions only at large radii

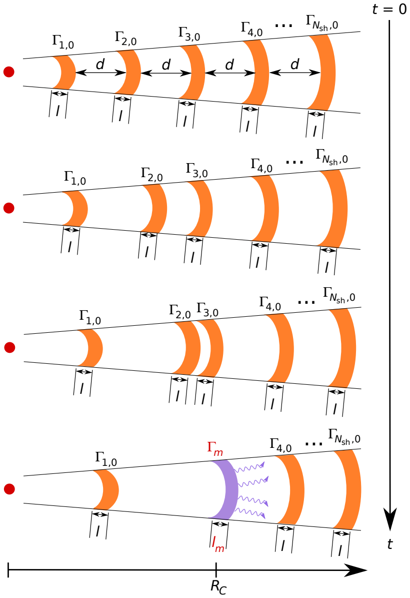
$A_\Gamma \gg 1$

Spread too large, too many collisions at low radii

$A_\Gamma \approx 1$

Just right, burst has high efficiency of conversion of kinetic to radiated energy

Propagating and colliding the shells



During propagation:

- speeds, masses, widths do not change (only in collisions)
- the new, merged shells continue propagating and can collide again

Evolution stops when either:

- a single shell is left; or
- all remaining shells have reached the circumburst medium ($\gtrsim 6 \times 10^{11}$ km)

final number of collisions
$\approx$
number of initial shells ($\gtrsim 1000$)

S. Kobayashi, T. Piran, R. Sari, *ApJ* **490**, 92 (1997)
F. Daigne, R. Mochkovitch, *MNRAS* **296**, 275 (1998)

How is the new prediction different?

- The top-contributing collisions are at the photosphere
- Pion production efficiency there is independent of Γ:

$$f_{p\gamma}^{\rm ph} \sim 5 \cdot \frac{\varepsilon}{0.25} \cdot \frac{\epsilon_e}{0.1} \cdot \frac{1~{\rm keV}}{\epsilon'_{\gamma,{\rm break}}}$$

  ε: energy dissipation efficiency
  ϵ_e: fraction of dissipated energy as e.m. output (photons)

- $\Rightarrow$ Time-integrated neutrino fluence dominated is independent of Γ:

$$\mathcal{F}_\nu \propto \frac{N_{\rm coll}\left(f_{p\gamma} \gtrsim 1\right)}{N_{\rm coll}^{\rm tot}} \times \min\left[1, f_{p\gamma}^{\rm ph}\right] \times \frac{\epsilon_p}{\epsilon_e} \times E_{\gamma-{\rm tot}}^{\rm iso}$$

- Compare to standard predictions, which have a $\langle\Gamma\rangle^{-4}$ dependence
- Raising ϵ_p automatically decreases ϵ_e, so the photosphere grows, but still ~ 10 photospheric collisions dominate

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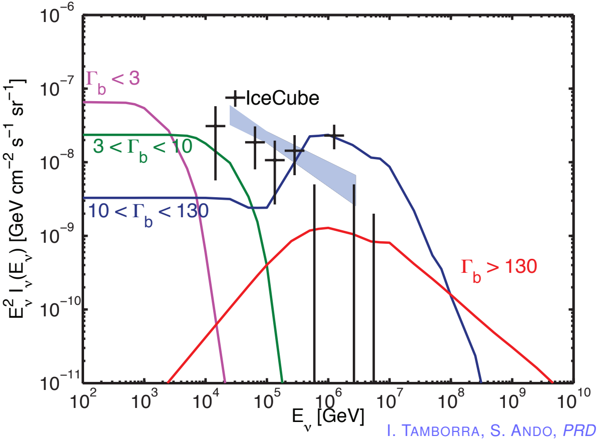
- Compare to standard predictions, which have a $\langle\Gamma\rangle^{-4}$ dependence
- Raising ϵ_p automatically decreases ϵ_e, so the photosphere grows, but still $\sim$ 10 photospheric collisions dominate

What about low-luminosity and choked GRBs?

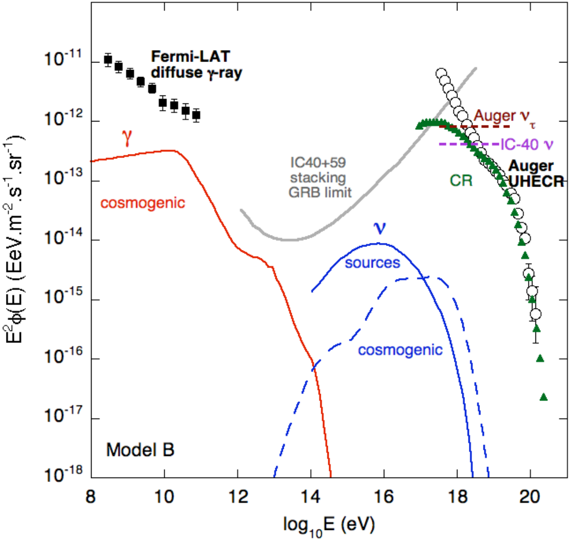
- Low-luminosity and choked GRBs might be in the same family as high-luminosity long GRBs
- Due to lower jet speeds (Γ_b), they do not break out
- They might explain the TeV region of the IceCube diffuse ν flux:



I. Tamborra, S. Ando, *PRD* **93**, 053010 (2016)

A detailed treatment of UHECR acceleration

From Globus, Allard, Mochkovitch, Parizot *MNRAS* 2015 —



- Multiple internal collisions
- Numerical simulation of CR acceleration at mildly relativistic shocks
- Accelerated *p* to Fe
- ν spectrum harder as shock propagates
- Results compatible with Bustamante+ 2015