Gamma-ray bursts: high-energy neutrino predictions in the IceCube era

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Two fifty-year-old mysteries

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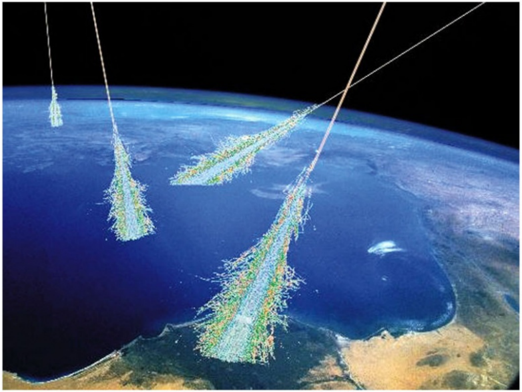
ultra-high-energy cosmic rays (UHECRs)



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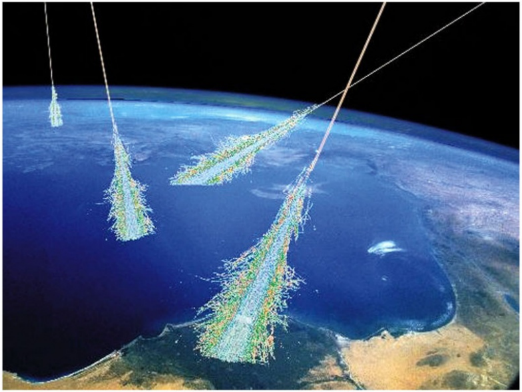

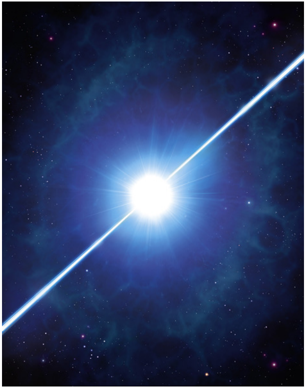

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gamma-ray bursts (GRBs)

The mystery

We do not know the origin of UHECRs and GRBs

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GRBs are the sources of the UHECRs
– and neutrinos are the smoking gun

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GRBs are the sources of the UHECRs
– and neutrinos are the smoking gun

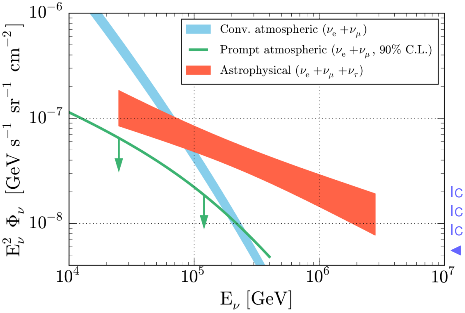
Our result

It is possible, *and testable*, but the connection between UHECRs, GRBs, and neutrinos is **not** as simple as we thought

A new player: high-energy astrophysical neutrinos

The era of neutrino astronomy has begun!

– IceCube has reported 54 events with 30 TeV – 2 PeV in 4 years

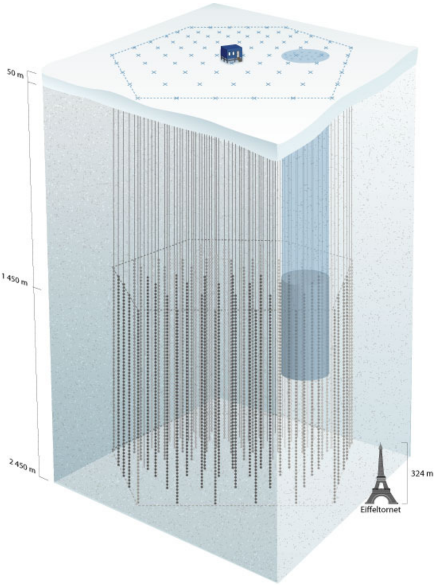


IceCube, *PRL* **111**, 021103 (2013)
IceCube, *Science* **342**, 1242856 (2013)
IceCube, *PRL* **113**, 101101 (2014)
◀ IceCube, *ApJ* **809**, 98 (2015)

Diffuse per-flavor astrophysical flux [IceCube 2015]:

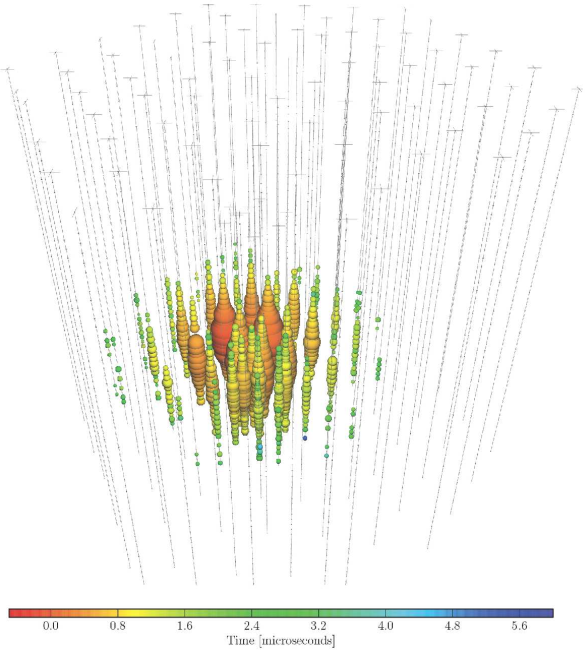
$$\Phi_\nu = \left(6.7^{+1.1}_{-1.2} \cdot 10^{-18}\right) \left(\frac{E}{100\ \text{TeV}}\right)^{-(2.5\pm0.09)} \text{GeV}^{-1}\ \text{cm}^{-2}\ \text{s}^{-1}\ \text{sr}^{-1}$$

Detecting the neutrinos: IceCube



IceCube: km^3 in-ice South Pole Čerenkov detector

- νN interactions ($N = n, p$) create particle showers
- 86 strings with 5160 digital optical modules (DOMs)
- depths between 1450 m and 2450 m


0.0
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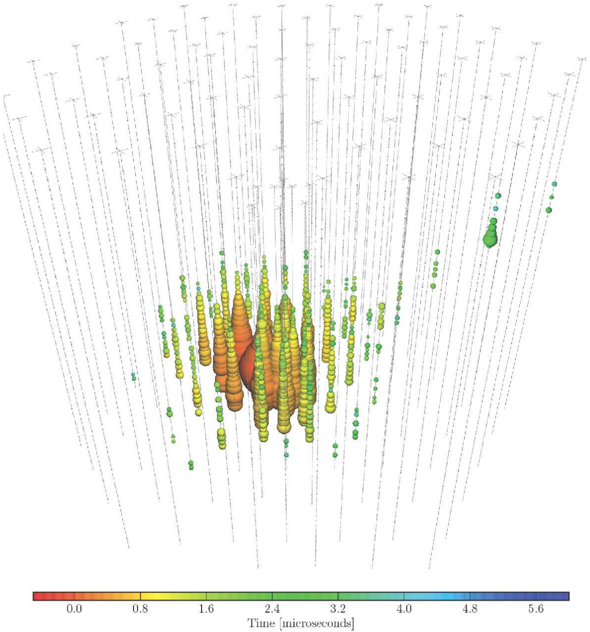
3.2
4.0
4.8
5.6
Time [microseconds]


0.0
0.8
1.6
2.4
3.2
4.0
4.8
5.6
Time [microseconds]

What we know / don't know

What we know

- compatible with isotropy
- power-law $\propto E^{-2.5}$
- not coincident with transient sources (*e.g.*, GRBs)
- not correlated with known sources
- flavor composition: compatible with equal proportion of ν_e, ν_μ, ν_τ
- also: no prompt atmospheric neutrinos

What we don't know

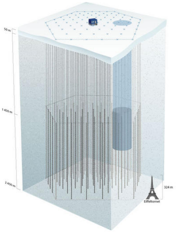
- **what are the sources?**
- what is the production mechanism (pp, $p\gamma$)?
- is there a cut-off at 2 PeV?
- what is the Galactic contribution, if any?
- **what is the precise relation to UHE cosmic rays?**
- is there new physics?
- what is the precise flavor composition of the flux?

. . . but we have good ideas on all

Why is *now* a good time to do this?

better, bigger detectors + loads of data + bright future

neutrinos



- IceCube: diffuse flux of HE astrophysical ν's
- No point sources *yet*
- GRBs: low bg due to time and direction cuts
- IceCube-Gen2

GRBs

- *Fermi*: $\sim$ 250 GRBs yr^{-1} in 8 keV – 40 MeV
- $\sim$ 12 GRBs yr^{-1} in 20 MeV – 300 GeV
- different wavelengths: INTEGRAL, *Swift*
- 1000's GRBs detected so far

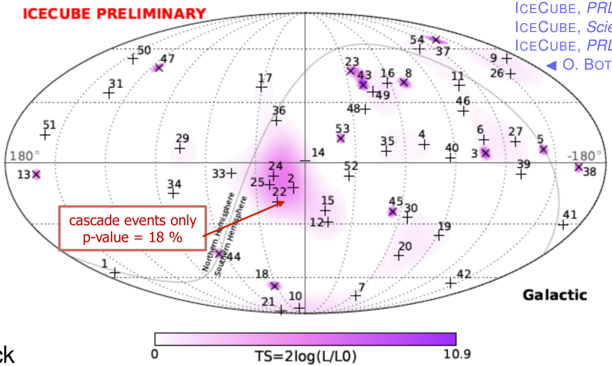
UHECRs

- Auger: 69 events > 57 EeV
- Telescope Array: 72 events
- surface + fluorescence
- future: LHAASO, JEM-EUSO

What about the sources of HE neutrinos?

Arrival directions compatible with an isotropic distribution –

IceCube, *PRL* **111**, 021103 (2013)
IceCube, *Science* **342**, 1242856 (2013)
IceCube, *PRL* **113**, 101101 (2014)
◀ O. Botner, IPA 2015



+: shower
×: muon track

– no association with sources found yet

- GRB searches by IceCube (2012, 2014): no associated ν's found
- GRBs contribute only few % of the observed diffuse flux

[I. Tamborra, S. Ando, *JCAP* **1509**, 036 (2015)]

Why is it still interesting to study GRBs?

1. They are the best candidates for detection of coincident e.m.–neutrino emission at 100's TeV – PeV ν energies
2. Neutrinos from GRB afterglows are expected to be important at EeV energies (observable by radio neutrino detectors)

   [E. Waxman, J. N. Bahcall, *PRL* **78**, 2292 (1997)] [K. Murase, *PRD* **76**, 123001 (2007)]
3. Dark, "failed" GRBs might contribute an important part to the diffuse flux seen by IceCube

   [P. Mészaros, E. Waxman, *PRL* **87**, 171102 (2001)]

Here we will focus on issue 1
and explore the prompt emission of neutrinos in GRBs

GRBs – good candidates for UHE CR & ν sources

GRBs are among the best candidate sources for CRs *and* ν's:

- radiated energy of $\sim 10^{52} - 10^{53}$ erg
- intense magnetic fields of $\sim 10^5$ G
- magnetically-confined *p*'s shock-accelerated to $\sim 10^{12}$ GeV
- plus: low backgrounds (for ν's) due to small time window

Problem: experiments (IceCube, ANTARES) are starting to strongly constrain the simplest joint emission models

Solution: we need to build more realistic models!

GRBs – good candidates for UHE CR & ν sources

10^{20} erg		H bomb
10^{26} erg		killer asteroid
10^{40} erg		Death Star
10^{33} erg s^{-1}		Sun
10^{41} erg s^{-1}		supernova
10^{45} erg s^{-1}		galaxy

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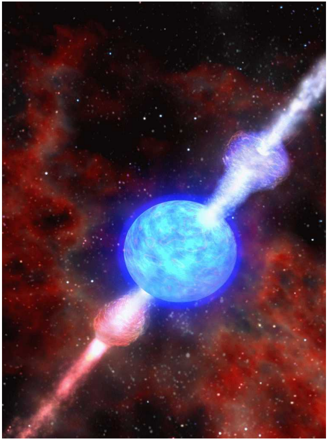
GRBs – what are they?

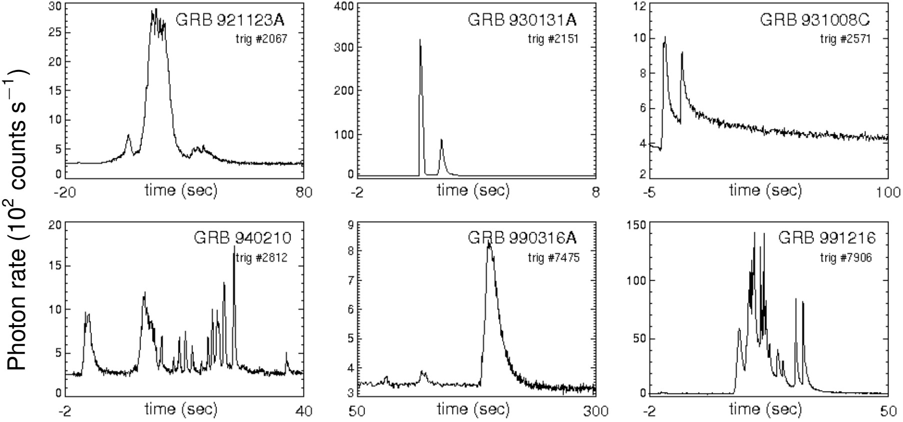
GRBs: the most luminous explosions in the Universe

- **brief** flashes of gamma rays: from 0.1 s to few 100's s
- isotropically distributed in the sky
- they are **far**: most occur at $\sim$ 1 Gpc from us ($z \approx 2$)
- they are **rare**: $\sim$ 0.3 Gpc^{-3} yr^{-1}
- two populations:
  - **short-duration ($<$ 2 s):** neutron star-neutron star or NS-black hole mergers
  - **long-duration ($>$ 2 s):** associated to hypernovae
- powered by matter accretion onto a black hole



GRBs – a zoo of light curves

GRB light curves come in different shapes:

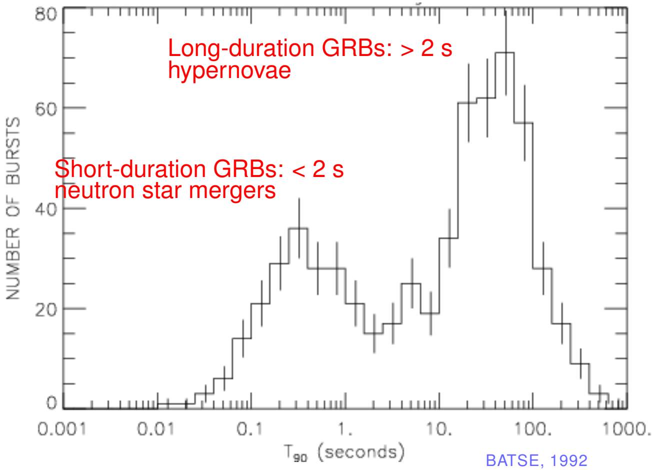


BATSE

variability timescale (width of pulses) $\equiv t_v \approx 1$ ms

GRBs – two populations

Two populations of GRBs:

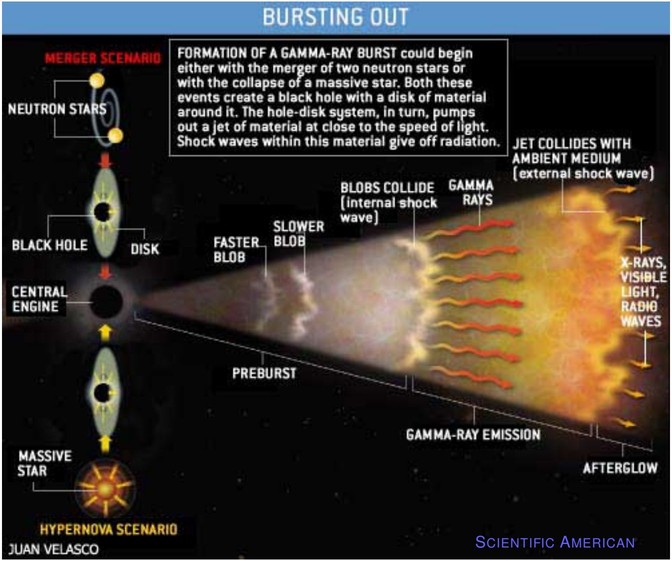


BATSE, 1992

T_{90}: time during which 90% of gamma-ray energy is recorded

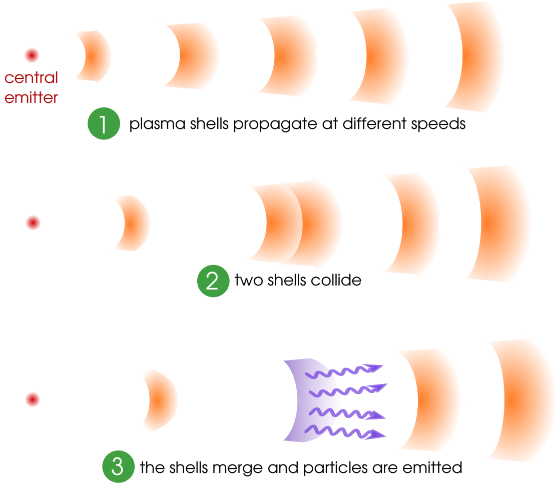
GRBs explained – the fireball model

Developed by Mészáros, Stecker, Piran, Waxman, *et al.* in the 1990s



Internal collisions

Relativistically-expanding blobs of plasma collide with each other, merge, and emit UHE particles —



Producing the UHE ν's, CRs, γ rays – a first look

Joint production of UHECRs, ν's, and γ's:

power law $\sim E^{-\alpha_p}$ — broken power law

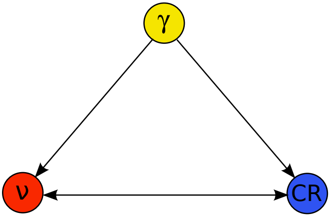
$$p\gamma \rightarrow \Delta^+(1232) \rightarrow \begin{cases} n\pi^+ , & \text{BR} = 1/3 \\ p\pi^0 , & \text{BR} = 2/3 \end{cases}$$

$$\pi^+ \rightarrow \mu^+\nu_\mu \rightarrow \bar{\nu}_\mu e^+ \nu_e \nu_\mu$$

$$\pi^0 \rightarrow \gamma\gamma$$

$$n \text{ (escapes)} \rightarrow p e^- \bar{\nu}_e$$

(Δ^+: ~50% of all $p\gamma$ interactions)



After propagation, with flavor mixing:

$$\nu_e : \nu_\mu : \nu_\tau : p = 1 : 1 : 1 : 1$$

("one ν_μ per cosmic ray")

This neutron model of CR emission is now strongly disfavoured

IceCube, *Nature* **484**, 351 (2012)
M. Ahlers *et al. Astropart. Phys.* **35**, 87 (2011)

What are the ingredients?

To calculate the ν flux from a GRB, we need:

- its gamma-ray luminosity L_γ^{iso} [erg s^{-1}] [measured]
- its variability timescale t_v [s], from the light curve [measured]
- the break energy of its photon spectrum $\epsilon_{\gamma,\text{break}}$ [MeV] [measured]
- the bulk Lorentz factor of its jet Γ [estimated]
- the energy in electrons, magnetic field, protons [estimated]

Now let us cook up the neutrinos ▶

Normalizing neutrinos with observed gamma rays

energy in neutrinos $\propto$ energy in gamma rays

In detail, for each GRB,

$$\int_0^\infty dE_\nu E_\nu F_\nu(E_\nu) = \frac{1}{8} \underbrace{\left[1 - (1 - \langle x_{p\to\pi} \rangle)^{\Delta R/\lambda_{p\gamma}}\right]}_{f_\pi} \frac{1}{f_e} \int_{1\text{ keV}}^{10\text{ MeV}} d\epsilon_\gamma \epsilon_\gamma F_\gamma(\epsilon_\gamma)$$

f_π: fraction of total proton energy transferred to pions
ΔR: size of the emitting region
$\lambda_{p\gamma}$: mean free path for $p\gamma$ interactions
$\langle x_{p\to\pi} \rangle$: avg. fraction of p energy transferred to a π in one interaction
f_e^{-1}: ratio of energy in protons to energy in photons/electrons

$$\frac{\Delta R}{\lambda_{p\gamma}} = \left(\frac{L_\gamma^{\text{iso}}}{10^{52}\,\text{erg s}^{-1}}\right) \left(\frac{0.01}{t_v}\right) \left(\frac{10^{2.5}}{\Gamma}\right)^4 \left(\frac{\text{MeV}}{\varepsilon_{\gamma,\text{break}}}\right)$$

The original recipe: conventional fireball model

Observed photon spectrum [GeV^{-1} cm^{-2}]

$$F_\gamma(\varepsilon_\gamma) \propto \begin{cases} (\varepsilon_\gamma/\varepsilon_{\gamma,\text{break}})^{-\alpha_\gamma} & , \varepsilon_\gamma < \varepsilon_{\gamma,\text{break}} \\ (\varepsilon_\gamma/\varepsilon_{\gamma,\text{break}})^{-\beta_\gamma} & , \varepsilon_\gamma \geq \varepsilon_{\gamma,\text{break}} \end{cases}$$

$$\alpha_\gamma = 1,\ \beta_\gamma = 2.2,\ \varepsilon_{\gamma,\text{break}} = 1 \text{ MeV}$$

+

Assumed proton spectrum

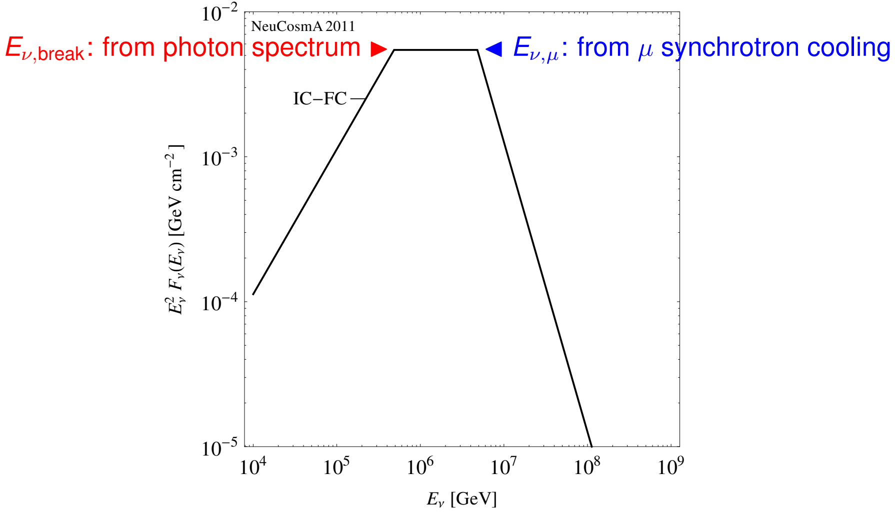
$$N_p(E_p) \propto E_p^{-2}$$

=

Neutrino spectrum, via Δ resonance

$$F_\nu(E_\nu) \propto \begin{cases} \left(\frac{E_\nu}{E_{\nu,\text{break}}}\right)^{-\alpha_\nu} & , E_\nu < E_{\nu,\text{break}} \\ \left(\frac{E_\nu}{E_{\nu,\text{break}}}\right)^{-\beta_\nu} & , E_{\nu,\text{break}} \leq E_\nu < E_{\nu,\mu} \\ \left(\frac{E_\nu}{E_{\nu,\text{break}}}\right)^{-\beta_\nu} \left(\frac{E_\nu}{E_{\nu,\mu}}\right)^{-2} & , E_\nu \geq E_{\nu,\mu} \end{cases}$$

Neutrino spectrum – conventional fireball



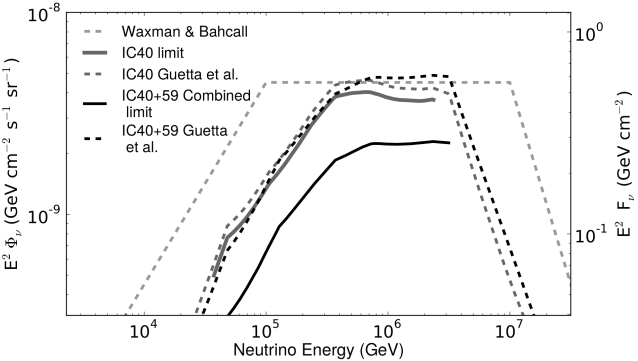
E. WAXMAN, J. N. BAHCALL, *PRL* **78**, 2292 (1997)

D. GUETTA *et al.*, *Astropart. Phys.* **20**, 429 (2004)

Neutron model of UHECR emission under tension?

In 2012, IceCube ruled this analytical version of the fireball model –

- assumed a fixed baryonic loading of 10
- extrapolated diffuse ν flux from 117–215 GRBs ("quasi-diffuse")
- analytical calculation – in tension with upper bounds



IceCube, *Nature* **484**, 351 (2012)
M. Ahlers *et al. Astropart. Phys.* **35**, 87 (2011)
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NeuCosmA: (revised) GRB particle emission – I

Go to the source frame and calculate particle production there:

$$\underbrace{N'_p\left(E'_p\right)}_{\text{proton density at the source [GeV}^{-1}\text{ cm}^{-3}]} \overset{\text{NeuCosmA}}{\otimes} \underbrace{N'_\gamma\left(E'_\gamma\right)}_{\text{photon density at the source}}$$

$$= \underbrace{Q'_\nu\left(E'_\nu\right)}_{\text{ejected neutrino spectrum [GeV}^{-1}\text{ cm}^{-3}\text{ s}^{-1}]}$$

- From Fermi shock acceleration: $N'_p\left(E'_p\right) \propto E'^{-\alpha_p}_p e^{-E'_p/E'_{p,\text{max}}}$
- Photon density at source has same shape as observed:

$$N'_\gamma\left(E'_\gamma\right) = \begin{cases} \left(E'_\gamma/E'_{\gamma,\text{break}}\right)^{-\alpha_\gamma} & , E'_{\gamma,\text{min}} \leq E'_\gamma < E'_{\gamma,\text{break}} \\ \left(E'_\gamma/E'_{\gamma,\text{break}}\right)^{-\beta_\gamma} & , E'_\gamma \geq E'_{\gamma,\text{break}} \end{cases}$$

$$\alpha_\gamma = 1,\ \beta_\gamma = 2.2,\ E'_{\gamma,\text{min}} = 0.2\text{ eV},\ E'_{\gamma,\text{break}} = 1\text{ keV}$$

NeuCosmA: (revised) GRB particle emission – II

Normalize the densities at the source – for one shell collision:

- Photons:

$$\underbrace{\int E'_\gamma \, N'_\gamma \left(E'_\gamma\right) \, dE'_\gamma}_{\text{total energy density in photons}} = \frac{E'^{\text{iso}}_{\gamma-\text{sh}}}{V'_{\text{iso}}}$$

- Protons:

baryonic loading (energy in p's / energy in e's + γ's), *e.g.*, 10

$$\underbrace{\int E'_p \, N'_p \left(E'_p\right) \, dE'_p}_{\text{total energy density in protons}} = \frac{1}{f_e} \frac{E'^{\text{iso}}_{\gamma-\text{sh}}}{V'_{\text{iso}}}$$

We will calculate the ν flux from one collision
– then multiply by the number of collisions: $T_{90}/t_v = 100 - 1000$

NeuCosmA: (revised) GRB particle emission – III

NeuCosmA calculates the injected/ejected spectrum of secondaries (π, K, n, ν, etc.):

$x \equiv E'/E'_p$

$y \equiv E'_p E'_\gamma / \left(m_p c^2\right)$

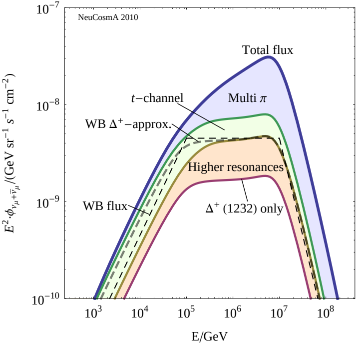
$$Q'\left(E'\right) = \int_{E'}^{\infty} \frac{dE'_p}{E'_p} N'_p\left(E'_p\right) \int_0^{\infty} c \, dE'_\gamma \, N'_\gamma\left(E'_\gamma\right) R\left(x, y\right)$$

response function

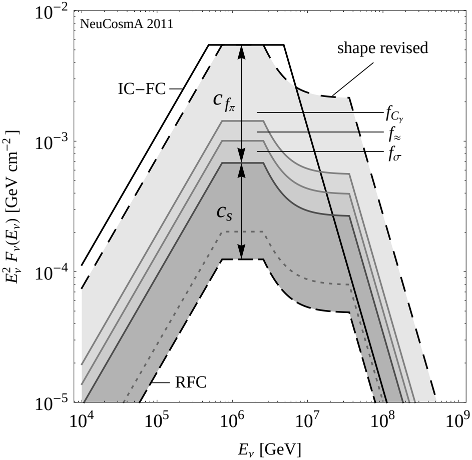
R contains cross sections, multiplicities for different channels

What does NeuCosmA include?

- $p\gamma \to \Delta^+(1232) \to \pi^0, \pi^+, \ldots$
- extra K, n, π^-, multi-π production modes
- synchrotron losses of secondaries
- adiabatic cooling
- full photon spectrum
- neutrino flavor transitions



NeuCosmA – how the neutrino spectrum changes – I

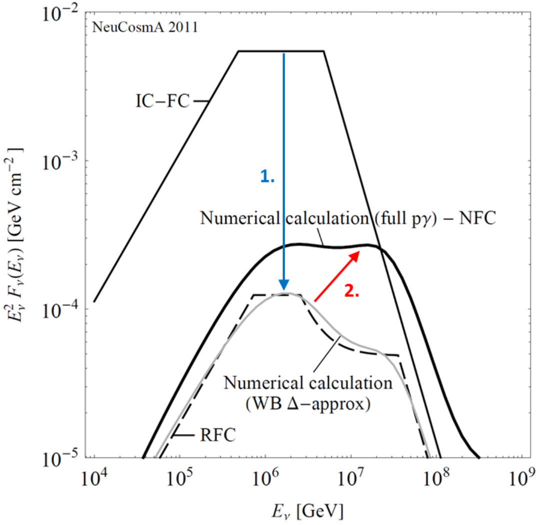


S. Hümmer, P. Baerwald, W. Winter, *Phys. Rev. Lett.* **108**, 231101 (2012)

Corrections to the analytical model:

- shape revised:
  - shift of first break (correction of photohadronic threshold)
  - different cooling breaks for μ's and π's
  - $(1+z)$ correction on the variability scale of the GRB
- Correction cf_π to π prod. efficiency:
  - f_{C_γ}: full spectral shape of photons
  - $f_\approx = 0.69$: rounding error in analytical calculation
  - $f_\sigma \simeq 2/3$: from neglecting the width of the Δ-resonance
- Correction c_S:
  - energy losses of secondaries
  - energy dependence of the mean free path of protons

NeuCosmA – how the neutrino spectrum changes – II



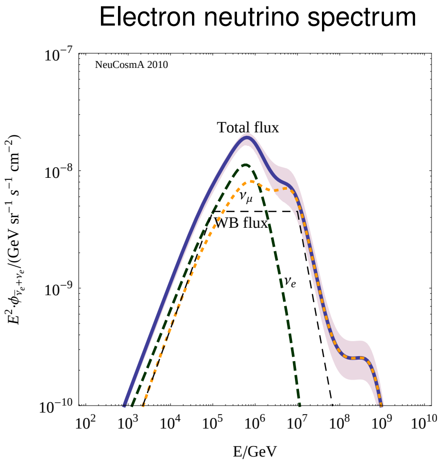
For example, GRB080603A:

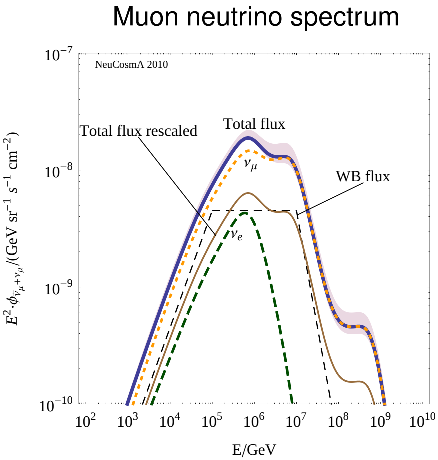
1. Correction to analytical model (IC-FC → RFC)
2. Change due to full numerical calculation

IC-FC: IceCube-Fireball Calculation
RFC: Revised Fireball Calculation
NFC: Numerical Fireball Calculation

S. Hümmer, P. Baerwald, W. Winter, *Phys. Rev. Lett.* **108**, 231101 (2012)

NeuCosmA – neutrino spectra including flavor mixing





P. Baerwald, S. Hümmer, W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

Characteristic double peak structure from μ and π decay in both flavors, additional peak from K^+ decay at 10^8 to 10^9 GeV

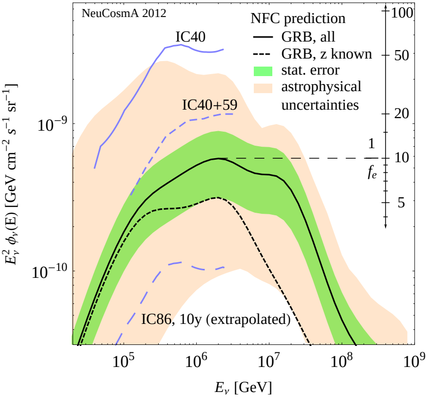
The new prediction of the quasi-diffuse GRB ν flux

Repeat the IceCube GRB neutrino analysis, with NeuCosmA –

- Same GRB sample and parameters as used by IceCube
- Calculate the associated neutrino flux for each burst and the stacked flux $F_\nu(E_\nu)$
- Quasidiffuse flux:

$$\phi_\nu(E_\nu) = F_\nu(E_\nu)\frac{1}{4\pi}\frac{1}{n}\frac{667 \text{ bursts}}{\text{yr}}$$

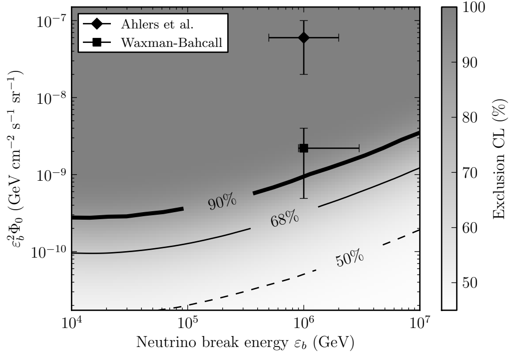
ν flux $\sim$ 1 order of magnitude lower!



S. Hümmer, P. Baerwald, W. Winter, *PRL* **108**, 231101 (2012)

Improved IceCube GRB bounds (2014)

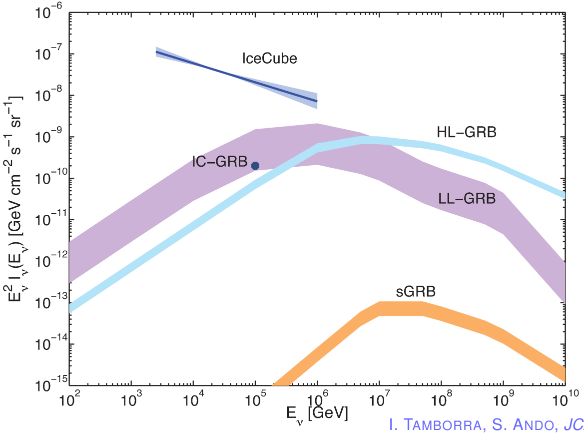
- Only upgoing ν_μ's with > 1 TeV used
- Four years of data (IC-40, -59, -79, -86)
- Larger GRB catalogue (506 bursts)
- One coincident event found, with low statistical significance
- $\lesssim 1\%$ of the diffuse flux can be from prompt GRB emission



ICECUBE, *ApJ* **805**, L5 (2015)

Contribution of GRBs to the diffuse ν flux

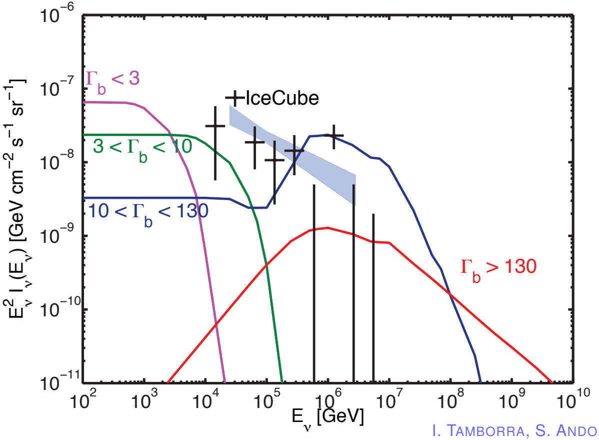
- Three populations: high-luminosity long GRBs (HL-GRB), low-luminosity long GRBs (LL-GRB), short GRBs (sGRB)
- Sub-PeV: GRBs contribute a few % to the IceCube diffuse flux
- PeV: contribution could be higher



I. Tamborra, S. Ando, *JCAP* **1509**, 036 (2015)

What about low-luminosity and choked GRBs?

- Low-luminosity and choked GRBs might be in the same family as high-luminosity long GRBs
- Due to lower jet speeds (Γ_b), they do not break out
- They might explain the TeV region of the IceCube diffuse ν flux:



I. Tamborra, S. Ando, 1512.01559

Going beyond the neutron model

The neutron model hinges on:

1. *p*'s magnetically confined, only *n*'s escape
2. *p*'s interact at most once, *n*'s do not (*optically thin source*)

However, under the "one ν_μ per CR" hypothesis, GRBs are disfavored as the sole sources of UHECRs (AHLERS *et al.*).

M. AHLERS, M. GONZÁLEZ-GARCÍA, F. HALZEN *Astropart. Phys.* **35**, 87 (2011)

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1. *p*'s magnetically confined, only *n*'s escape
2. *p*'s interact at most once, *n*'s do not (*optically thin source*)

However, under the "one ν_μ per CR" hypothesis, GRBs are disfavored as the sole sources of UHECRs (AHLERS *et al.*).

What if 1 and 2 are violated?

- *p*'s "leak out", not accompanied by (direct) ν production
- multiple *p* interactions enhance the ν flux
- in *optically thick sources*, only *n*'s at the borders escape

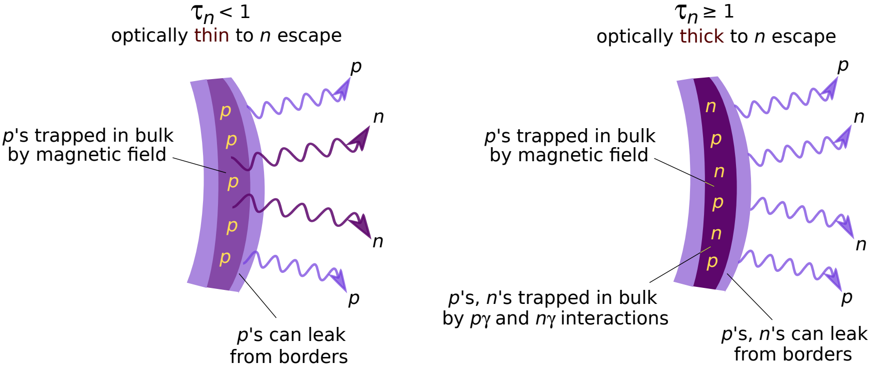
M. AHLERS, M. GONZÁLEZ-GARCÍA, F. HALZEN *Astropart. Phys.* **35**, 87 (2011)

Going beyond the neutron model

We have improved the model – now UHECRs escape as either:

- neutrons, which decay into protons outside the source; or
- protons that leak out without interacting inside the source

Relative contributions determined by $\tau_n \equiv \left(t_{p\gamma}^{-1} / t_{\text{dyn}}^{-1} \right) \Big|_{E'_{p,\max}}$



A two-component model of UHECR emission

Two important points:

1. $E'_{p,\mathrm{max}}$ is determined by energy-loss processes:
   $t'_{\mathrm{acc}}\left(E'_{p,\mathrm{max}}\right) = \min\left[t'_{\mathrm{dyn}}, t'_{\mathrm{syn}}\left(E'_{p,\mathrm{max}}\right), t'_{p\gamma}\left(E'_{p,\mathrm{max}}\right)\right]$

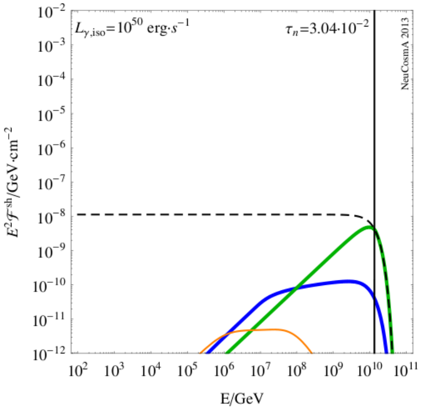
2. Photons can be trapped in the source by Thomson scattering:

$$e + \gamma \to e + \gamma$$

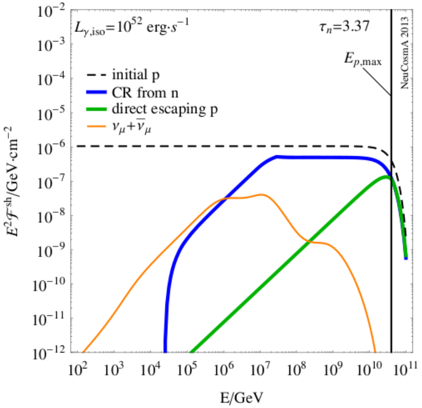
Photosphere: radius where $\tau_{e\gamma}\left(E'_\gamma\right) = 1$ for all E'_γ

Sample neutrino fluences –

Optically **thin** source

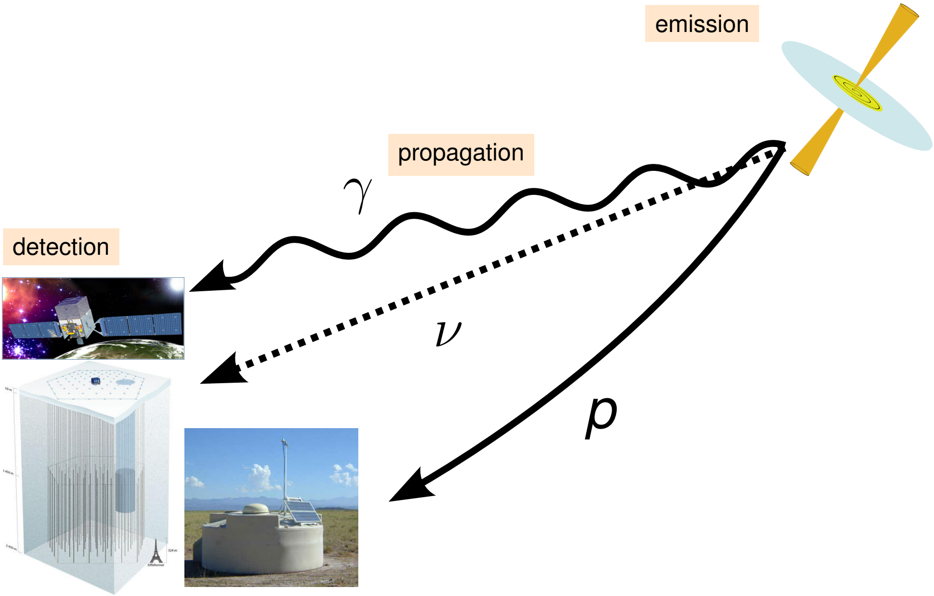


Optically **thick** source



P. Baerwald, MB, W. Winter, *ApJ* **768**, 186 (2013)

From the sources to us



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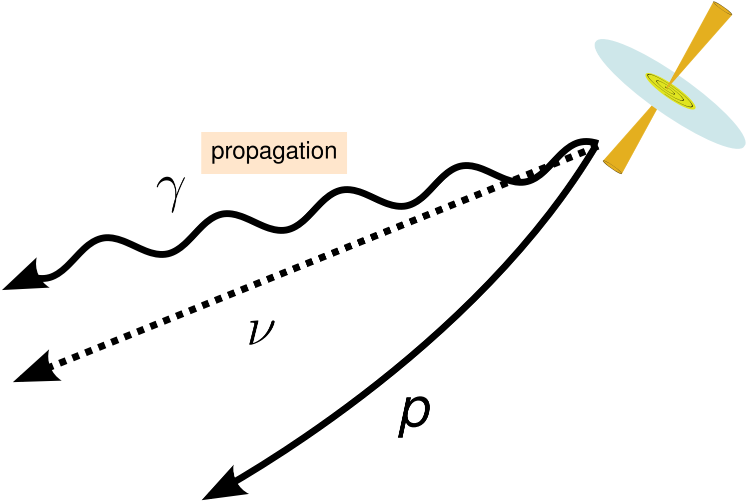
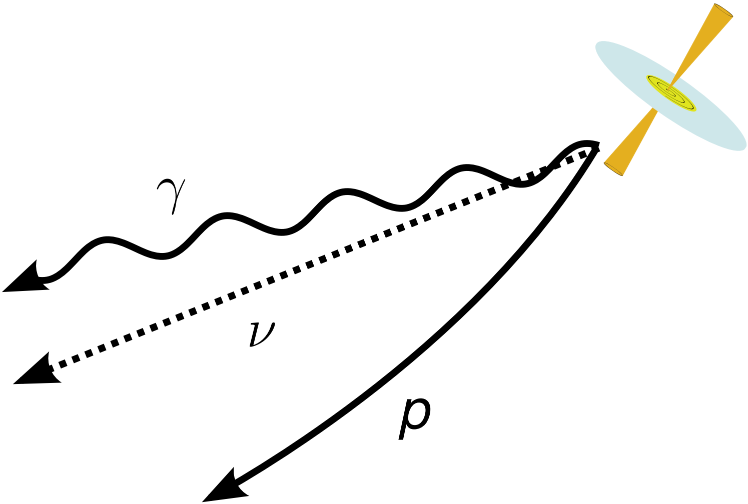


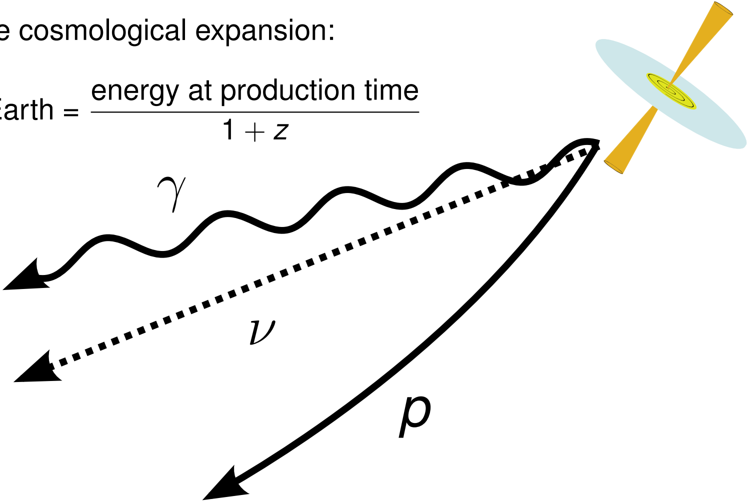
From the sources to us



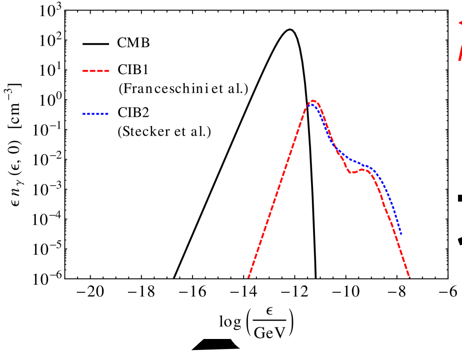
From the sources to us

Because of the cosmological expansion:

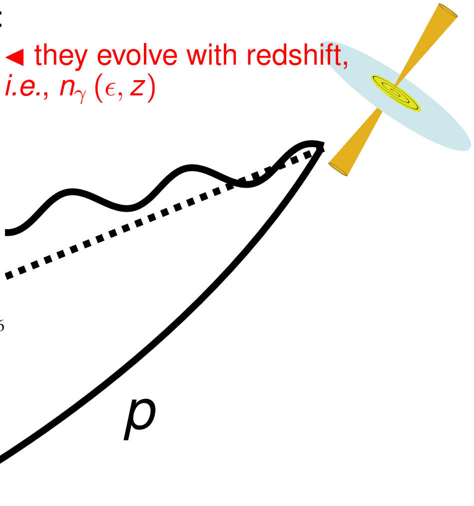
$$\text{energy at Earth} = \frac{\text{energy at production time}}{1+z}$$



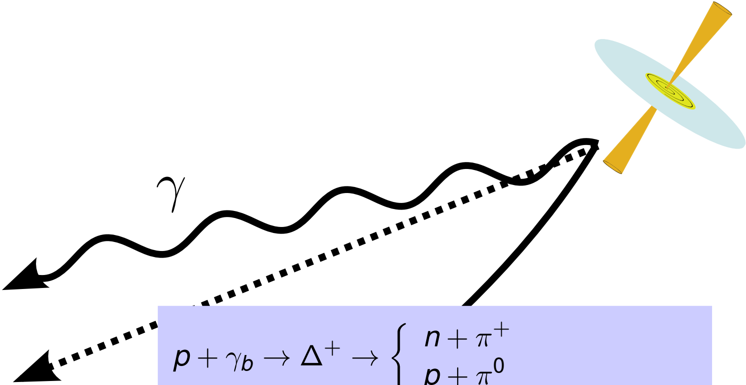
Cosmological photon backgrounds:



◀ they evolve with redshift, *i.e.*, $n_\gamma(\epsilon, z)$



From the sources to us



$$p + \gamma_b \rightarrow \Delta^+ \rightarrow \begin{cases} n + \pi^+ \\ p + \pi^0 \end{cases}$$

$$\pi^0 \rightarrow \gamma + \gamma$$

Cosmogenic ν's: $\pi^+ \rightarrow e^+ + \nu_e + \bar{\nu}_\mu + \nu_\mu$

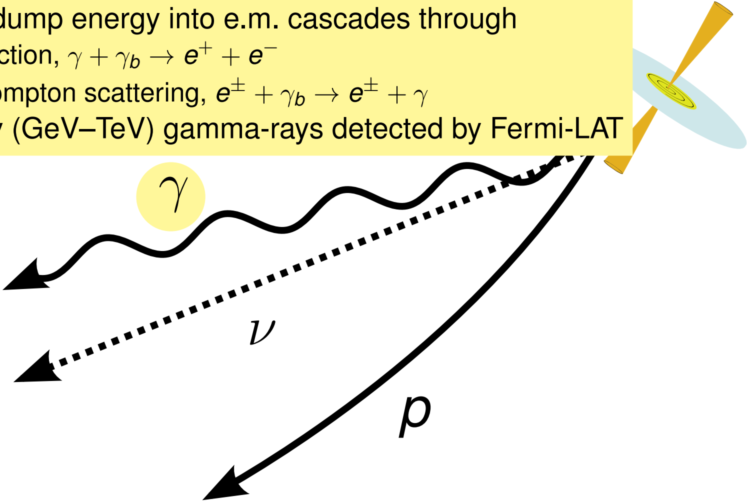
$n \rightarrow p + e^- + \bar{\nu}_e$

From the sources to us

γ's and $e^{\pm}$'s dump energy into e.m. cascades through

- pair production, $\gamma + \gamma_b \to e^+ + e^-$
- inverse Compton scattering, $e^{\pm} + \gamma_b \to e^{\pm} + \gamma$

Lower-energy (GeV–TeV) gamma-rays detected by Fermi-LAT



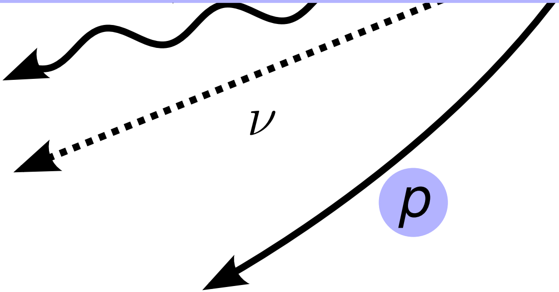
From the sources to us

p's are deflected by extragalactic magnetic fields
⇒ except for the most energetic ones, they are **not** expected to point back to the sources — Pierre Auger found weak correlation with known AGN positions

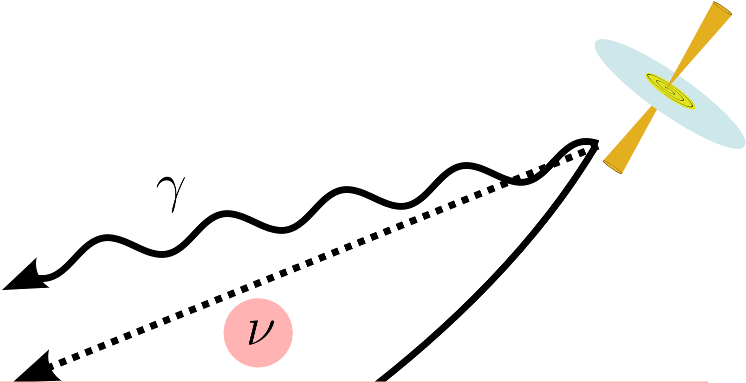
They lose energy through:

- pair production, $p + \gamma_b \rightarrow p + e^+ + e^-$
- photohadronic interactions, $p\gamma_b$

— depend on the redshift evolution of the cosmological γ backgrounds



From the sources to us



Initial UHE ν flavor fluxes: $\nu_e : \nu_\mu : \nu_\tau = 1 : 2 : 0$

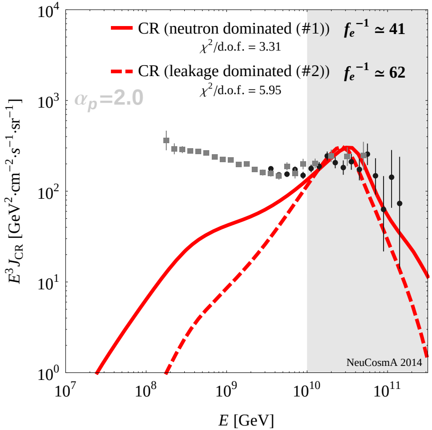
Probability of $\nu_\alpha \to \nu_\beta$ transition: $P_{\alpha\beta}(E_0, z)$

Flavor oscillations redistribute the fluxes
– at Earth: $\nu_e : \nu_\mu : \nu_\tau \approx 1 : 1 : 1$ (might be changed by exotic physics!)

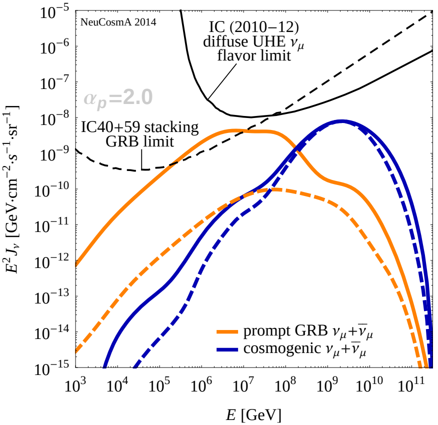
What do the diffuse fluxes look at Earth, then?

Neutron model vs. two-component model:
prompt and cosmogenic ν's

UHECRs



Neutrinos



P. Baerwald, MB, W. Winter, *ApJ* **768**, 186 (2013)
P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)
See also: H. He, K. Murase, *et al.*, *ApJ* **752**, 29 (2012)

Using UHECR + neutrino constraints

We can already limit the parameter space by using the UHECR observations and ν upper bounds:

P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)

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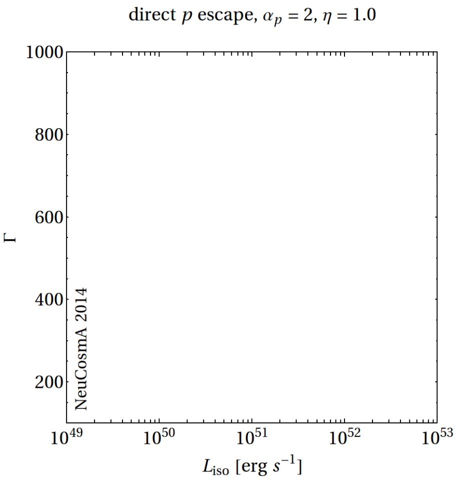

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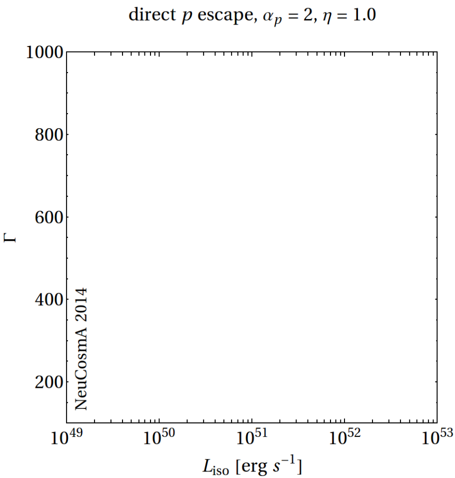

P. BAERWALD, MB, W. WINTER, *Astropart. Phys.* **62**, 66 (2015)

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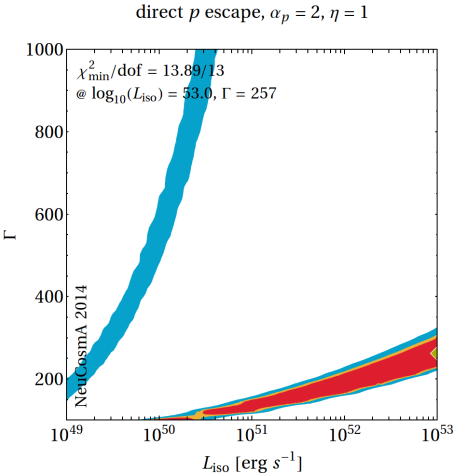

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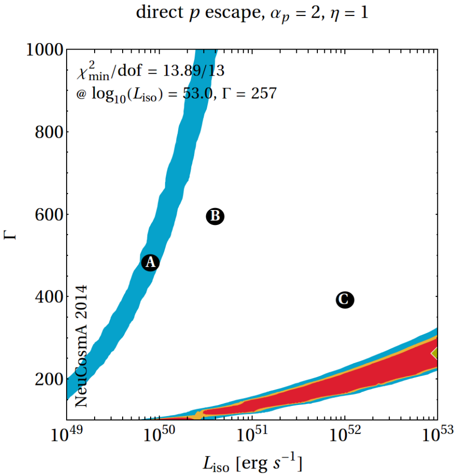

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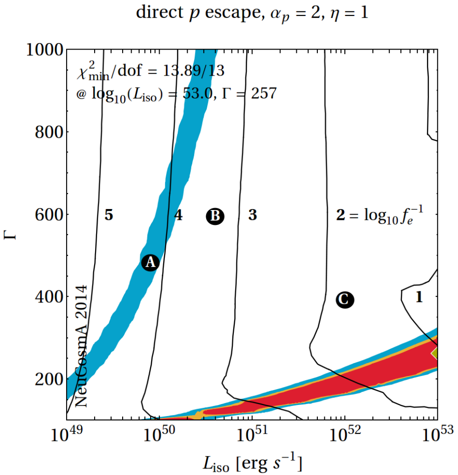

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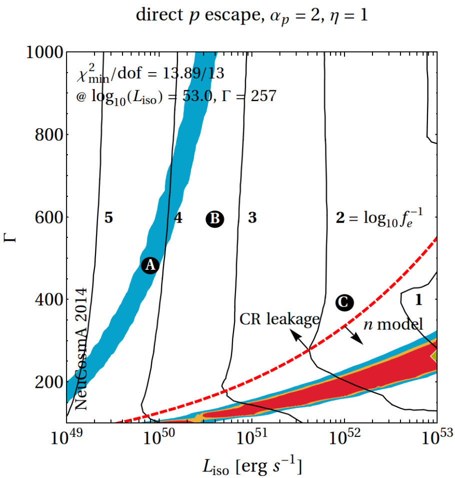
5. Identify the region corresponding to pure *n* escape and to *n* escape + CR leakage



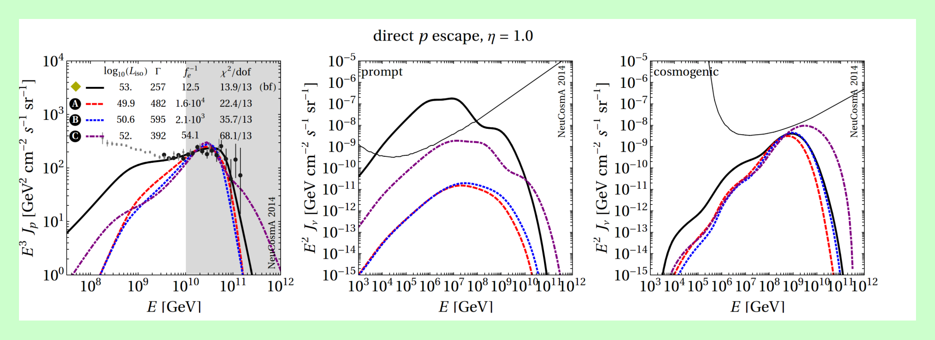
P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)

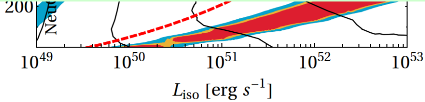
Using UHECR + neutrino constraints

We can already limit the parameter space by using the UHECR observations and ν upper bounds:

1. Generate the UHECR spectrum at every

direct p escape, $\alpha_p = 2$, $\eta = 1$





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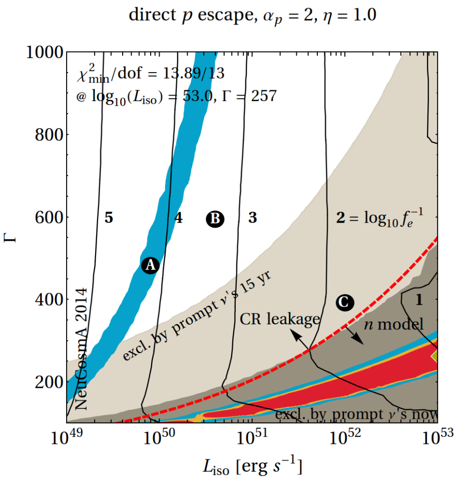
6. Find the region where the number of prompt ν_μ's is $>$ 2.44, *i.e.*, the excluded region at 90% C.L.



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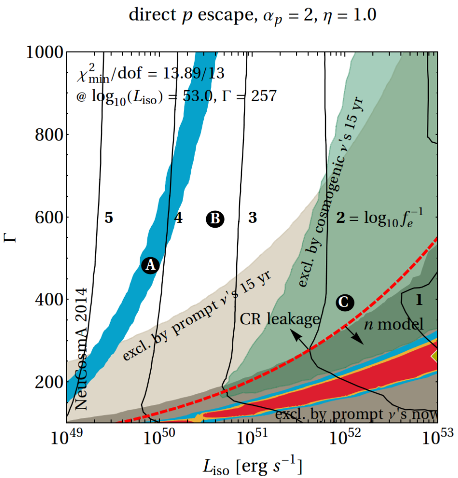
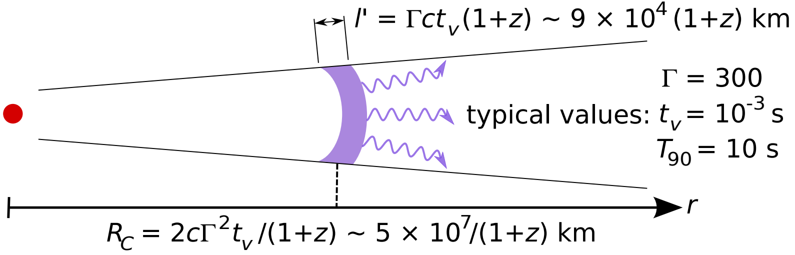
5. Identify the region corresponding to pure *n* escape and to *n* escape + CR leakage
6. Find the region where the number of prompt ν_μ's is > 2.44, *i.e.*, the excluded region at 90% C.L.
7. After 15 yr of exposure and no detection, cosmogenic neutrinos also exclude



P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)

The *static* burst

So far, we have talen *all* internal collisions to occur at the same radius:



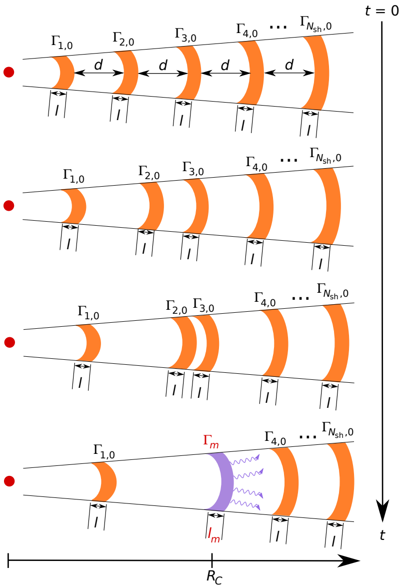
- *average* speed Γ inferred from afterglow observations
- "variability timescale" t_v measured from the light curve
- redshift z measured for the host galaxy

Static burst: made up of $T_{90}/t_v \sim 100 - 1000$ identical collisions

Going further: an evolving burst

Consider an evolving fireball –

- the fireball expands with time
- $\sim$ 1000 shells propagate with different speeds
- they have different masses
- they collide at different radii (collisions no longer identical)



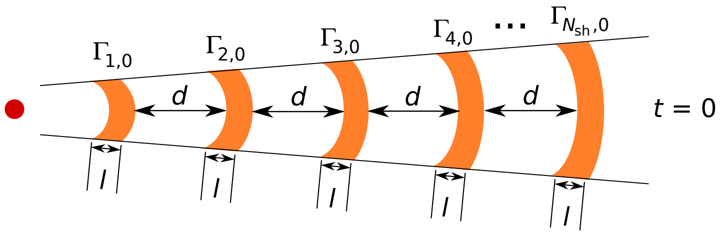
Why does this matter?

The particle (γ, p) densities fall as the fireball expands – particle production conditions change with time/radius

S. Kobayashi, T. Piran, R. Sari, *ApJ* **490**, 92 (1997)
F. Daigne, R. Mochkovitch, *MNRAS* **296**, 275 (1998)
D. Guetta, M. Spada, E. Waxman, *ApJ* **557**, 399 (2001)
N. Globus, *et al.*, *MNRAS* **451**, 751 (2015)

Initialising the burst simulation

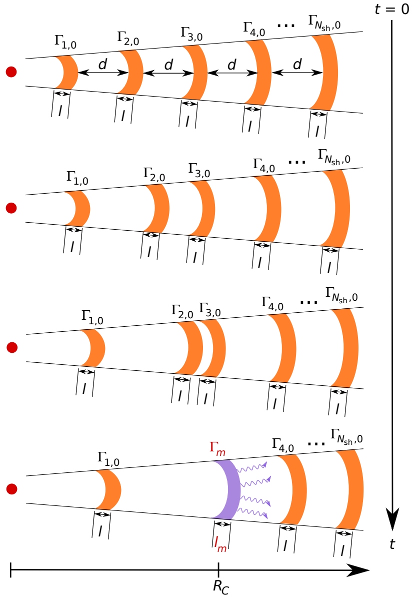
Initial number of plasma shells in the jet: $\gtrsim 1000$



Initial values of shell parameters:

- Width of shells and separation between them: $l = d = c \cdot \delta t_{eng}$
- Equal kinetic energy for all shells ($\sim 10^{52}$ erg)
- Shell speeds $\Gamma_{k,0}$ follow a distribution (log-normal or other)

Propagating and colliding the shells



During propagation:

- speeds, masses, widths do not change (only in collisions)
- the new, merged shells continue propagating and can collide again

Evolution stops when either:

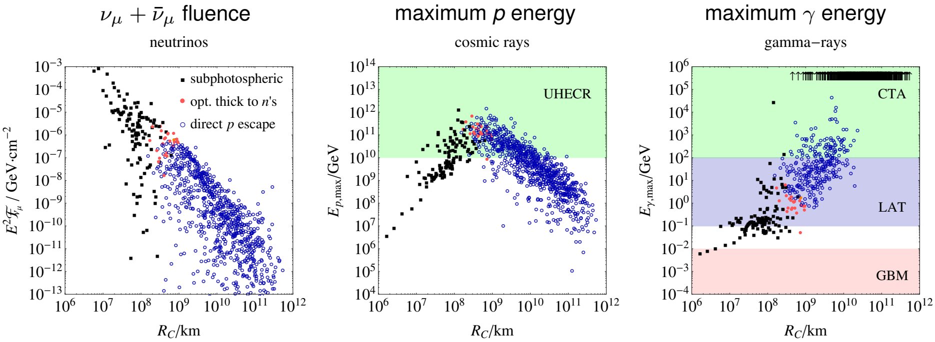
- a single shell is left; or
- all remaining shells have reached the circumburst medium ($\gtrsim 6 \times 10^{11}$ km)

final number of collisions
$\approx$
number of initial shells ($\gtrsim 1000$)

S. Kobayashi, T. Piran, R. Sari, *ApJ* **490**, 92 (1997)
F. Daigne, R. Mochkovitch, *MNRAS* **296**, 275 (1998)

Tracking each collision individually

Each collision occurs in a different emission regime –

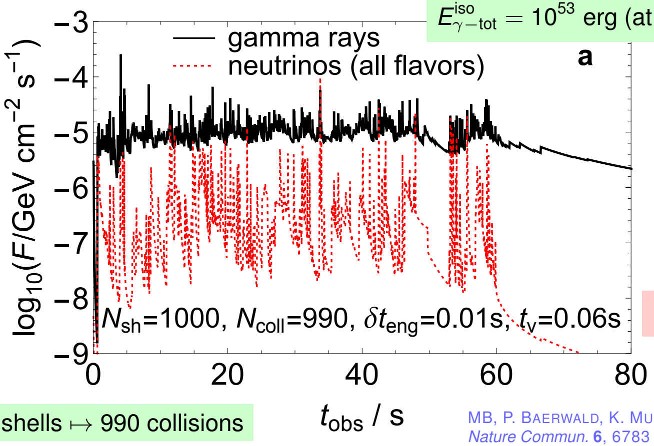


(observer's frame) (source frame)

MB, P. Baerwald, K. Murase, W. Winter, *Nat. Commun.* **6**, 6783 (2015)

An emission pulse is assigned to each collision
– their superposition yields a synthetic light curve:

Total energy in gamma-rays:
$E^{\text{iso}}_{\gamma-\text{tot}} = 10^{53}$ erg (at source)



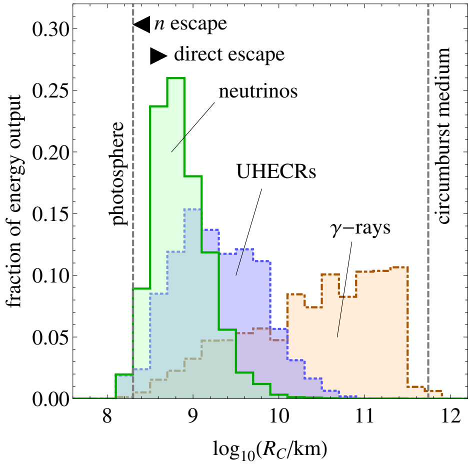
$T_{90} \approx 59$ s

1000 initial shells $\mapsto$ 990 collisions

MB, P. Baerwald, K. Murase, W. Winter
Nature Commun. **6**, 6783 (2015)

Different particles come from different jet regions

Emission of different species peaks at different collision radii –



Why?

As the fireball expands, photon and proton densities fall

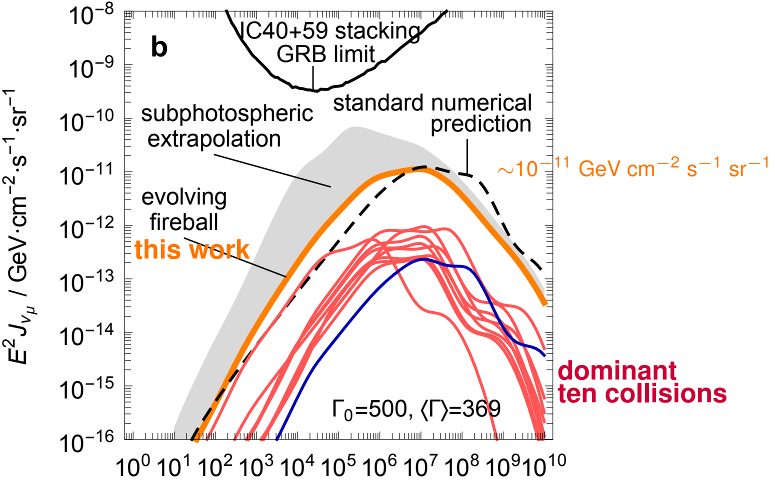
Why does it matter?

GRB parameters derived from gamma-ray observations might not be adequate to describe ν and UHECR emission

MB, P. Baerwald, K. Murase, W. Winter, *Nat. Commun.* **6**, 6783 (2015)

A robust minimal diffuse ν flux from GRBs

- Take the simulated burst as stereotypical
- Quasi-diffuse neutrino flux, assuming 667 identical GRBs per year:



MB, P. Baerwald, K. Murase, W. Winter
Nature Commun. **6**, 6783 (2015)

How is the new prediction different?

- The top-contributing collisions are at the photosphere
- Pion production efficiency there is independent of Γ:

$$f_{p\gamma}^{\text{ph}} \sim 5 \cdot \frac{\varepsilon}{0.25} \cdot \frac{\epsilon_e}{0.1} \cdot \frac{1\ \text{keV}}{\epsilon'_{\gamma,\text{break}}}$$

ε: energy dissipation efficiency
ϵ_e: fraction of dissipated energy as e.m. output (photons)

- $\Rightarrow$ Time-integrated neutrino fluence dominated is independent of Γ:

$$\mathcal{F}_\nu \propto \frac{N_{\text{coll}}\left(f_{p\gamma} \gtrsim 1\right)}{N_{\text{coll}}^{\text{tot}}} \times \min\left[1, f_{p\gamma}^{\text{ph}}\right] \times \frac{\epsilon_p}{\epsilon_e} \times E_{\gamma-\text{tot}}^{\text{iso}}$$

- Compare to standard predictions, which have a $\langle\Gamma\rangle^{-4}$ dependence
- Raising ϵ_p automatically decreases ϵ_e, so the photosphere grows, but still ~ 10 photospheric collisions dominate

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  ε: energy dissipation efficiency
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- $\Rightarrow$ Time-integrated neutrino fluence dominated is independent of Γ:

$$\mathcal{F}_\nu \propto \frac{\overbrace{N_{\text{coll}}\left(f_{p\gamma} \gtrsim 1\right)}^{\sim 10}}{\underbrace{N_{\text{coll}}^{\text{tot}}}_{\sim 1000}} \times \min\left[1, f_{p\gamma}^{\text{ph}}\right] \times \overbrace{\frac{\epsilon_p}{\epsilon_e}}^{10} \times \overbrace{E_{\gamma-\text{tot}}^{\text{iso}}}^{10^{53} \text{ erg}}$$

- Compare to standard predictions, which have a $\langle\Gamma\rangle^{-4}$ dependence
- Raising ϵ_p automatically decreases ϵ_e, so the photosphere grows, but still ~ 10 photospheric collisions dominate

What makes a GRB bright in neutrinos? (Preliminary)

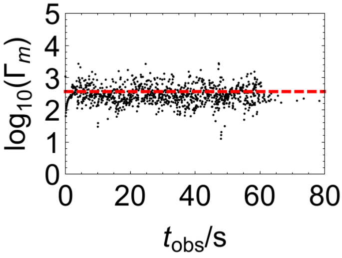
- Goal: to use the morphology of the GRB gamma-ray light curve alone to assess whether the burst is neutrino-bright
- The central engine determines the features of the light curve

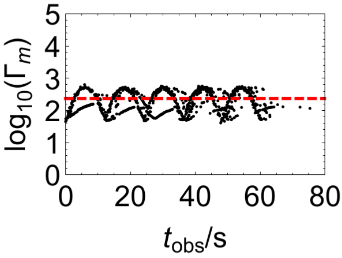
GRB A: undisciplined engine

- Engine emits shells with log-normal distribution of Γ
- Broad Γ distribution ($A_\Gamma = 1$)

GRB B: disciplined engine

- Engine emits shells with oscillating Γ
- Narrow Γ distribution ($A_\Gamma = 0.1$)





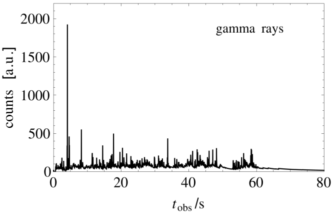
How do the gamma-ray and neutrino light curves look?

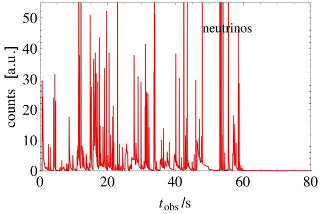
GRB A: undisciplined engine

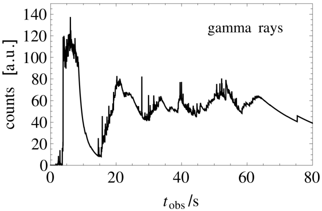
- Fast variability dominates
- No broad pulses

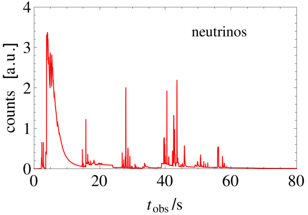
GRB B: disciplined engine

- Broad pulses dominate
- Fast variability on top





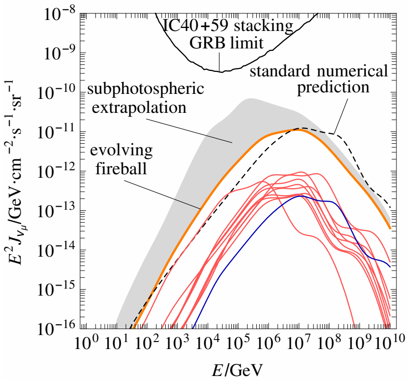


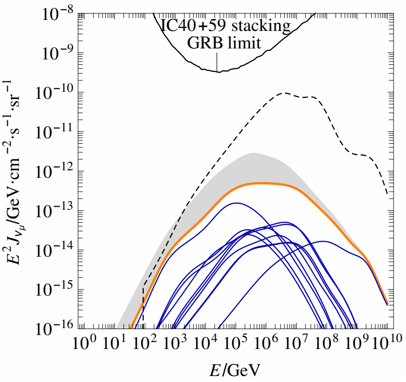


So which burst is neutrino-bright?

Compare the quasi-diffuse neutrino fluxes:

GRB A: undisciplined engine	GRB B: disciplined engine
$\sim 10^{-11}$ GeV cm^{-1} s^{-1} sr^{-1}	$\sim 5 \cdot 10^{-13}$ GeV cm^{-1} s^{-1} sr^{-1}





∴ An undisciplined engine makes a GRB neutrino-bright

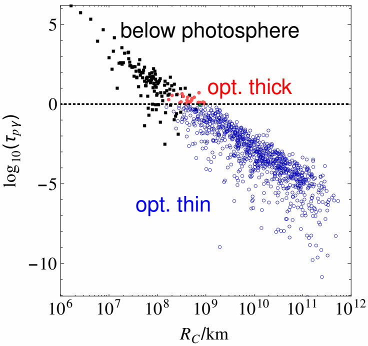
Why?

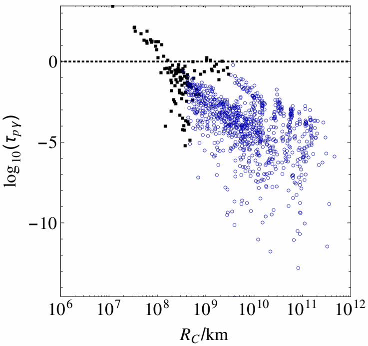
GRB A: undisciplined engine

- Shells have very different speeds
- Collide quickly, close to engine
- High p and γ densities
- $\sim$ 10 optically-thick bursts near the photosphere

GRB B: disciplined engine

- Shells have similar speeds
- Collide far from engine
- Low p and γ densities
- All (superphotospheric) collisions are optically thin





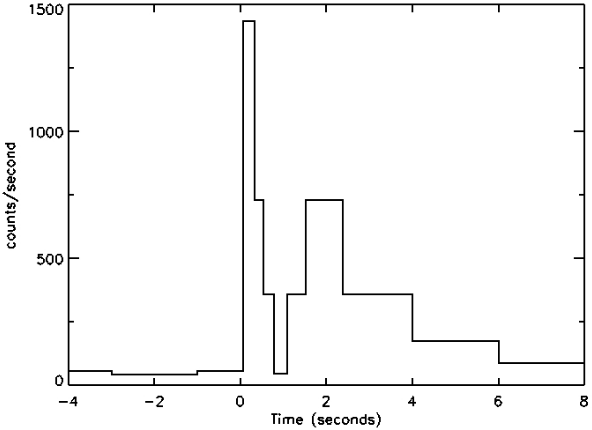
Conclusions . . . and the future

- GRBs *are* good UHECR and ν source candidates
- *But* the CR-ν-γ connection is trickier than originally thought
- Contribute to the diffuse HE astrophysical ν flux at the few % level
- Likely to be the first point neutrino sources to be resolved
- Need next-gen neutrino telescopes (IceCube-Gen2, KM3NeT) . . .
- . . . while Auger, Telescope Array, *etc.* gather extensive UHECR statistics
- By exploiting gamma-ray light curve morphology, we may be able to improve selection cuts

Backup slides

GRBs – the first one detected

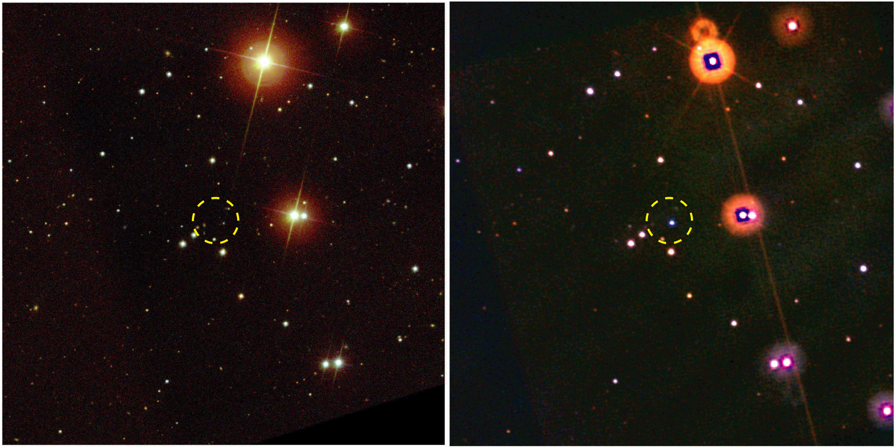
First GRB detected: July 2, 1967, 14:19 UTC



Detected by *Vela* 3, 4a, 4b (found on archival data)

What do they *look* like?

e.g., GRB060218 seen by *Swift*



SDSS, Swift Collab., Sloan Foundation, NSF, NASA

A two-component model of CR emission

Optical depth:

$$\tau_n = \left.\frac{t_{p\gamma}^{-1}}{t_{\text{dyn}}^{-1}}\right|_{E_{p,\max}} = \begin{cases} \lesssim 1\,, & \text{optically **thin** source} \\ > 1\,, & \text{optically **thick** source} \end{cases}$$

Particles can escape from within a shell of thickness λ'_{mfp}:

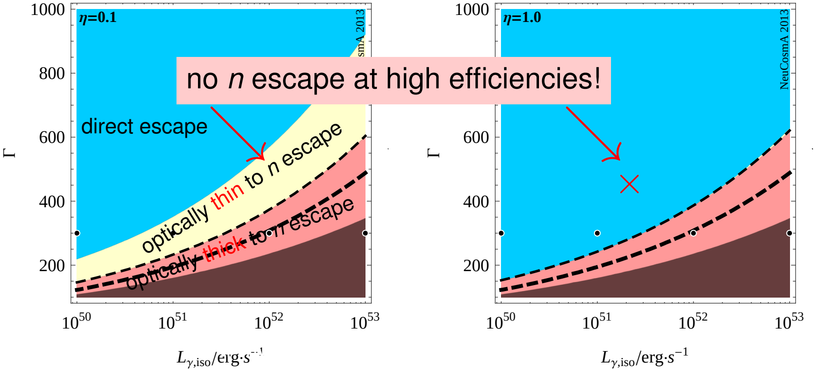
$$\left.\begin{aligned} \lambda'_{p,\text{mfp}}\left(E'\right) &= \min\left[\Delta r', R'_L\left(E'\right), ct'_{p\gamma}\left(E'\right)\right] \\ \lambda'_{n,\text{mfp}}\left(E'\right) &= \min\left[\Delta r', ct'_{p\gamma}\left(E'\right)\right] \end{aligned}\right\} f_{\text{esc}} = \frac{\lambda'_{\text{mfp}}}{\Delta r'}$$

fraction of escaping particles

We *need* direct proton escape

Scan of the GRB emission parameter space –

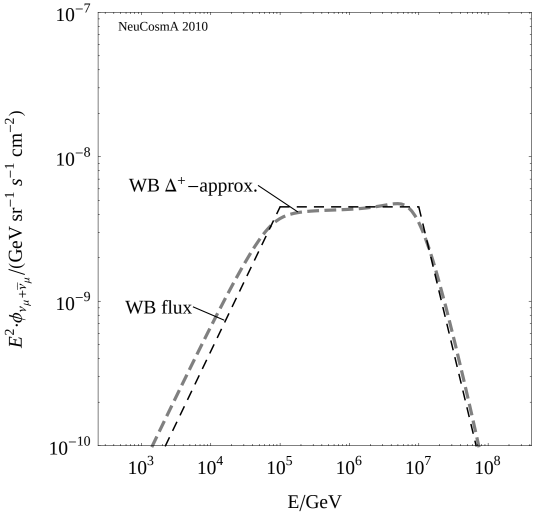
acceleration efficiency ⟶ $\eta = 0.1$ $\eta = 1.0$



P. Baerwald, MB, and W. Winter, *ApJ* **768**, 186 (2013)

we need high efficiencies ⇒ direct proton escape *is* required

NeuCosmA – the full photohadronic cross section

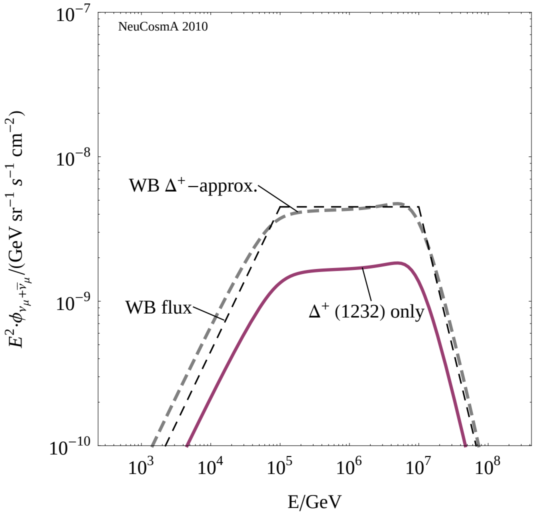


NeuCosmA – the full photohadronic cross section

Contributions to $(\nu_\mu + \bar{\nu}_\mu)$ flux from $\pi^\pm$ decay divided in:

- Δ(1232)-resonance

P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

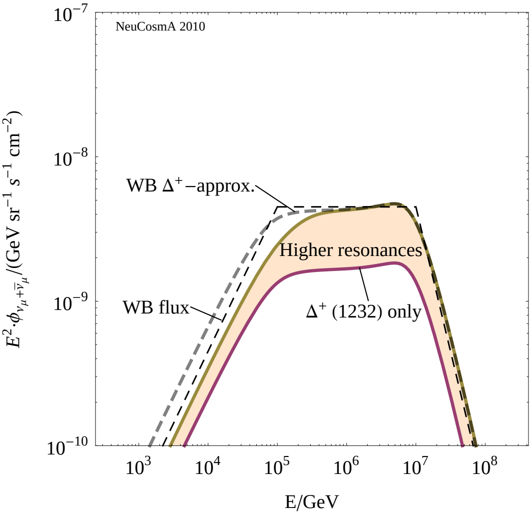


NeuCosmA – the full photohadronic cross section

Contributions to $(\nu_\mu + \bar{\nu}_\mu)$ flux from $\pi^\pm$ decay divided in:

- $\Delta(1232)$-resonance
- Higher resonances

P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)



NeuCosmA – the full photohadronic cross section

Contributions to $(\nu_\mu + \bar{\nu}_\mu)$ flux from $\pi^\pm$ decay divided in:

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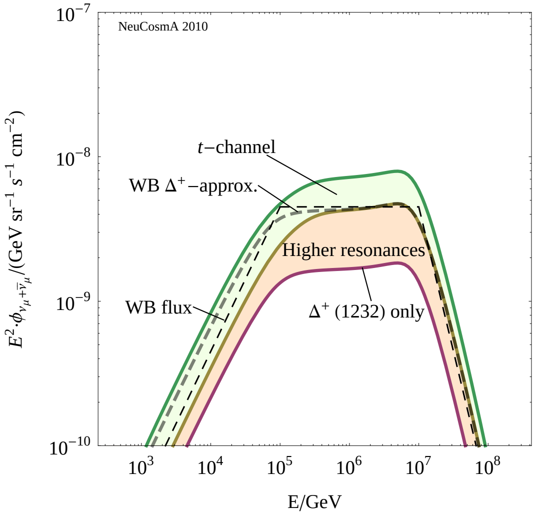
P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

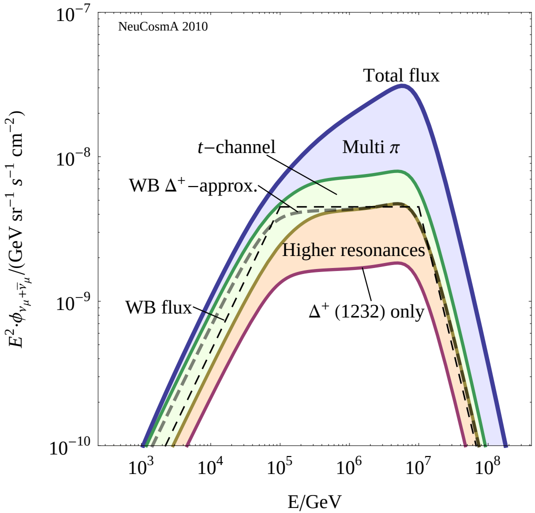


NeuCosmA – the full photohadronic cross section

Contributions to $(\nu_\mu + \bar{\nu}_\mu)$ flux from $\pi^\pm$ decay divided in:

- Δ(1232)-resonance
- Higher resonances
- *t*-channel (direct production)
- High energy processes (multiple π)

P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)



Especially "Multi π" contribution leads to change of flux shape; neutrino flux higher by up to a factor of 3 compared to WB treatment

NeuCosmA – further particle decays

$$\pi^+ \rightarrow \mu^+ + \nu_\mu$$
$$\mu^+ \rightarrow e^+ + \nu_e + \bar{\nu}_\mu$$

$$\pi^- \rightarrow \mu^- + \bar{\nu}_\mu$$
$$\mu^- \rightarrow e^- + \bar{\nu}_e + \nu_\mu$$

$$K^+ \rightarrow \mu^+ + \nu_\mu$$

$$n \rightarrow p + e^- + \bar{\nu}_e$$

NeuCosmA – further particle decays

$$\pi^+ \rightarrow \mu^+ + \nu_\mu$$
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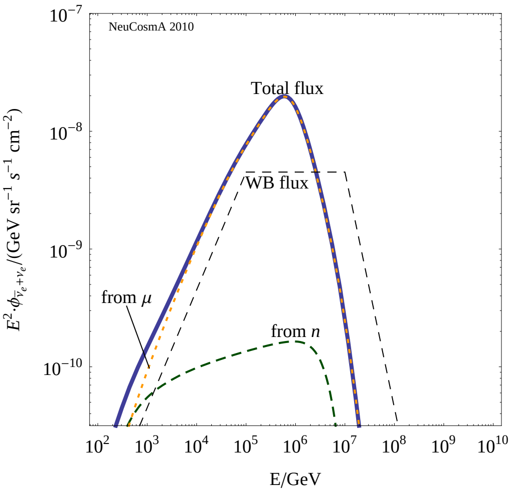
$$\mu^- \rightarrow e^- + \bar{\nu}_e + \nu_\mu$$

$$K^+ \rightarrow \mu^+ + \nu_\mu$$

$$n \rightarrow p + e^- + \bar{\nu}_e$$

Resulting ν_e flux (at the observer)



P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

NeuCosmA – further particle decays

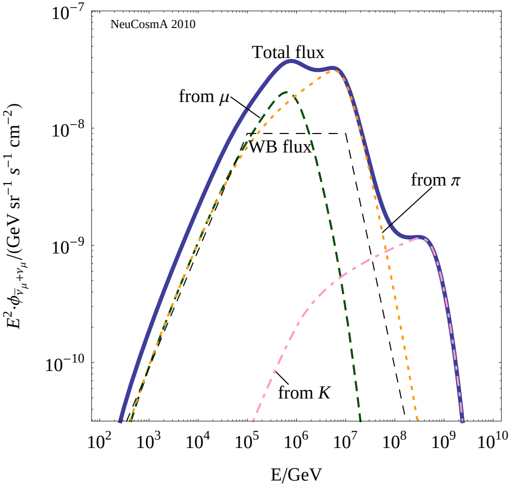
$$\pi^+ \rightarrow \mu^+ + \nu_\mu$$
$$\mu^+ \rightarrow e^+ + \nu_e + \bar{\nu}_\mu$$

$$\pi^- \rightarrow \mu^- + \bar{\nu}_\mu\,,$$
$$\mu^- \rightarrow e^- + \bar{\nu}_e + \nu_\mu$$

$$K^+ \rightarrow \mu^+ + \nu_\mu$$

$$n \rightarrow p + e^- + \bar{\nu}_e$$

Resulting ν_μ flux (at the observer)



P. Baerwald, S. Hümmer, and W. Winter, *Phys. Rev.* **D83**, 067303 (2011)

Propagating the UHECRs to Earth

We use a Boltzmann equation to transport protons to Earth:

- Comoving number density of protons ($\mathrm{GeV}^{-1}\ \mathrm{cm}^{-3}$):

$$Y_p(E,z) = n_p(E,z) / (1+z)^3 \ ,$$

with n_p the real number density

- Transport equation (comoving source frame):

$$\dot{Y}_p = \partial_E (HEY_p) + \partial_E (b_{e^+e^-} Y_p) + \partial_E (b_{p\gamma} Y_p) + \mathcal{L}_{\mathrm{CR}}$$

adiabatic losses

pair production losses

photohadronic losses

CR injection from sources

$$Q_{\mathrm{CR}}(E) \propto E^{-\alpha_p} e^{-E/E_{p,\max}}$$

The need for km-scale neutrino telescopes

Expected ν flux from cosmological accelerators (Waxman & Bahcall 1997 & 1998):

$$E^2\Phi_\nu \sim 10^{-8} \frac{f_\pi}{0.2} \left(\frac{\dot{\varepsilon}_{\text{CR}}^{[10^{10},10^{12}]}}{10^{44}\ \text{erg Mpc}^{-3}\ \text{yr}^{-1}} \right)\ \text{GeV cm}^{-2}\ \text{s}^{-1}\ \text{sr}^{-1}$$

Integrated flux above 1 PeV:

$$\Phi_\nu(> 1\ \text{PeV}) \sim \int_{1\ \text{PeV}}^{\infty} \frac{10^{-8}}{E^2}\, dE \sim 10^{-20}\ \text{cm}^{-2}\ \text{s}^{-1}\ \text{sr}^{-1}$$

Number of events from half of the sky (2π):

$$N_\nu \simeq 2\pi \cdot \Phi_\nu(> 1\ \text{PeV}) \cdot 1\ \text{yr} \cdot A_{\text{eff}} \approx \left(2.4 \times 10^{-10}\ \text{cm}^{-2}\right) A_{\text{eff}}\ ,$$

where A_{eff} is the effective area of the detector

To detect $N_\nu > 1$ events per year, we need an area of

$$A_{\text{eff}} \gtrsim 0.4\ \text{km}^2$$

Therefore, we need km-scale detectors, like IceCube

Cosmogenic neutrinos

We have seen that protons interact with the cosmological photon fields (CMB, etc.), *e.g.*,

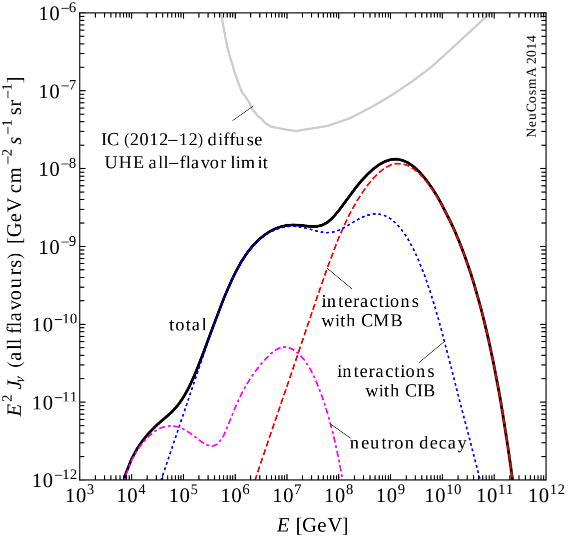
$$p + \gamma \to \Delta^+ \to \pi^+ + n \, ,$$

and neutrinos are created in the decays of the secondaries:

$$\pi^+ \to \mu^+ + \nu_\mu$$
$$\mu^+ \to \bar{\nu}_\mu + \nu_e + e^+$$
$$n \to p + e^- + \bar{\nu}_e$$

These are called *cosmogenic neutrinos*

Cosmogenic neutrinos



P. Baerwald, MB, W. Winter, *Astropart. Phys.* **62**, 66 (2015)

Interaction with the photon backgrounds – I

- Energy loss rate (GeV s^{-1}):

$$b(E) \equiv \frac{dE}{dt}$$

- For pair production $p\gamma \longrightarrow pe^+e^-$:

$$b_{e^+e^-}(E, z) = -\alpha r_0^2 \left(m_e c^2\right)^2 c \int_2^\infty d\xi n_\gamma \left(\frac{\xi m_e c^2}{2\gamma}, z\right) \frac{\phi(\xi)}{\xi^2}$$

  - n_γ: isotropic photon background (GeV^{-1} cm^{-3})
  - ξ: photon energy in units of $m_e c^2$
  - proton energy: $E = \gamma m_p c^2$ ($\gamma \gg 1$)
  - $\phi(\xi)$: (tabulated) integral in energy of outgoing e^-

G. BLUMENTHAL, *Phys. Rev.* **D 1**, 1596 (1970)
H. BETHE, W. HEITLER, *Proc. Roy. Soc.* **A146**, 83 (1934)

Interaction with the photon backgrounds – II

Photohadronic interactions – $p\gamma$ interaction rate (s^{-1} per particle):

$$\Gamma_{p\gamma\to p'b}(E,z) = \frac{1}{2}\frac{m_p^2}{E^2}\int_{\frac{\epsilon_{\rm th} m_p}{2E}}^{\infty} d\epsilon \frac{n_\gamma(\epsilon,z)}{\epsilon^2}\int_{\epsilon_{\rm th}}^{2E\epsilon/m_p} d\epsilon_r \epsilon_r \sigma^{\rm tot}_{p\gamma\to p'b}(\epsilon_r)$$

1. For given values of E and z, NeuCosmA calculates the cooling rate $t_{p\gamma}^{-1} \equiv -(1/E)\, b_{p\gamma}$ (s^{-1}) as

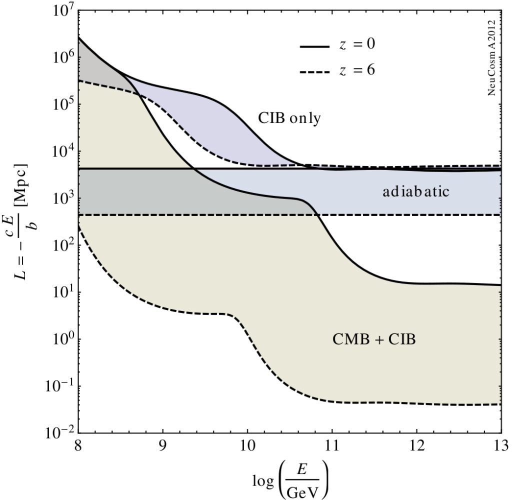
   $$t_{p\gamma}^{-1}(E,z) = \sum_i^{\text{all channels}} \Gamma^i_{p\to p}(E,z)\, K^i\ ,$$

   with $K^i E$ the loss of energy per interaction
2. From this, we calculate back $b_{p\gamma}$ (GeV s^{-1}) ...
3. ...and the corresponding energy-loss term in the transport equation, $\partial_E(b_{p\gamma} Y_p)$.

S. Hümmer, M. Rüger, F. Spanier, W. Winter, *Astrophys. J.* **721**, 630 (2010) [1002.1310]

Interaction lengths

Note that $L_{\mathrm{CIB}} \gg L_{\mathrm{CMB}}$:



Matches, *e.g.*, H. Takami, K. Murase, S. Nagataki, K. Sato, *Astropart. Phys.* **31**, 201 (2009) [0704.0979]

UHE ν's in the GRB internal shock model

Secondary injection of neutrons, neutrinos (GeV^{-1} cm^{-3} s^{-1})

$$Q'\left(E'\right) = \int_{E'}^{\infty} \frac{dE'_p}{E'_p} N'_p\left(E'_p\right) \int_0^{\infty} c d\varepsilon' N'_\gamma\left(\varepsilon'\right) R\left(E', E'_p, \varepsilon'\right)$$

Normalisation to the observed GRB photon flux F_γ

$$\int \varepsilon' N'_\gamma\left(\varepsilon'\right) d\varepsilon' = \frac{E'^{\text{sh}}_{\text{iso}}}{V'_{\text{iso}}} \propto F_\gamma \; , \quad \int E'_p N'_p\left(E'_p\right) dE'_p = \frac{1}{f_e} \frac{E'^{\text{sh}}_{\text{iso}}}{V'_{\text{iso}}} \propto \frac{F_\gamma}{f_e}$$

Fluence per shell, at Earth (GeV^{-1} cm^{-2})

$$\mathcal{F}^{\text{sh}} = t_v V'_{\text{iso}} \frac{\left(1+z\right)^2}{4\pi d_L^2} Q'$$

UHE ν's in the GRB internal shock model

Secondary injection of neutrons, neutrinos ($\mathrm{GeV}^{-1}\ \mathrm{cm}^{-3}\ \mathrm{s}^{-1}$)

$$Q'\left(E'\right) = \int_{E'}^{\infty} \frac{dE'_p}{E'_p} N'_p\left(E'_p\right) \int_0^{\infty} c d\varepsilon' N'_\gamma\left(\varepsilon'\right) R\left(E', E'_p, \varepsilon'\right)$$

- Photon density, shock rest frame ($\mathrm{GeV}^{-1}\ \mathrm{cm}^{-3}$):

$$N'_\gamma\left(\varepsilon'\right) \propto \begin{cases} \left(\varepsilon'\right)^{-\alpha_\gamma}, & \varepsilon'_{\gamma,\mathrm{min}} = 0.2\ \mathrm{eV} \leq \varepsilon' \leq \varepsilon'_{\gamma,\mathrm{break}} \\ \left(\varepsilon'\right)^{-\beta_\gamma}, & \varepsilon'_{\gamma,\mathrm{break}} \leq \varepsilon' \leq \varepsilon'_{\gamma,\mathrm{max}} = 300 \times \varepsilon'_{\gamma,\mathrm{min}} \end{cases}$$

$$\varepsilon'_{\gamma,\mathrm{break}} = \mathcal{O}\left(\mathrm{keV}\right), \alpha_\gamma \approx 1, \beta_\gamma \approx 2$$

- Proton density:

$$N'_p\left(E'_p\right) \propto \left(E'_p\right)^{-\alpha_p} \times \exp\left[-\left(E'_p/E'_{p,\mathrm{max}}\right)^2\right] \quad \left(\alpha_p \approx 2\right)$$

Maximum proton energy limited by energy losses:

$$t'_{\mathrm{acc}}\left(E'_{p,\mathrm{max}}\right) = \min\left[t'_{\mathrm{dyn}}, t'_{\mathrm{syn}}\left(E'_{p,\mathrm{max}}\right), t'_{p\gamma}\left(E'_{p,\mathrm{max}}\right)\right]$$

UHE ν's in the GRB internal shock model

Secondary injection of neutrons, neutrinos (GeV^{-1} cm^{-3} s^{-1})

$$Q'(E') = \int_{E'}^{\infty} \frac{dE'_p}{E'_p} N'_p(E'_p) \int_0^{\infty} c d\varepsilon' N'_\gamma(\varepsilon') R(E', E'_p, \varepsilon')$$

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UHE ν's in the GRB internal shock model

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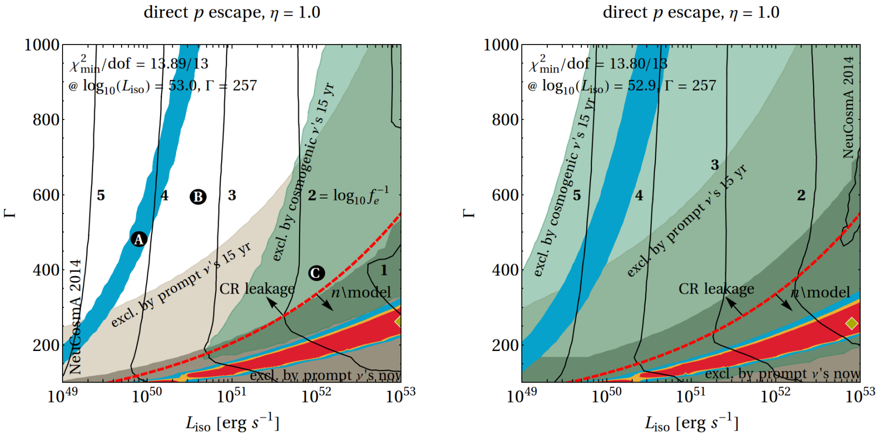
$$\int \varepsilon' N'_\gamma\left(\varepsilon'\right) d\varepsilon' = \frac{E'^{\rm sh}_{\rm iso}}{V'_{\rm iso}} \propto F_\gamma \; , \quad \int E'_p N'_p\left(E'_p\right) dE'_p = \frac{1}{f_e} \frac{E'^{\rm sh}_{\rm iso}}{V'_{\rm iso}} \propto \frac{F_\gamma}{f_e}$$

Fluence per shell, at Earth (GeV^{-1} cm^{-2})

$$\mathcal{F}^{\rm sh} = t_v V'_{\rm iso} \frac{\left(1+z\right)^2}{4\pi d_L^2} Q'$$

Star-formation-rate vs. GRB redshift evolution

The exclusion from cosmogenic ν's grows if the number of GRBs evolves more strongly with redshift:



$n_{\mathrm{GRB}}(z) \propto \rho_{\mathrm{SFR}}(z)$ (star formation rate)

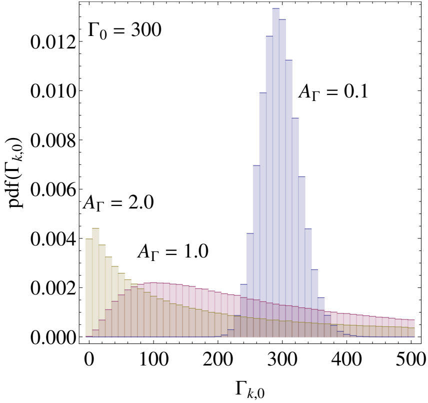
$n_{\mathrm{GRB}}(z) \propto \rho_{\mathrm{SFR}}(z) \times (1+z)^{1.2}$

Initial distribution of shell speeds

Distribution of initial shell speeds (Lorentz factors):

$$\ln\left(\frac{\Gamma_{k,0}-1}{\Gamma_0-1}\right) = A_\Gamma x$$

x follows a Gaussian distribution, $P(x)\,dx = dx\; e^{-x^2/2}/\sqrt{2\pi}$



$A_\Gamma < 1$

speeds too similar, collisions only at large radii

$A_\Gamma \gg 1$

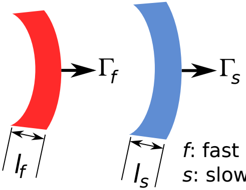
spread too large, too many collisions at low radii

$A_\Gamma \approx 1$

just right, burst has high efficiency of conversion of kinetic to radiated energy

Anatomy of an internal collision

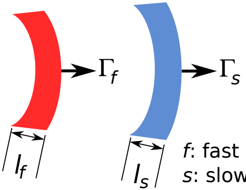
1. Propagation



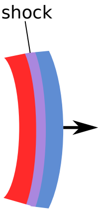
f: fast shell
s: slow shell

Anatomy of an internal collision

1 Propagation

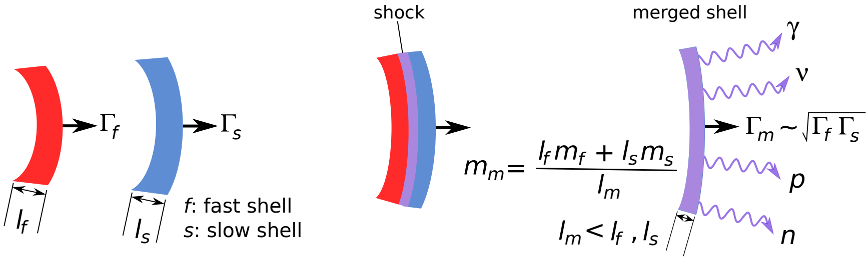


2 Collision



Anatomy of an internal collision

1 Propagation — 2 Collision — 3 Radiation



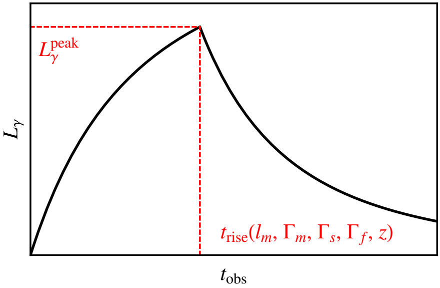
Part of the initial kinetic energy radiated as γ's, ν's, p's, and n's:

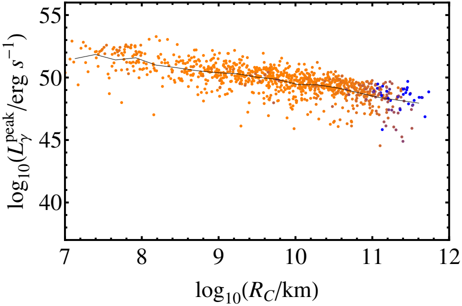
$$E_{\text{coll}}^{\text{iso}} = \left(E_{\text{kin},f}^{\text{iso}} - E_{\text{kin},m}^{\text{iso}} \right) + \left(E_{\text{kin},m}^{\text{iso}} - E_{\text{kin},s}^{\text{iso}} \right)$$

1/12	1/12	5/6
$\underbrace{\epsilon_e E_{\text{coll}}^{\text{iso}}}$	$\underbrace{\epsilon_B E_{\text{coll}}^{\text{iso}}}$	$\underbrace{\epsilon_p E_{\text{coll}}^{\text{iso}}}$
energy in photons	energy in magnetic fields	energy in baryons

Gamma-ray and neutrino pulses

A fast-rise-exponential-decay (FRED) gamma-ray pulse is emitted in every collision:

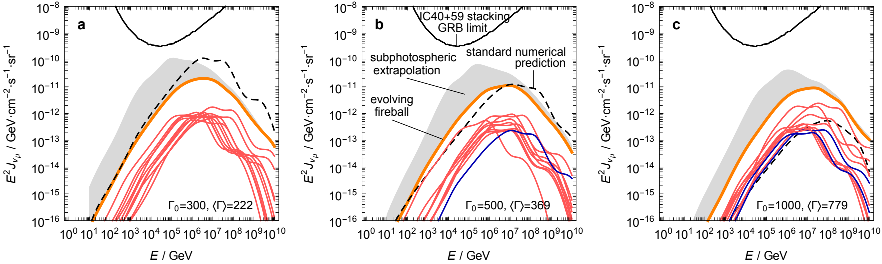




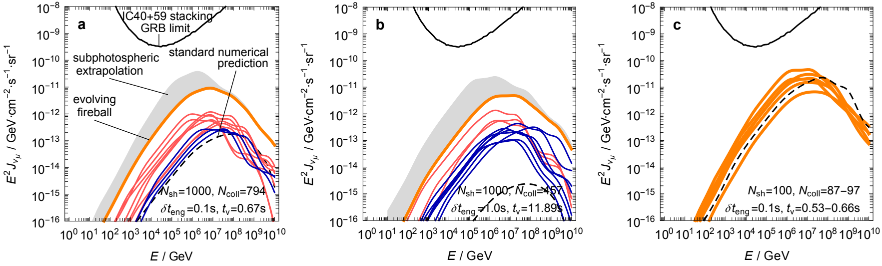
$L_\gamma^{\rm peak} \sim R_C^{-2}$

The prediction *is* robust

Simulations show only weak dependence of the **flux** on the boost Γ . . .

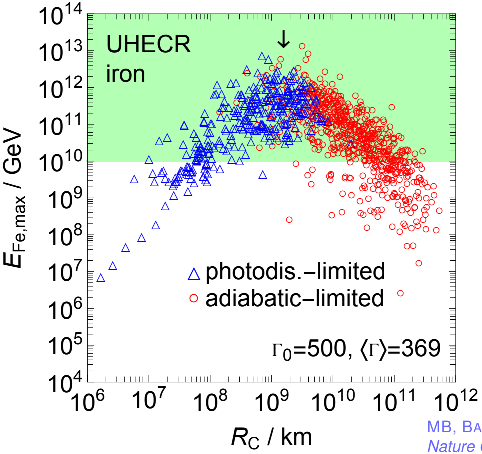


. . . and on the GRB engine variability time $\delta t_{\rm eng}$



Accelerating iron

- Photodisintegration destroys nuclei close to the center ($\sim 10^8$ km)
  e.g., ANCHORDOQUI *et al.*, *Astropart. Phys.* **29**, 1 (2008)
- However, they can survive at large radii:



MB, BAERWALD, MURASE, WINTER
Nature Commun. **6**, 6783 (2015)